On the electron-induced isotope fractionation in low temperature $^{32}\text{O}_2/^{36}\text{O}_2$ ices—ozone as a case study

B. Sivaraman, A. M. Mebel, N. J. Mason, D. Babikov and R. I. Kaiser

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The formation of six ozone isotopomers and isotopologues, $^{16}\text{O}^{16}\text{O}$, $^{16}\text{O}^{18}\text{O}$, $^{18}\text{O}^{16}\text{O}$, $^{16}\text{O}^{16}\text{O}^{18}$, $^{18}\text{O}^{18}\text{O}^{16}$, and $^{18}\text{O}^{18}\text{O}^{18}$, has been studied in electron-irradiated solid oxygen $^{16}\text{O}_2$ and $^{18}\text{O}_2$ (1 : 1) ices at 11 K. Significant isotope effects were found to exist which involved enrichment of $^{18}\text{O}$-bearing ozone molecules. The heavy $^{18}\text{O}^{18}\text{O}^{18}$ species is formed with a factor of about six higher than the corresponding $^{18}\text{O}^{18}\text{O}^{16}$ isotope. Likewise, the heavy $^{18}\text{O}^{18}\text{O}^{18}$ species is formed with abundances of a factor of three higher than the lighter $^{18}\text{O}^{18}\text{O}^{16}$ counterpart. No isotope effect was observed in the production of $^{18}\text{O}^{18}\text{O}^{16}$ versus $^{18}\text{O}^{16}\text{O}^{16}$. Such studies on the formation of distinct ozone isotopomers and isotopologues involving non-thermal, non-equilibrium chemistry by irradiation of oxygen ices with high energy electrons, as present in the magnetosphere of the giant planets Jupiter and Saturn, may suggest that similar mechanisms may contribute to the $^{18}\text{O}$ enrichment on the icy satellites of Jupiter and Saturn such as Ganymede, Rhea, and Dione. In such a Solar System environment, energetic particles from the magnetospheres of the giant planets may induce non-equilibrium reactions of suprathermal and/or electronically excited atoms under conditions, which are quite distinct from isotopic enrichments found in classical, thermal gas phase reactions.

1. Introduction

The origin of the oxygen isotopic fractionation ($^{18}\text{O}$ vs. $^{16}\text{O}$) in molecules in our Solar System such as ozone ($\text{O}_3$), water ($\text{H}_2\text{O}$), carbon monoxide (CO), and carbon dioxide (CO$_2$) has been a topic of experimental and theoretical studies since it was discovered nearly 30 years ago. First measured in chondritic meteorites, such isotopic enrichments have been found in Martian rocks and are now routinely utilized to identify Martian meteorite material that are deposited on Earth. In 1965, Taylor suggested a classification of rocks in the solar system according to their $^{18}\text{O}/^{16}\text{O}$ ratio. Recently Ozima et al. classified 400 bulk meteorite samples using their oxygen isotopic data and commented upon the similarity between solar and planetary oxygen. NASA’s ‘Genesis’ and ‘Stardust’ missions measured oxygen isotopes in the solar wind and from comet Wild 2 by collecting (and returning) ‘dust’ samples to the Earth. Such studies have provided vital information on the chemical origins of our Solar System and may reveal how the different planets were formed. Oxygen isotopes may also be used to understand the formation of the Earth–Moon system. Also, the study of oxygen isotopic variations is not confined to the Solar System; current space-based telescopes are capable of developing maps of isotopic variations in star forming regions; for example the Odin satellite has recently revealed the presence of molecular oxygen in the interstellar medium. This can also be extended to determine $^{18}\text{O}/^{16}\text{O}$ ratios in abundant, oxygen-bearing molecules such as carbon monoxide and—if detected—potentially ozone.

