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Physical properties of the Sh2-104 H II region as seen by Herschel


Context. Sh2-104 is a Galactic H II region with a bubble morphology, detected at optical and radio wavelengths. It is considered the first observational confirmation of the collect-and-collapse model of triggered star-formation.

Aims. We aim to analyze the dust and gas properties of the Sh2-104 region to better constrain its evolutionary stage and local future generations of stars. In addition, we investigate the relationship between the dust emissivity index β and the dust temperature, $T_d$.

Methods. Using Herschel PACS and SPIRE images at 100, 160, 250, 350 and 500 μm, we determine $T_d$ and β throughout Sh2-104, fitting the spectral energy distributions (SEDs) obtained from aperture photometry. With the SPIRE Fourier-transform spectrometer (FTS) we obtained spectra at different positions in the Sh2-104 region. We detect J-ladders of $^{12}$CO and $^{13}$CO, with which we derive the gas temperature and column density. We also detect proxies of ionizing flux as the [N II] $P_{0} - P_{1}$ and [CI] $P_{2} - P_{1}$ transitions.

Results. We find an average value of $β \approx 1.5$ throughout Sh2-104, as well as a $T_d$ difference between the photodissociation region (PDR, ~25 K) and the interior (~40 K) of the bubble. We recover the anti-correlation between $β$ and dust temperature reported numerous times in the literature. The relative isotopologue abundances of CO appear to be enhanced above the standard ISM values, but the obtained value is very preliminary and is still affected by large uncertainties.

Key words. stars: formation – ISM: bubbles – H II regions – dust, extinction – infrared: ISM – ISM: individual objects: Sh2-104

1. Introduction

Sharpless 104 (Sh2-104, Sharpless 1959) is an optically visible Galactic H II region with a bubble morphology, excited by an O6V star (Crampton et al. 1978; Lahulla 1985). It is located ~4 kpc from the Sun (Deharveng et al. 2003), with galactic coordinates 74.7620; +0.60 (J2000).

Deharveng et al. (2003) proposed Sh2-104 as a strong candidate for massive triggered star formation through the collect-and-collapse process (Elmegreen & Lada 1977). The ionized region is also visible at radio wavelengths (Fich 1993), and an ultracompact (UC) H II region, coincident with the IRAS 20160+3636 source, lies at its eastern border (Condon et al. 1998).

We present new submm images and spectra taken towards Sh2-104 with the Herschel Space Observatory (Pilbratt et al. 2010). These observations allow us to map a wavelength range not easily accessed before, providing new insights into the dust and gas properties of Sh2-104.

2. Observations

The Herschel observations were taken on 2009 December 17 simultaneously with PACS (Poglitsch et al. 2010) and SPIRE (Griffin et al. 2010), as part of the guaranteed-time key-projects “Evolution of Interstellar Dust” of SPIRE (Abergel et al. 2010), and HOBYS of PACS (Motte et al. 2010). A 15′ × 15′ region was imaged with PACS at 100 and 160 μm (at resolutions of 10″ and 14″), and with SPIRE at 250, 350 and 500 μm (at resolutions of 18″, 25″ and 36″). Spectra were taken with the SPIRE-FTS long and short wavelength receivers (SLW and SSW, respectively) at seven different positions with sparse sampling, covering the 194–671 μm range. The resolution at the receivers’ central pixels varies between 17–19″ for SSW and 29–42″ for SLW. The data were reduced with the HIPE software version 2.0 with the latest standard calibration (Swinyard et al. 2010).

Figure 1 shows a color-composite image of Sh2-104 with PACS 100 μm (blue), SPIRE 250 μm (green) and SPIRE 500 μm (red). Different regions of interest, addressed in the following sections, are superimposed. We can see that the interior of the bubble is brighter in the PACS band, showing the hotter temperatures of the material in this region. On the other hand, outside...
the bubble the material is colder and emits stronger in the SPIRE bands.

3. Dust properties

We assume the dust emission in Sh2-104 can be modeled by an (optically thin) gray-body and that the emissivity of the dust grains can be fitted with a power law

\[ S_\nu \propto \Omega B_\nu (T) \kappa_\nu \left( \frac{\nu}{\nu_0} \right)^\beta N, \]

where \( S_\nu \) is the measured flux density at frequency \( \nu \), \( \Omega \) is the observing beam solid angle, \( B_\nu (T) \) is the Planck function for temperature \( T \) at frequency \( \nu \), \( \kappa_\nu \left( \frac{\nu}{\nu_0} \right)^\beta \) is the dust opacity, \( \beta \) is the dust emissivity index, and \( N \) is the dust column density (Hildebrand 1983). The value of \( \beta \) is believed to range between \(-1\) and \(-2\), but is an open issue that is still discussed, as is its dependence on the dust temperature (see e.g., Ossenkopf & Henning 1994). Dupac et al. (2003) and Désert et al. (2008) found an inverse relationship between \( \beta \) and \( T_{\text{dust}} \). On the other hand, Shetty et al. (2009) suggested that this result arises from noise and the combination of multiple emission components along the line of sight.

