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How to cite:

Clements, D. L.; Rigby, E.; Maddox, S.; Dunne, L.; Mortier, A.; Pearson, C.; Amblard, A.; Auld, R.; Baes, M.; Bonfield, D.; Burgarella, D.; Buttiglione, S.; Cava, A.; Cooray, A.; Dariush, A.; de Zotti, G.; Dye, S.; Eales, S.; Frayer, D.; Fritz, J.; Gardner, Jonathan P.; Gonzalez-Nuevo, J.; Herranz, D.; Ibar, E.; Ivison, R.; Jarvis, M. J.; Lagache, G.; Leeuw, L.; Lopez-Caniego, M.; Negrello, M.; Pascale, E.; Pohlen, M.; Rodighiero, G.; Samui, S.; Serjeant, S.; Sibthorpe, B.; Scott, D.; Smith, D. J. B.; Temi, P.; Thompson, M.; Valtchanov, I.; van der Werf, P. and Verma, A. (2010). *Herschel*-ATLAS: extragalactic number counts from 250 to 500 microns. *Astronomy & Astrophysics*, 518, article no. L8.

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LETTER TO THE EDITOR

Herschel-ATLAS: Extragalactic number counts from 250 to 500 microns[★]

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Received 3 March 2010 / Accepted 12 May 2010

ABSTRACT

Aims. The *Herschel*-ATLAS survey (H-ATLAS) will be the largest area survey to be undertaken by the *Herschel* Space Observatory. It will cover 550 sq. deg. of extragalactic sky at wavelengths of 100, 160, 250, 350 and 500 μm when completed, reaching flux limits (5σ) from 32 to 145 mJy. We here present galaxy number counts obtained for SPIRE observations of the first ~ 14 sq. deg. observed at 250, 350 and 500 μm .

Methods. Number counts are a fundamental tool in constraining models of galaxy evolution. We use source catalogs extracted from the H-ATLAS maps as the basis for such an analysis. Correction factors for completeness and flux boosting are derived by applying our extraction method to model catalogs and then applied to the raw observational counts.

Results. We find a steep rise in the number counts at flux levels of 100–200 mJy in all three SPIRE bands, consistent with results from BLAST. The counts are compared to a range of galaxy evolution models. None of the current models is an ideal fit to the data but all ascribe the steep rise to a population of luminous, rapidly evolving dusty galaxies at moderate to high redshift.

Key words. galaxies: evolution – galaxies: statistics – infrared: galaxies – submillimeter: galaxies

[★] *Herschel* is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.

1. Introduction

The determination of the number of sources observed as a function of brightness is fundamental in observational astronomy.

These so-called number counts have been used in the past to determine the structure of our Galaxy (Herschel 1784), rule out cosmological models (Ryle & Clark 1961), and examine the evolution of galaxies and quasars (e.g. Smail et al. 1995; Hasinger et al. 1993). Number counts have revealed that very rapid evolution must take place among infrared luminous populations to account for the large number of sources seen at faint fluxes in the submillimetre (e.g. Hughes et al. 1998; Smail et al. 1997; Eales et al. 2000) and to produce the cosmic infrared background (Puget et al. 1996; Fixsen et al. 1998; Hauser & Dwek 2001). The detailed origin and nature of this evolution remains unclear following observations in the mid-IR (e.g. Elbaz et al. 1999; Papovich et al. 2004; Le Floch et al. 2009; Babbedge et al. 2006) and far-IR (e.g. Frayer et al. 2006) using ISO and *Spitzer*, and extensive followup studies of submm sources (e.g. Chapman et al. 2005; Clements et al. 2008; Dye et al. 2008). Issues needing to be addressed include spectral energy distribution (SED) evolution and the redshift distribution of the evolving luminous populations. Understanding these processes is a key goal of the *Herschel* Space Observatory (Pilbratt et al. 2010) and of the various large extragalactic surveys that it is undertaking. We here report the first number counts obtained by the *Herschel*-ATLAS (H-ATLAS) survey (Eales et al. 2010) which set new constraints on models of galaxy evolution in the far-IR/submm.