The $^{18}\text{O}$ enrichment in ozone has received particular attention. In the Earth’s atmosphere, ozone is formed by the Chapman mechanism (eqn (1) and (2)), where M is a third body (usually nitrogen in the terrestrial atmosphere) necessary to stabilize the ozone molecule, which would otherwise decompose back to the reactants due to the internal energy it has upon formation. This process has been studied extensively. Each of the three stable oxygen isotopes ($^{18}\text{O}$,$^{17}\text{O}$,$^{16}\text{O}$) may take part in the Chapman reactions producing a wide range of isotopomers and isotopologues all of which have been measured in the terrestrial stratosphere. Within the terrestrial atmosphere the ratio of oxygen isotopes have been measured to be $^{18}\text{O}/^{17}\text{O} = 5.5$, $^{16}\text{O}/^{18}\text{O} = 500$, $^{16}\text{O}/^{17}\text{O} = 2750$. Kinetic isotope effects and symmetry induced isotopic fractionation then suggest different rate constants for eqn (1) and (2) depending on the isotopes involved leading to distinct formation probabilities for each of the isotopomers and isotopologues. Initially, these molecules were thought to form via a cyclic ozone intermediate, whose ring opened to the acyclic form. However, this interpretation has been challenged, and fast isotopic exchange reactions involving $^{16}\text{O}$ enriched molecular oxygen ($^{16}\text{O}^{16}\text{O}$) have to be considered as well.

$$\text{O}_2(X^3\Sigma_g^-) + h\nu \rightarrow \text{O}(^3\Pi) + \text{O}(^3\Pi) \quad (1)$$

$$\text{O}(^3\Pi) + \text{O}_2(X^3\Sigma_g^-) + M \rightarrow \text{O}_3(1\Sigma_g^+) + M \quad (2)$$

Isotopic enrichment processes may also occur within oxygen-rich ices in the Solar System. However, the underlying mechanisms and reaction conditions can differ dramatically. In classical experiments, isotope enrichment processes were studied in the gas phase under thermal conditions where the thermal reactants have a well-defined Maxwell–Boltzmann distribution. However, high energy electrons induce a nonequilibrium chemistry in which multiple effects occur such as...
the generation of secondary electrons, the formation of electronically excited atoms and the generation of ground state oxygen atoms with excess kinetic energy. It must also be noted that these solid/ice experiments are not conducted under ‘single collision’ conditions as encountered in crossed beam experiments. Therefore, it is often difficult to untangle to what extend the products originate from, for instance, reactions of thermalized atoms or suprathermal atoms, where the latter can overcome reaction barriers easily.

Solid ozone has been detected on the frozen surfaces of Jupiter’s moon Ganymede and on Saturn’s satellites Rhea and Dione.\(^{24,25}\) The formation of ozone on these surfaces has then been attributed to bombardment by ions and energetic electrons from planetary magnetospheres.\(^{26}\) However, to date, it has not been possible to determine the possible role of any isotopic fractionation of ozone in such low temperature ices. Therefore, we have investigated \(^{16}\)O and \(^{18}\)O isotopic enrichment processes in ozone formation in low temperature oxygen ices \((^{16}\text{O}_2 \text{ and } ^{18}\text{O}_2)\) irradiated by energetic electrons. These experiments provide a simulation of the interaction of magnetospheric electrons with oxygen-bearing ices (oxygen) and may elucidate the underlying mechanisms forming ozone isotopomers and isotopologues in low temperature, oxygen-rich Solar System ices.