To determine the dust temperature structure of Sh2-104, we performed aperture photometry measurements on selected areas in the field. The apertures are shown in Fig. 1 and sample the interior of the bubble and the photodissociation region (PDR), including the UCH II region associated with IRAS 20160+3636. We used a single aperture (not shown) to account for the background emission. We fitted the PACS+SPIRE emission for all regions. These data represent the cold emission component, therefore we fitted a single temperature. We did this first allowing \( \beta \) to vary, and later fixing \( \beta = 1.5 \). The resulting \( T_{\text{dust}} \) and \( \beta \) values are shown in Table 1, their uncertainties are the formal 1\( \sigma \) values from the fitting procedure. Figure 2 shows examples of the fitted spectral energy distributions (SEDs) for the regions UCHII and PDR 1. Apart from the clear difference in \( T_{\text{dust}} \), it can be seen that for UCHII the PACS+SPIRE emission does not allow to sample the peak of the SED, which is reflected in a larger error in the fit. This is also seen in the four Interior # regions, suggesting the presence of a warmer emission component that contributes to the shorter-\( \lambda \) emission. Observations in the mid-IR can provide a constraint on this component, and a 2-temperatures fitting would be more appropriate to determine the \( T_{\text{dust}} \) and \( \beta \) values.

The temperatures throughout the PDR are between \( \sim 20 \) and \( \sim 30 \) K, and the UCH II region is marginally warmer. The region outside the bubble (region Outer) is the coldest, while the regions mapping the interior of the bubble are hotter, with an average temperature of \( \sim 40 \) K. The average \( \beta \) obtained is \( \sim 1.5 \), therefore the temperatures obtained from the fit with a fixed \( \beta = 1.5 \) do not significantly differ from those obtained with a free \( \beta \) (see Table 1).

![Composite image of Sh2-104. The field is \( \sim 13' \times 13' \) and shows PACS 100 \( \mu \text{m} \) emission in blue, SPIRE 250 \( \mu \text{m} \) in green and SPIRE 500 \( \mu \text{m} \) in red. Outlined are the regions where aperture photometry was applied, the region used for background substraction is off the map. Blue and red circles mark the pointings of the central pixels of SPIRE-FTS SSW and SLW, respectively. Their diameter represents their respective average resolution of 19'' and 35''.](image)

![Example of the SEDs obtained for regions UCHII (filled red squares) and PDR 1 (open blue squares). A dashed line is the fit with \( \beta = 1.5 \) and a solid line is the fit with \( \beta \) allowed to vary, with the resulting \( (T_{\text{dust}}; \beta) \) values shown in the respective color. The temperature difference between the regions is apparent, as is the difference in the peak sampling of the SED for PDR 1 and UCHII.](image)