2. Observations and source extraction

The overall H-ATLAS survey is described by Eales et al. (2010). Briefly, it is a survey that will cover ~ 550 sq. deg. of extragalactic sky at wavelengths of 100, 160, 250, 350 and 500 μm to 5σ depths of 32–145 mJy. The observations discussed here cover a ~ 14 sq. deg. region, $<3\%$ of the total area, which were obtained during the science demonstration phase (SDP) of *Herschel* operations. This field is centred on RA = 09:05:30.0 Dec = 00:30:00.0. The field lies in the centre of one of the Galaxy And Mass Assembly (GAMA) optical survey fields (Driver et al. 2009), GAMA9, so plentiful complementary data from UV to near-IR, including optical spectra, is available.

H-ATLAS uses parallel mode observations which provide data simultaneously from both the PACS (Poglitsch et al. 2010) and SPIRE (Griffin et al. 2010) instruments. The time-line data is reduced using a naive mapping technique after removing instrumental temperature variations from the time-line data (Pascale et al. in prep., see also Griffin et al. 2008). Noise maps were generated using the two cross-scan measurements to estimate the noise per detector pass, and then for each pixel the noise is scaled by $\sqrt{\text{number of detector passes}}$. PACS data and results will be discussed elsewhere.

Sources were identified in the SPIRE maps by first subtracting a local background, estimated from the peak of the histogram of pixel values in 30×30 blocks, corresponding to $2.5'$ for the 250 μm map, and $5'$ for the 350 and 500 μm maps. Larger blocks, with 60×60 and 120×120 pixels were tried, but these produced no significant differences ($> \sim 1\sigma$) in derived counts. The background at each pixel was then estimated using a bi-cubic interpolation between the coarse grid of backgrounds. The background subtracted maps were then filtered by the estimated PSF, including inverse variance weighting, where the noise for each pixel was estimated from the noise map.

The maps from all three bands are then combined with weights set by the local inverse variance, and also the prior expectation of the SED of the galaxies. We tried a flat-spectrum prior, where equal weight is given to each band and also 250 μm weighting, where only the 250 μm was included. At the depth of the filtered maps source confusion is a significant problem,

and the higher resolution of the 250 μm maps outweighed the signal-to-noise gain from adding in the other bands. So for our current data we chose to use the 250 μm only prior for all our catalogues. We plan to revisit this issue in future data-releases, but also compared the number counts produced by our chosen method and the flat spectrum prior and found no significant differences. Further information on our background subtraction and source extraction methods can be found in Maddox et al. (2010) and Rigby et al. (in prep).

All local peaks are identified in the 250 μm PSF filtered map as potential sources, and a Gaussian is fitted to each peak. This provides an estimate of the position at the sub-pixel level, and an estimate of the peak value, which gives the best flux estimate for a point source. The fluxes in other bands were estimated by using a bi-cubic interpolation to the position given by the 250 μm map. This extraction method was validated against a range of source fluxes in our simulations, and will be discussed in detail in Rigby et al. (in prep). To produce a catalogue of reliable sources, we selected only sources that are detected at the 5σ level in any of the bands. In calculating the σ for each source, we use the relevant noise map, and add the confusion noise to this in quadrature. The average 1σ instrumental noise values for the PSF filtered maps are 4, 4 and 5.7 mJy/beam respectively in the 250, 350 and 500 μm bands. We estimated the confusion noise from the difference between the variance of the maps and the expected variance due to instrumental noise, and find that the 1σ confusion noise is 5, 6 and 7 mJy/beam at 250, 350 and 500 μm . The resulting total 5σ limits are 33 36 and 45 mJy/beam. It should also be noted that overall flux calibration uncertainties are at the 15% level for SPIRE (Swinyard et al. 2010).

3. The H-ATLAS number counts

We use our extracted catalogs as the basis for the determination of the H-ATLAS number counts in the SPIRE bands. Additionally, those objects identified with optical galaxies whose size extends beyond the instrumental beam, have fluxes from an explicit aperture extraction with size equal to the instrument beam added in quadrature to the measured object size. One further source, a merger of two galaxies where the optical radius is too small to include the second component, had an aperture assigned by hand. An additional very bright, extended source, known to be galactic in origin (Thompson et al. in prep), was excluded from our analysis. The resulting catalog of source fluxes in the three bands is cut so that only sources detected at 5σ or above in a given band are used for the counts.