2. Experimental

The experiments were performed within an ultra high vacuum (UHV) chamber operating at base pressures of about \(5 \times 10^{-11}\) Torr.\(^{27}\) Briefly, the ice samples were prepared on a polished silver target attached to a two stage closed loop helium cryostat capable of cooling the target down to 11 K. A mixture of \(^{16}\text{O}_2\) (Gaspro; 99.999\%) and \(^{18}\text{O}_2\) (98\% \(^{18}\)O) was prepared in a gas mixing chamber with a ratio of 1.0: (1.0 ± 0.01) and deposited onto the cooled target for 300 s at a chamber pressure of \(10^{-7}\) Torr. Subliming the freshly deposited mixtures and correcting for the isotope-dependent electron impact ionization cross section, a mass spectrometric quantification of the subliming molecules at \(m/z = 32\) \(^{16}\text{O}_2^+\) and \(m/z = 36\) \(^{18}\text{O}_2^+\) suggests that there is no isotopic fractionation when condensing the gases onto the target at 11 K; therefore, within the error limits of our mass spectrometer, the \(^{16}\text{O}_2\) to \(^{18}\text{O}_2\) ratio in the ices is 1.00 ± 0.01. After deposition, the ices were irradiated with 5 keV electrons for one hour. The electron beam, generated using a commercially available electron gun, was operated at 100 nA current and scanned over 1.8 ± 0.2 cm\(^2\) of the target; accounting for the extraction efficiency of 78\%, this exposes the target to \(1.75 \times 10^{15}\) electrons during the irradiation period. After irradiation the sample was left isothermally at 11 K to monitor possible reactions of trapped oxygen atoms at 11 K. Thereafter, the ices were warmed at a rate of 0.5 K min\(^{-1}\) to 293 K. Radiation-induced chemical changes in the ice films were monitored on line and \textit{in situ} using a Fourier transform infrared (FTIR) spectrometer operating in reflection–absorption–reflection mode with a reflection angle of 75° from the normal relative to the mirror surfaces. FTIR spectra were collected over the range of 6000–500 cm\(^{-1}\) with an integration time of 2.5 min and at 2 cm\(^{-1}\) resolution.

3. Results

3.1. Infrared spectra

In the gas phase, the \(C_2\) symmetric ozone molecule has three fundamental vibrational modes: the \(v_1\) bending mode at 701 cm\(^{-1}\) (lowest intensity), the \(v_2\) symmetric stretching vibration at 1103 cm\(^{-1}\) of medium intensity, and the most intense \(v_3\) asymmetric stretching mode at 1042 cm\(^{-1}\).\(^{27}\) In the irradiated ices, ozone was monitored \textit{via} its most intense \(v_3\) fundamental around 1000 cm\(^{-1}\). The broad spectral band lying between 960 cm\(^{-1}\) to 1050 cm\(^{-1}\) was deconvoluted into six distinct peaks corresponding to the six possible ozone isotopomers and isotopologues.\(^{28–30}\) Bands at 978 cm\(^{-1}\), 988 cm\(^{-1}\), 1002 cm\(^{-1}\), 1012 cm\(^{-1}\), 1022 cm\(^{-1}\), and 1035 cm\(^{-1}\) were assigned to \(^{16}\text{O}^{18}\text{O}^{16}\text{O}, ^{18}\text{O}^{16}\text{O}^{16}\text{O}, ^{16}\text{O}^{18}\text{O}^{16}\text{O}, ^{16}\text{O}^{16}\text{O}^{18}\text{O}, ^{16}\text{O}^{16}\text{O}^{16}\text{O}\) and \(^{18}\text{O}^{16}\text{O}^{16}\text{O}\), respectively (Fig. 1a). These data are in excellent agreement with literature values.\(^{28–30}\)