<table>
<thead>
<tr>
<th>Region</th>
<th>( T_{\text{dust}} ) (K)</th>
<th>( \beta )</th>
<th>( T_{\text{dust}} ) (K)</th>
<th>( \beta )</th>
</tr>
</thead>
<tbody>
<tr>
<td>CS</td>
<td>25 ( \pm ) 3</td>
<td>1.5 ( \pm ) 0.3</td>
<td>25.5 ( \pm ) 0.9</td>
<td></td>
</tr>
<tr>
<td>Interior 1</td>
<td>43 ( \pm ) 12</td>
<td>1.2 ( \pm ) 0.4</td>
<td>35.7 ( \pm ) 2.1</td>
<td></td>
</tr>
<tr>
<td>Interior 2</td>
<td>39 ( \pm ) 9</td>
<td>1.3 ( \pm ) 0.3</td>
<td>33.5 ( \pm ) 1.7</td>
<td></td>
</tr>
<tr>
<td>Interior 3</td>
<td>47 ( \pm ) 15</td>
<td>1.0 ( \pm ) 0.4</td>
<td>34.2 ( \pm ) 1.9</td>
<td></td>
</tr>
<tr>
<td>Interior 4</td>
<td>35 ( \pm ) 8</td>
<td>1.3 ( \pm ) 0.4</td>
<td>30.4 ( \pm ) 1.4</td>
<td></td>
</tr>
<tr>
<td>Outer</td>
<td>17 ( \pm ) 2</td>
<td>1.7 ( \pm ) 0.4</td>
<td>17.3 ( \pm ) 0.4</td>
<td></td>
</tr>
<tr>
<td>PDR 1</td>
<td>20 ( \pm ) 2</td>
<td>1.7 ( \pm ) 0.4</td>
<td>21.7 ( \pm ) 0.6</td>
<td></td>
</tr>
<tr>
<td>PDR 2</td>
<td>22 ( \pm ) 3</td>
<td>1.7 ( \pm ) 0.3</td>
<td>24.1 ( \pm ) 0.8</td>
<td></td>
</tr>
<tr>
<td>PDR 3</td>
<td>23 ( \pm ) 3</td>
<td>1.7 ( \pm ) 0.4</td>
<td>24.6 ( \pm ) 0.8</td>
<td></td>
</tr>
<tr>
<td>PDR 4</td>
<td>22 ( \pm ) 3</td>
<td>1.5 ( \pm ) 0.3</td>
<td>22.3 ( \pm ) 0.7</td>
<td></td>
</tr>
<tr>
<td>PDR 5</td>
<td>22 ( \pm ) 3</td>
<td>1.5 ( \pm ) 0.3</td>
<td>22.9 ( \pm ) 0.7</td>
<td></td>
</tr>
<tr>
<td>PDR 6</td>
<td>24 ( \pm ) 3</td>
<td>1.8 ( \pm ) 0.3</td>
<td>26.0 ( \pm ) 0.9</td>
<td></td>
</tr>
<tr>
<td>PDR 7</td>
<td>25 ( \pm ) 3</td>
<td>1.7 ( \pm ) 0.3</td>
<td>26.9 ( \pm ) 1.0</td>
<td></td>
</tr>
<tr>
<td>PDR 8</td>
<td>26 ( \pm ) 4</td>
<td>1.7 ( \pm ) 0.4</td>
<td>28.6 ( \pm ) 1.2</td>
<td></td>
</tr>
<tr>
<td>UCHII</td>
<td>29 ( \pm ) 5</td>
<td>1.8 ( \pm ) 0.3</td>
<td>32.7 ( \pm ) 1.5</td>
<td></td>
</tr>
</tbody>
</table>

![Table 1. Results from aperture photometry.](image)
Figure 3 shows the distribution of the spectral indices $\beta$ vs. dust temperature, along with the relationships found by Dupac et al. (2003) and Désert et al. (2008). Within the uncertainties, our results agree with both relationships. We can also identify two groups, one with temperatures between 35 and 47 K and an emissivity index between 1.0 and 1.3, and the other with $T = 17–29$ K and $\beta = 1.5–1.8$. These two groups show that on average higher $\beta$ values are preferentially associated with colder material.

Although anti-correlation between $\beta$ and $T_{\text{dust}}$ is reported in the literature (e.g., Dupac et al. 2003; Yang & Phillips 2007; Désert et al. 2008), it is yet not clear which physical processes are behind it. Other authors studying this relationship examine the emission from different regions scattered in the sky. In contrast, we are finding this anti-correlation in the analysis of one contiguous complex object at a specific location in the sky. A $\beta \propto T_{\text{dust}}$ relationship may indicate a change of the dust properties (Stepnik et al. 2003). Grain finessiness in particular increases the emissivity index while keeping a relatively low temperature.

Fluffy grains result from grain coagulation and growth. The grain coagulation timescale and feasibility depend on factors such as the existence of ice mantles, grain size, and relative grain velocities.

We are finding the highest $\beta$ values and lowest $T_{\text{dust}}$ values toward the PDR of Sh2-104, which would imply then that the fluffiest and largest grains are located in the PDR. The question remains whether the PDR dust coagulated after the creation of the H II region, or if the birth of the H II region has destroyed the already coagulated dust located in the ionized cavity.

### 4. Gas properties

The seven pointings of the central pixels observed with SPIRE-FTS are marked with red (SLW) and blue (SSW) circles in Fig. 1. In total we detected the transitions described in Table 2. The richest spectrum (Fig. 4) was obtained towards pointing E, which targets the UCH II region associated with IRAS 20160+3636. The most prominent features are the CO J-ladder and the [N II] $^3P_1-^3P_0$ transition, an important ionized-gas coolant and a proxy for the H$\alpha$ flux (see e.g., Oberst et al. 2006).