Completeness, reliability and the effects of both flux boosting (e.g. Coppin et al. 2006) and blending were assessed using simulations of our observations. These have the same noise properties as the processed maps and include a realistic cirrus background, based on IRAS measurements. Simulated sources were generated using the model of Negrello et al. (2007), with an additional reduction in flux density of 15% to provide better agreement with the measured counts. The source extraction method described above was applied to these images, and the resulting simulated catalogue was also cut at 5σ in each band. Flux correction factors matching input and output counts, and accounting for the effects of both boosting/blending and incompleteness, were derived from the simulations. Where greater than 5%, these corrections were applied to the real counts to produce our final numbers. The band mainly affected by this is the 500 μm channel where flux boosting and blending are most significant due to the large beams. There is also a correction to the lowest flux bin of the 250 μm band as a result of incompleteness.

Table 1. H-ATLAS galaxy counts.

Flux (mJy)	No. of Galaxies	Flux Correction	Int. Counts $N(>S)$ per sq. deg.	Diff. Counts gal. sr ⁻¹ Jy ^{1.5}
250 μm				
800	10	1	0.7 ± 0.2	2200 ± 700
450	4	1	1 ± 0.3	1200 ± 600
350	14	1	2 ± 0.4	2300 ± 600
250	16	1	3 ± 0.5	1100 ± 300
175	25	1	5 ± 1.0	1500 ± 300
125	115	1	10 ± 1.0	2900 ± 300
93	69	1	20 ± 1	4200 ± 500
83	126	1	25 ± 2	5700 ± 500
73	232	1	45 ± 2	7600 ± 500
63	418	1	70 ± 2	9600 ± 500
56	345	1	100 ± 3	12000 ± 650
52	506	1	150 ± 5	14000 ± 600
47	710	1	200 ± 5	16000 ± 600
42	1093	1	300 ± 5	18500 ± 600
38	1702	1	374 ± 5	21500 ± 500
34	898	1.1	436 ± 6	23000 ± 800
350 μm				
700	7	1	0.5 ± 0.2	800 ± 300
250	8	1	1.0 ± 0.3	600 ± 200
175	8	1	1.6 ± 0.3	500 ± 200
125	27	1	3.5 ± 0.5	700 ± 150
91	16	1	4.6 ± 0.5	900 ± 250
82	43	1	7.6 ± 1	2000 ± 300
71	92	1	14.0 ± 1	2800 ± 300
61	195	1	28 ± 1	4200 ± 300
55	158	1	39 ± 2	5000 ± 400
50	256	1	56 ± 2	6600 ± 400
46	458	1	88 ± 3	9400 ± 450
41	669	1	134 ± 3	$11\ 000 \pm 400$
38	594	1	176 ± 4	$13\ 000 \pm 550$
500 μm				
250	4	1	0.28 ± 0.15	285 ± 150
150	7	1	0.76 ± 0.25	140 ± 50
95	3	1	0.97 ± 0.3	190 ± 100
82	13	1	1.88 ± 0.4	370 ± 100
70	26	0.89	3.68 ± 0.5	580 ± 100
63	16	0.84	4.79 ± 0.6	580 ± 150
60	19	0.82	6.11 ± 0.7	760 ± 200
56	43	0.82	9.1 ± 0.8	1100 ± 200
52	76	0.83	14.4 ± 1.0	1300 ± 150
48	50	0.84	17.8 ± 1.0	1900 ± 300

Notes. Correction column gives the effective flux correction factor required to match the output to input counts from the simulations, and the fluxes quoted are after correction.

The Euclidean normalised number counts derived from this data and using the corrections shown in Table 1 are shown in Fig. 1, alongside a variety of models and data from BLAST (Patanchon et al. 2009) and, at 350 μ m, Khan et al. (2007). Table 1 also provides raw counts, normalised differential counts and integral counts.