A weak \(v_2\) band was also observed in the irradiated ices (Fig. 1b). However, only the \(^{16}\text{O}^{18}\text{O}^{16}\text{O}\) and \(^{18}\text{O}^{16}\text{O}^{16}\text{O}\) isotopologues could be observed at 662 cm\(^{-1}\) and 677 cm\(^{-1}\), respectively.\(^{28,30,31}\) Two bands, at 688 cm\(^{-1}\) and 673 cm\(^{-1}\), corresponding to the \(^{16}\text{O}^{18}\text{O}^{16}\text{O}\) and \(^{18}\text{O}^{16}\text{O}^{16}\text{O}\) species, only appeared when the ice was warmed to 45 K. We note that a previous study in our groups suggested that the irradiated oxygen ices trap oxygen atoms at 11 K\(^{27}\) such that, as the temperature increases, these oxygen atoms become mobile and can react with an oxygen atom or molecule to form molecular oxygen and ozone respectively. During the heating process additional bands of \(^{16}\text{O}^{18}\text{O}^{16}\text{O}\) and \(^{16}\text{O}^{16}\text{O}^{16}\text{O}\) appeared at 699 cm\(^{-1}\) and 707 cm\(^{-1}\), respectively. Finally, the weak \(v_1\) fundamental was observed at 1061 cm\(^{-1}\) and 1090 cm\(^{-1}\) (Fig. 1c); these absorptions correspond to the \(^{16}\text{O}^{16}\text{O}^{16}\text{O}\) and \(^{16}\text{O}^{16}\text{O}^{16}\text{O}\) isotopologues as evident from literature data.\(^{30}\)

3.2. Mass spectrometry

To monitor the subliming molecules, \(O_2\) and \(O_3\) containing both \(^{18}\text{O}\) and \(^{16}\text{O}\) were followed in real time by recording the ion currents of the mass-to-charge ratios \((m/z)\) of 32 \((^{16}\text{O}_2^+)\), 36 \((^{18}\text{O}_2^+)\), 48 \((^{16}\text{O}_3^+)\), 50 \((^{16}\text{O}_2^{18}\text{O}^+)\), 52 \((^{18}\text{O}_2^{16}\text{O}^+)\), and 54 \((^{18}\text{O}_3^+)\). In the actual irradiation phase neither atomic oxygen, molecular oxygen, nor ozone were observed in the gas phase since, the detection limit of our mass spectrometer has been cited by the manufacturer to have partial pressures of about \(1 \times 10^{-15}\) torr, we conclude that no desorption occurs during irradiation. However dramatic changes were observed during the heating of the irradiated sample (Fig. 2). \(^{32}\)O\(_2\) and \(^{36}\)O\(_2\) were found to desorb when the temperature reaches 28 K and their yield is maximized at 35 K. All the ozone isotopomers and isotopologues were found to start to sublime from the ice at about 59 K. Simultaneous monitoring of the FTIR spectra and the ion currents in the gas phase reveals all ozone molecules had desorbed at 63 ± 1 K.


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4. Discussion

4.1. Ozone formation during irradiation phase

Ozone was observed to form within the ices immediately after the electron irradiation commenced. Previous experiments on pure molecular oxygen ices (\(16\text{O}_2\)), which were irradiated under identical conditions as the isotopically mixed systems, suggested that upon electron exposure, an oxygen molecule undergoes unimolecular decomposition to form two oxygen atoms (\(16\text{O}_2\)) (reaction (3); term symbols are omitted for clarity). The released oxygen atom was found to react with a neighboring oxygen molecule to form an ozone molecule (\(16\text{O}_3\)) via eqn (4). Quantitatively, the temporal evolution of the ozone (\(16\text{O}_3\)) column density could be fit with (pseudo) first order kinetics.\(^{27}\) Under these irradiation conditions, the back-dissociation of ozone to molecular oxygen and an oxygen atom was not found to play any significant role.\(^{27}\)

\[
\begin{align*}
16\text{O}_2 & \rightarrow 16\text{O} + 16\text{O} \\
16\text{O}_2 + 16\text{O} & \rightarrow 16\text{O}_3
\end{align*}
\]
18O2 → 18O + 16O
18O2 + 18O → 18O3
16O2 + 18O → 16O218O
18O2 + 16O → 16O218O