<table>
<thead>
<tr>
<th>Transition</th>
<th>Rest $\lambda$ ($\mu$m)</th>
<th>$E_{\text{up}}$ (K)</th>
</tr>
</thead>
<tbody>
<tr>
<td>CO(4–3)</td>
<td>650.3</td>
<td>55.4</td>
</tr>
<tr>
<td>CO(5–4)</td>
<td>520.2</td>
<td>83.0</td>
</tr>
<tr>
<td>CO(6–5)</td>
<td>433.6</td>
<td>116.3</td>
</tr>
<tr>
<td>CO(7–6)</td>
<td>371.7</td>
<td>155.0</td>
</tr>
<tr>
<td>CO(8–7)</td>
<td>325.2</td>
<td>199.3</td>
</tr>
<tr>
<td>CO(9–8)</td>
<td>289.1</td>
<td>249.1</td>
</tr>
<tr>
<td>CO(10–9)</td>
<td>260.2</td>
<td>304.4</td>
</tr>
<tr>
<td>CO(11–10)</td>
<td>236.6</td>
<td>365.3</td>
</tr>
<tr>
<td>CO(12–11)</td>
<td>216.9</td>
<td>431.7</td>
</tr>
<tr>
<td>[N II] $^3P_1-^3P_0$</td>
<td>205.2</td>
<td>70.2</td>
</tr>
</tbody>
</table>

In the simple hypothesis of optically thin emission, we plotted $^{12}$CO and $^{13}$CO excitation diagrams, following the formulation of Johansson et al. (1984). An example is shown in Fig. 5 for pointing E. The slope of the linear fit for $^{12}$CO results in $T_{\text{L}} = 246 \pm 2$ K, and $T_{\text{H}} = 170 \pm 53$ K for $^{13}$CO. The total column densities derived are $N_{^{12}\text{CO}} = 10^{16.36\pm0.01}$ and $N_{^{13}\text{CO}} = 10^{15.4\pm0.1}$ cm$^{-2}$, respectively. The fits in Fig. 5 do not include all the measurements. The downturn seen in the lower levels of $^{12}$CO is interpreted as the optically thick/thin regime turnover, and is most likely a real physical feature and not an instrumental or calibration effect, because it is only seen for that species and not, for example, for $^{13}$CO. If the lines were emitted from an optically thick medium, their intensity would be underestimated, thus their respective column densities would also be a lower limit. Therefore, the optically thin assumption would hold for $^{13}$CO only for transitions higher than $J = 9 \rightarrow 8$, and those are the points included in the fit.

For $^{13}$CO on the other hand only the lines with wavelengths in the SPIRE-FTS SLW range are detected. These correspond to the $J = 5 \rightarrow 4$ to $J = 9 \rightarrow 8$ transitions. The $^{13}$CO lines in the SPIRE-FTS SSW wavelength range ($J = 10 \rightarrow 9$ to $J = 14 \rightarrow 13$) are detected as upper limits, because the line positions are found displaced from their expected positions, indicating that we are probably seeing some “outlier” noise features rather than the lines themselves. Therefore, we fit and show in Fig. 5 only the five $^{13}$CO transitions detected with SPIRE-FTS SLW. Following the reasoning of the previous paragraph, the $^{13}$CO transitions are most likely optically thin.
The two distinct gas temperatures obtained with 12CO and 13CO suggest two different gas components or a stratification of the emitting region. The colder component is traced by the optically-thin 13CO transitions at energy levels for which 12CO is optically thick, while the hotter component is traced by the more energetic, optically-thin 12CO transitions. Therefore it is likely that the colder gas is located at greater depths in the PDR than the hotter gas.

Assuming similar emitting volumes and beam filling factors as well as optically thin emission, our CO and 13CO column density values imply an abundance ratio [12CO]/[13CO], which is several times lower than the reported elemental value of 12C/13C ~ 69 (Wilson 1999). This would imply an enhancement of the 13CO isotopologue abundance.

However, several factors can contribute to the low 12C/13C abundance ratio we find. Perhaps the most important one would be the assumption of optically thin emission for 12CO. We used the high-energy transitions to derive its column density, and although in a first analysis they appear to be optically thin, it might not be the case (see e.g., Habart et al. 2010). To address this issue we will present and analyse PDR models of the Sh2-104 region in a forthcoming paper.

5. Summary

With Herschel PACS and SPIRE data we have analysed dust and gas properties of the bubble-shaped H II region Sh2-104. Aperture photometry of PACS+SPIRE images allowed us to derive the dust emissivity index β and the dust temperature throughout Sh2-104. We found two different groups, one at colder temperatures and higher β, and the other at warmer temperatures and lower β. We recover the inverse relationship found by Dupac et al. (2003) and Désert et al. (2008) with respect to the elemental value, but the uncertainties of different assumptions are still too large to confirm that result. In a follow-up paper we will show models of the PDR of Sh2-104, which will provide better constraints on the gas temperature, density and column density structure.

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