4. Discussion

We compare the H-ATLAS counts to a range of models and to existing data from BLAST (Patanchon et al. 2009), SHARC2 (Khan et al. 2007) at 350 μ m, and the Serjeant & Harrison (2005) study of the local far-IR/submm luminosity function. The models include those of Lagache et al. (2004), Pearson et al. (2009), Negrello et al. (2007), Lacey et al. (2009) and Wilman et al. (2010) and cover a wide range of modelling techniques and assumptions. The Lagache et al. (2004) model is a backward-evolution model in which two template SEDs, one for

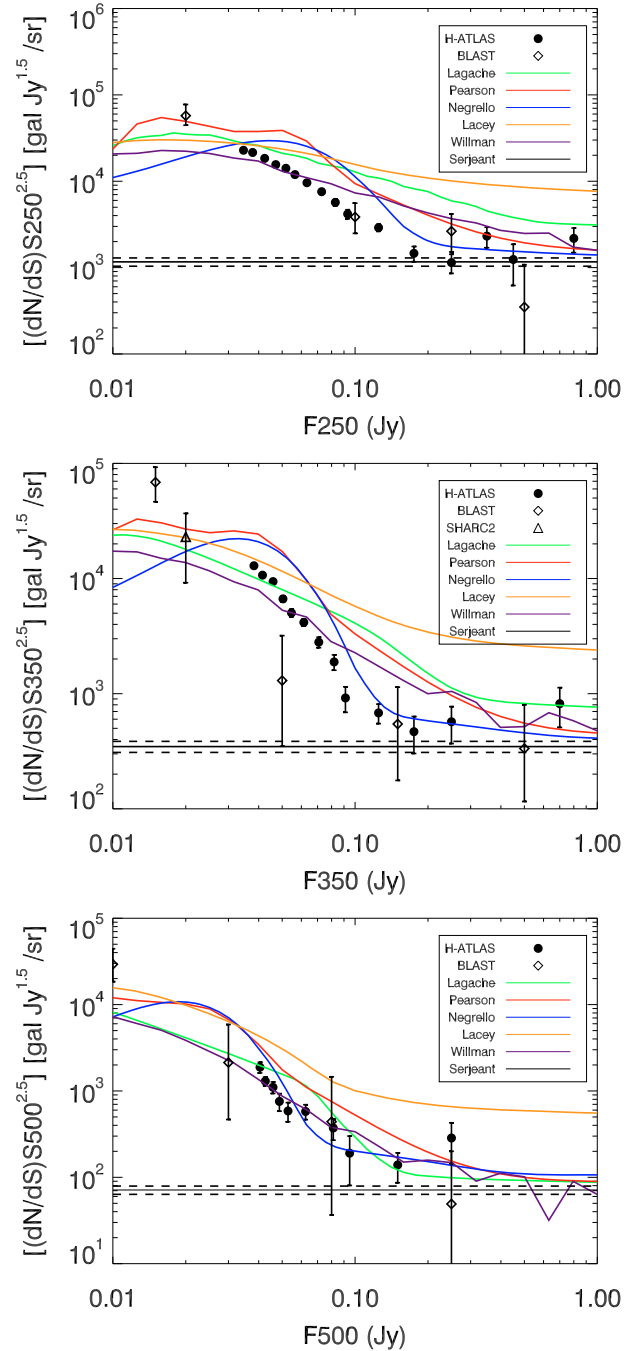


Fig. 1. H-ATLAS Euclidean normalised differential number counts compared to models and other data. Data from this work are shown as solid dots. Data from BLAST is shown as open diamonds. The solid horizontal line shows the asymptotic local differential counts from Serjeant & Harrison (2005) with dashed lines showing the $\pm 11\%$ fractional uncertainty. Models are shown as follows: Lagache et al. (2004) – green; Pearson & Khan (2009) – red; Negrello et al. (2007) – blue; Lacey et al. (2008) – orange; Wilman et al. (2010) – purple. *Top*: 250 μ m counts. *Middle*: 350 μ m counts, with the 350 μ m point of Khan et al. (2007) shown as an open triangle. *Bottom*: 500 μ m counts.

normal galaxies and another for starbursts, are evolved independently from $z = 0$. The Pearson et al. (2009) model is a similar backward-evolution model but instead uses seven luminosity dependent template SEDs. The rate of evolution is also luminosity dependent with stronger evolution in the more luminous sources (bright mode evolution in the terminology of Pearson et al. (2009)). The Negrello et al. (2007) count model includes