Analogous to these previous experiments, in a mixed isotope oxygen ice of 16O2 and 18O2, the production of 16O and 18O atoms via eqn (3) and (5) is expected. Here, the free 16O and 18O atoms can each react with 18O3 and 16O2 to form 18O2 (16O2/18O2; eqn (4)), 18O3 (18O2/18O2; eqn (6)), 16O218O (16O2/18O2; eqn (7)), and 18O216O (16O2/18O2; eqn (8)). It should be noted that 16O18O and 18O16O can be formed either in C2v (symmetric) or C4 (asymmetric) structures, i.e. 16O18O16O and 18O16O18O as well as 16O16O18O and 18O18O16O. Here, the reactive oxygen atom formally ‘inserts’ into the oxygen–oxygen bond giving the C2v structures or can add to a terminal oxygen atom to form the C4 molecules. The temporal evolution of all of these six ozone isotopomers and isotopologues are shown in Fig. 3 and are derived by plotting the column density of the deconvoluted ν3 bands as depicted in Fig. 1a versus the time; the column densities are derived utilizing the corresponding integrated absorption coefficients as outlined in reference.27

It should be noted that in order to obtain the relative abundances of these isotopomers and isotopologues it is important to correct each band area for the integrated absorption coefficient of the ν3 band for that species. Unfortunately absorption coefficients for the different ozone isotopomers and isotopologues have not been measured so they were calculated utilizing the hybrid density functional B3LYP method23,33 with Dunning’s aug-cc-pVTZ basis set34 using the GAUSSIAN 98 quantum chemical program package.35 It is important to note that computed, absolute absorption coefficients of individual vibrational normal modes may have uncertainties as large as 20%.36,37 Nevertheless if these coefficients are calculated using the same level of theory, the relative values of the absorption coefficients of different isotopically substituted ozone molecules will have much smaller uncertainties of only a few % and, since we are investigating the relative isotopic enrichment, this method is justified. Here, the integrated absorption coefficients for the ν3 mode of the isotopomers and isotopologues have been computed to be 3.46 × 10−17 cm molecule−1 (18–18–18), 3.32 × 10−17 cm molecule−1 (18–18–16), 3.90 × 10−17 cm molecule−1 (16–16–16), 3.54 × 10−17 cm molecule−1 (16–16–18), 3.63 × 10−17 cm molecule−1 (16–18–18), and 3.73 × 10−17 cm molecule−1 (18–16–18).

A simple inspection of the column densities of the distinct isotopomers and isotopologues at the end of the irradiation exposure depicts surprising results. First, the 18–18–18 species formed via eqn (6) has column densities a factor of six higher than the corresponding 16–16–16 molecule (eqn (4)). On the other hand, the production of the C2v symmetric 16–16–16 and 18–18–18 species are—within the error limits—identical. As a matter of fact these production rates are similar to the 16–16–18 isotopologues. Finally, the C4 species 18–18–18 are higher than the 16–16–18 molecule, but close to the column density of the 18–18–18 species. These data alone suggest that even in low temperature Solar System analog ices, energetic electrons can induce significant isotope effects leading to a preferential formation of the heavier 18–18–18 versus 16–16–16 and a significantly enhanced formation of the heavy 18–18–16 versus 16–16–18. No isotopic effect was found in case of 16–18–16 versus 18–16–18.

In an attempt to understand the underlying reaction mechanisms, we fitted the column densities as compiled in Fig. 3 using two different kinetic models for ozone formation via (pseudo) first order kinetics (eqn (9); mechanism M1) or through a two-step mechanism (eqn (10); mechanism M2). M1 could characterize an end-on addition of the oxygen atom to the oxygen molecule; M2 might represent a consecutive reaction mechanism in which the oxygen atom first adds to the ring forming a cyclic ozone intermediate and then ring opens; alternatively, this two-step process could resemble the formation of free oxygen atoms forming first a 18O16O molecular oxygen species (step I) which then reacts with 16O or 18O to ozone (step II).