contributions from a range of populations. These include submm galaxies evolving according to a physical model (Granato et al. 2004), radio sources, as modelled by DeZotti et al. (2005), and late-type (starburst plus normal spiral) galaxies. Unlike the others, the counts of SMGs in this model also include a contribution from strongly-lensed galaxies, quantified following the recipes of Perrotta et al. (2002). The Lacey model is calculated from a theoretical model of galaxy formation based on CDM (Lacey et al. 2010). It uses the GALFORM semi-analytical galaxy formation model in combination with the GRASIL model for the reprocessing of starlight by dust (Granato et al. 2000). The galaxy formation model is the same as that described in Baugh et al. (2005), and compared to *Spitzer* data in Lacey et al. (2008), and incorporates a top-heavy IMF in starbursts triggered by galaxy mergers. The Wilman et al. (2010) source counts are based on the large SKADS radio continuum simulation (Wilman et al. 2008) and uses the radio-FIR relation together with a model of the evolution of this relation, along with various prescriptions to account for the AGN and starburst populations, to predict the counts in the FIR/submm. All of these models are broadly consistent with existing data on galaxy counts, redshift distributions and luminosity functions from the mid-IR to radio bands, including data from IRAS, *Spitzer* and ground-based submm surveys.

A clear feature in the observed counts is a steep increase above an extrapolation of the bright counts to fainter fluxes. This increase starts at fluxes $\sim 100\text{--}200$ mJy. Despite being hinted at in earlier observations, *Herschel* allows this rise in counts to be unambiguously detected and studied in detail for the first time. The feature is seen in all three SPIRE bands, but is most pronounced in the longest wavelength, 350 and 500 μm , bands where the counts rise roughly twice as much as at 250 μm . Such a feature is predicted in a number of models though no model is a particularly good fit to the counts in all bands. The best performing model with respect to the shape of the counts is that of Negrello et al. (2007), but this model overpredicts the effect of the bump at 250 and 350 μm . The Negrello model also provides a better fit to the brighter counts than most of the other models, especially at 350 μm , but there is clearly much work to be done to improve the match between any of these models and the observed data. The number count rise represents the strongly evolving luminous ($>10^{11} L_{\odot}$) galaxy population that is a generic feature in all galaxy evolution models that are able to fit the submm galaxies. The details of this evolution, possibly combined with other effects such as gravitational lensing (a feature of the Negrello model), control the shape and position of the rise. Its strength as a function of wavelength will be dependent on the detailed redshift and possibly evolving SEDs of the rapidly-evolving population which can also be constrained by the colour of individual sources (see e.g. Amblard et al. 2010). Its prominence at 350 and 500 μm compared to 250 μm thus argues that the sources responsible lie at $z > 1.5$, as suggested by Granato et al. (2004). If they lie at lower redshifts they would require colder SEDs. It is also interesting to note that some models (e.g. that of Lacey) appear to be strongly excluded by the observations and that the largest discrepancy between predicted and observed counts is at the bright end, where the full H-ATLAS survey will be uniquely powerful. These models are also unable to reproduce the asymptotic local counts from Serjeant & Harrison (2005) based on IRAS data. This suggests that our knowledge of brighter, more local far-IR/submm galaxies is not as good as we might wish.

These new counts from H-ATLAS set the scene for a much better understanding of the enigmatic far-IR/submm population. The SDP data presented here represents less than 3% of the

total area that will be covered by H-ATLAS. Although this is only a small fraction of H-ATLAS, these data already include more than 5 times as many ($>5\sigma$) sources as previously published at these wavelengths. As further data is added, more detailed measurements towards the bright end will allow the effects of strong lensing and local sources on the counts to be determined. H-ATLAS is also large enough to totally overcome the effects of cosmic variance, and allow it to be studied by comparing counts in sub-regions.

5. Conclusions

We have used the first ~ 14 sq. deg. of the H-ATLAS project observed by *Herschel* to determine the number counts of far-IR/submm sources in the SPIRE bands at 250, 350 and 500 μm . We find that there is a steep rise in the counts at fluxes of 100–200 mJy. These observations are compared to a range of galaxy evolution models and data from BLAST and elsewhere. The H-ATLAS counts are found to be consistent with previous observations, but no theoretical model is an ideal fit. Some of the models are quite poor matches to the observations, with counts at fluxes brighter than ~ 200 mJy setting strongly excluding some models. The best fitting model of those considered is that of Negrello et al. (2007). These observations represent just 3% of the total area to be covered by the H-ATLAS survey. Once completed this project will clearly be a very powerful tool for understanding the far-IR/submm galaxy population.

References

- Amblard, A., et al. 2010, A&A, 518, L9
- Babbedge, T. S. R., et al. 2006, MNRAS, 353, 654
- Baugh, C., et al. 2005, MNRAS, 355, 1191
- Chapman, S. C., Blain, A. W., Smail, I., & Ivison, R. J. 2005, ApJ, 622, 722
- Clements, D. L., Vaccari, M., Babbedge, T., et al. 2008, MNRAS, 387, 247
- Coppin, K., Chapin, E. L., Mortier, A. M. J., et al. 2006, MNRAS, 372, 1621
- De Zotti, G., Ricci, R., Mesa, D., et al. 2005, A&A, 431, 893
- Driver, S. P., et al. 2009, Astron. Geophys., 50, 5.12
- Dye, S., Eales, S. A., Aretxaga, I., et al. 2008, MNRAS, 386, 1107
- Eales, S. A., Lilly, S., Webb, T., et al. 2000, AJ, 120, 2244
- Eales, S. A., Dunne, L., Clements, D., et al. 2010, PASP, 122, 499
- Elbaz, D., Cesarsky, C. J., Fadda, D., et al. 1999, A&A, 351, L37
- Fixsen, D. J., Dwek, E., et al. 1998, ApJ, 508, 123
- Le Floc'h, E., Aussel, H., Ilbert, O., et al. 2009, ApJ, 703, 222
- Frayser, D. T., Fadda, D., Yan, L., et al. 2006, AJ, 131, 250
- Gratta, G. L., de Zotti, G., et al. 2004, ApJ, 600, 580
- Granato, G. L., Lacey, C. G., Silva, L., et al. 2000, ApJ, 542, 710
- Griffin, M. G., Dowell, C. D., Lim, T., et al. 2008, Proc. SPIE, 7010, 70102Q
- Griffin, M. J., et al. 2010, A&A, 518, L3
- Hasinger, G., Burg, R., Giacconi, R., et al. 1993, A&A, 275, 1
- Hauser, M. G., & Dwek, E. 2001, ARAA, 39, 249
- Herschel, W. 1784, RSPT, 74, 437
- Hughes, D. G., Serjeant, S., Dunlop, J., et al. 1998, Nature, 394, 241
- Khan, S. A., Shafer, R. A., Serjeant, S., et al. 2007, ApJ, 665, 973
- Lacey, C. G., Baugh, C. M., Frenk, C. S., et al. 2008, MNRAS, 385, 1155
- Lacey, C. G., Baugh, C. M., et al. 2010, MNRAS, 405, 2
- Lagache, G., Dole, H., Puget, J.-L., et al. 2004, ApJS, 154, 112
- Maddox, S. J., et al. 2010, A&A, 518, L11
- Negrello, M., Perrotta, F., González-Nuevo, J., et al. 2007, MNRAS, 377, 1557
- Papovich, C., Dole, H., Egami, E., et al. 2004, ApJS, 154, 70
- Pataanchon, G., Ade, P. A. R., Bock, J. J., et al. 2009, ApJ, 707, 1750
- Pearson, C. P., & Khan, S. A. 2009, MNRAS, 299, L11
- Perrotta, F., Baccigalupi, C., Bartelmann, M., de Zotti, G., & Granato, G. L. 2002, MNRAS, 329, 445
- Pilbratt, G. L., et al. 2010, A&A, 518, L1
- Poglitsch, A., et al. 2010, A&A, 518, L2
- Puget, J.-L. P., Abergel, A., Bernard, J.-P., et al. 1996, A&A, 308, L5
- Ryle, M., Clark, & R. W. 1961, MNRAS, 122, 349
- Serjeant, S., & Harrison, D. 2005, MNRAS, 356, 192
- Smail, I., Hogg, D. W., Yan, L., & Cohen, J. G. 1995, ApJ, 449, L105
- Smail, I., Ivison, R. J., & Blain, A. W. 1997, ApJ, 490, L5
- Swinyard, B. M., et al. 2010, A&A, 518, L4
- Wilman, R. J., Miller, L., Jarvis, M. J., et al. 2008, MNRAS, 388, 1335
- Wilman, R. J., Jarvis, M. J., et al. 2010, MNRAS, in press [arXiv:1002.1112]