Again, as stated in the introduction, our experiments are designed to investigate if isotopic fractionation processes can occur in low temperature Solar System analog ices via non-equilibrium chemistry; therefore, the rate constants presented here refer to non-equilibrium rate constants, and cannot be compared with classical rate constants derived for isotopic enrichment studies in the gas phase. Eqn (9) and (10) mathematically describe a pseudo first order reaction A → B8 and a consecutive two step process via the reaction sequence A → B → C (eqn 10).27 Present to an unbiased and objective analysis, we fit the graphs of all isotopomers and isotopologues utilizing both models via M1 and M2. The rate constants of both fits are compiled in Tables 1 and 2, respectively.

\[ \frac{[\text{O}_3]}{[\text{A}]} = a_1(1 - e^{-kt}) \] (9)

\[ \frac{[\text{O}_3]}{[\text{A}]} = \frac{(a_2/(k_2 - k_1))(k_3(1 - e^{-kt_1}) - k_1(1 - e^{-kt_2}))}{(1 - e^{-kt})} \] (10)

The temporal evolutions can be fitted using eqn (9) (solid lines; Fig. 3; Table 1) utilizing rate constants between 2.7 and 4.2 × 10−2 s−1. Eqn (10) also presented a reasonable fit of the measured column density profiles. Therefore, a kinetic analysis of the data suggests that the ozone isotopes can be formed either by a pseudo first order or second order reaction mechanism—or both. Both 16–16–16 and 18–18–18 could be formed via reaction of 18O with 16O2 and 16O with 16O2, respectively, via a cyclic ozone intermediate. Alternatively, an 18O18O molecule formed by recombination of 16O and 18O could react with 18O or 16O.

There has been a considerable debate as to the possibility in forming and stabilizing a ring structure of O3. It is estimated that the cyclic form of ozone lies ~ 1 eV above the dissociation limit.38 Fleming et al.39 have further argued that the lifetime of cyclic form at low temperatures may be greater than 30 seconds. We therefore searched for the evidence of this cyclic form which is predicted to have spectroscopic signatures to be at 1099.6 (ν1), 782.9 (ν2, ν3) cm−1 (these values are for 48O2), but we found no evidence for this form of ozone suggesting either that it is not formed or its lifetime is much shorter than Fleming et al. predict. The 16O18O16O and 18O18O18O isotopologues may be formed by either the
end-on addition of $^{16}$O to $^{16}$O$^{16}$O$^{18}$O to $^{16}$O$^{18}$O. Due to large concentrations of their parent molecules ($^{16}$O$^{16}$O and $^{18}$O$^{18}$O) this is likely to be the dominant formation mechanism. It is also possible to conclude that the end-on addition to be the dominant pathway to produce $^{16}$O$^{16}$O$^{18}$O and $^{18}$O$^{18}$O$^{16}$O isotopologues, however we note that $^{16}$O$^{16}$O$^{16}$O is produced with a quarter of the efficiency of the $^{18}$O$^{18}$O$^{18}$O isotopologues.

4.2. Ozone formation during warm-up

After irradiation the sample was kept under isothermal conditions for an hour before the irradiated ice was warmed at the rate of 0.5 K min$^{-1}$. After the irradiation oxygen atoms are still trapped in the ice matrix. 27 Although O(1D) can be

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<th>Rate coefficient ($k$/s$^{-1}$)</th>
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<td>$^{18}$O$^{16}$O$^{16}$O</td>
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<table>
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<th>Product</th>
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<th>$k_2$/s$^{-1}$ (2nd step)</th>
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<td>$^{16}$O$^{16}$O$^{16}$O</td>
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<td>$3.5 \pm 0.8 \times 10^{-2}$</td>
<td>$4.7 \pm 4.3 \times 10^{-1}$</td>
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Fig. 3 Temporal evolution of ozone isotopomers and isotopologues in the ice during the irradiation at 10 K together with kinetic fits corresponding to pseudo first order reaction (given in eqn (3)) for (a) $^{18}$O$^{18}$O$^{18}$O, (b) $^{18}$O$^{18}$O$^{16}$O, (c) $^{16}$O$^{16}$O$^{18}$O, (d) $^{16}$O$^{16}$O$^{18}$O, (e) $^{16}$O$^{16}$O$^{16}$O, and a two step consecutive reaction for (f) $^{18}$O$^{16}$O$^{18}$O, given in eqn (4) respectively. Dotted lines show the fits using eqn (4) for a–d and using eqn (3) e–f.

Table 1 Rate coefficients obtained by fitting the temporal profiles of ozone isotopomers and isotopologues via eqn (9)

Table 2 Rate coefficients obtained by fitting the temporal profiles of ozone isotopomers and isotopologues via eqn (10)
produced in the matrix during irradiation, after 3600 s in the isothermal condition all these atoms will have relaxed into their ground state leaving the irradiated ice rich in atomic oxygen in its ground state ($^3P$). As the temperature is increased, these oxygen atoms start to diffuse forming (i) molecular oxygen due to recombination of oxygen atoms and (ii) ozone due to the addition of oxygen atoms to the nascent molecular oxygen. This explains our observation of an increased ozone production during temperature programmed desorption for all six ozone isotopomers and isotopologues (Fig. 4). After the oxygen molecules have desorbed from the surface at around 35 K, no further increase in ozone production is observed and only ozone is left on the surface. However at about 47 K, novel bands appear in the infrared spectra indicating a change of phase from the $\alpha$ (amorphous) to $\beta$ (crystalline) ozone (Fig. 1(b and c)). All ozone isotopomers and isotopologues have sublimed from the surface at around 63 K.

5. Conclusion

We have studied the formation of six ozone isotopomers and isotopologues in electron-irradiated solid oxygen $^{16}$O$_2$ and $^{18}$O$_2$ (1 : 1) at 11 K. These are: $^{16}$O$^{16}$O$^{16}$O, $^{18}$O$^{18}$O$^{16}$O, $^{16}$O$^{16}$O$^{18}$O, $^{18}$O$^{18}$O$^{16}$O, $^{16}$O$^{18}$O$^{16}$O, and $^{18}$O$^{16}$O$^{18}$O. We have shown that under such conditions there are profound isotope effects which are induced by non-equilibrium chemistry. The heavy $^{18}$O$^{18}$O$^{16}$O species is formed with factor of about six more than the corresponding $^{16}$O$^{16}$O$^{16}$O molecule. Likewise, the heavy $^{18}$O$^{18}$O$^{16}$O species is formed with abundances of about a factor of three higher than the lighter $^{16}$O$^{16}$O$^{16}$O counterpart. No isotope effect was observed in the production of $^{16}$O$^{16}$O$^{16}$O versus $^{18}$O$^{18}$O$^{16}$O. The production of distinct ozone isotopomers and isotopologues by irradiation of oxygen ices with high energy electrons is relevant in the magnetosphere of the giant planets Jupiter and Saturn, where similar mechanisms may lead to the $^{16}$O enrichment on the icy satellites of Jupiter and Saturn. Whether such $^{16}$O enrichment occurs in more complex ices (e.g. carbon dioxide, water, and oxygen rich ices more typical of planetary surfaces) remains to be seen in future laboratory experiments. Likewise, although we have presented solid evidence for the fact that energetic electrons can lead to a $^{18}$O enrichment in ozone in irradiated oxygen ices, the underlying reaction mechanisms and possible involvement of ‘cyclic ozone’ still remains to be resolved. Does a symmetry-induced isotopic fractionation present a valid answer even in low temperature Solar System ices? Or can a symmetry-based dynamical effect, as suggested for calcium-aluminium rich inclusions in meteorites, be the source of these enrichment processes? Likewise, these processes could also involve long-lived resonances in the delta zero-point energies. For completeness, we should also state that we have no evidence for radiation induced tetraoxygen molecules (O$_4$). These issues remain to be investigated theoretically in the future.

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