ExoMars TGO/NOMADUVIS vertical profiles of ozone: Part 1 – Seasonal variation and comparison to water

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Abstract. We present ~1.5 Mars Years (MY) of ozone vertical profiles, covering $L_s = 163^\circ$ in MY34 to $L_s = 320^\circ$ in MY35, a period which includes the 2018 global dust storm. Since April 2018, the Ultraviolet and Visible Spectrometer channel of the Nadir and Occultation for Mars Discovery (NOMAD) instrument aboard the ExoMars Trace Gas Orbiter has observed the vertical, latitudinal and seasonal distributions of ozone. Around perihelion, the relative abundance of both ozone and water (from coincident NOMAD measurements) increases with decreasing altitude below ~40 km. Around aphelion, localized decreases in ozone abundance exist between 25 and 35 km coincident with the location of modeled peak water abundances. High-latitude (>$\pm55^\circ$), high altitude (40–55 km) equinoctial ozone enhancements are observed in both hemispheres ($L_s \sim 350^\circ$–40°) and discussed in the companion article to this work (Khayat et al., 2021). The descending branch of the main Hadley cell shapes the observed ozone distribution at $L_s = 40^\circ$–60°, with the possible signature of a northern hemisphere thermally indirect cell identifiable from $L_s = 40^\circ$–80°. Morning terminator observations show elevated ozone abundances with respect to evening observations, with average ozone abundances between 20 and 40 km an order of magnitude higher at sunrise compared to sunset, attributed to diurnal photochemical partitioning along the line of sight between ozone and O or fluctuations in water abundance. The ozone retrievals presented here provide the most complete global description of Mars ozone vertical distributions to date as a function of season and latitude.

Plain Language Summary. We present over two years of new observations of the vertical distribution of ozone in the atmosphere of Mars. The ExoMars Trace Gas Orbiter spacecraft has been recording observations of the Martian atmosphere since 2018 to map the presence and changes in abundance of gases such as ozone by using the “Nadir and Occultation for Mars Discovery (NOMAD)” instrument. NOMAD continually observes the change in ozone abundance (among other gases) at different heights across much of the planet. These abundance profiles have revealed the presence of distinct layers of ozone enhancement at high altitudes in the atmosphere of Mars toward the polar regions and between spring and autumn in the southern hemisphere of Mars, discussed in detail in the companion article. We observe broad periods where often the abundance of ozone follows the abundance of water from ~10 km altitude up to ~50 km altitude, and other times when the two appear to be opposite in their variation with height. Our retrievals of ozone from NOMAD data provide the first coincident observations of ozone and water and provide previously unavailable information on the photochemistry of Mars.
1. Introduction

Ozone (O$_3$) is a highly reactive trace gas in the Martian atmosphere, where odd-hydrogen (OH, HO$_2$, H, H$_2$O$_2$) catalytic products of water vapor photolysis dominate Mars photochemistry (McElroy & Donahue, 1972; Parkinson & Hunten, 1972), and contribute the primary loss mechanism for Mars O$_3$. As a consequence, large spatial and temporal changes in Mars O$_3$ were quickly recognized to be associated with condensation-driven variations in atmospheric water vapor (Barth et al., 1973; Clancy & Nair, 1996; Kong & McElroy, 1977; Shimazaki & Shimizu, 1979). Mars atmospheric water variations result from seasonal/spatial variations in loss or supply to seasonal and residual polar ice caps (Fedorova et al., 2006; Jakosky & Farmer, 1982; Khayat et al., 2019; Pankine & Tamppari, 2019; Smith, 2004; Smith et al., 2009; Steel et al., 2014); and by the large spatial/temporal variations in ice cloud saturation conditions in the Mars atmosphere (Clancy et al., 1996; Montmessin et al., 2004; Navarro et al., 2014; Neary et al., 2020).

Mariner 7 made the first detection of O$_3$ in the Martian atmosphere (Barth & Hord, 1971) and later Mariner 9 showed O$_3$ to be seasonally variable (Barth et al., 1973). Subsequent missions and ground-based observations have characterized the geographical and seasonal variability of O$_3$ extensively (e.g., Blamont & Chassefière, 1993; Clancy et al., 1999, 2016; Fast et al., 2006; Lebonnois et al., 2006; Lefèvre & Krasnopolsky, 2017; Lefèvre et al., 2021; Montmessin & Lefèvre, 2013; Perrier et al., 2006). This has led to the well-established O$_3$ climatology of increased column abundances in the cold, dry atmospheric conditions of the fall and winter high-latitude regions; and reduced substantially in abundance over warmer, wetter low latitude and summer high-latitude regions. In addition, a distinct-low latitude, mid-altitude peak in O$_3$ arises from ~20 K colder atmospheric (and surface) temperatures present around Mars aphelion (Clancy & Nair, 1996), when Mars is furthest from the Sun in its eccentric orbit (at the current epoch, this occurs around Mars northern summer at a solar longitude, $L_s$, of 71°). These basic temporal and spatial behaviors of Mars atmospheric O$_3$ and H$_2$O column abundances are qualitatively reproduced in Mars global circulation models (GCM; e.g., Daerden et al., 2019; Holmes et al., 2017, 2020, 2018; Lefèvre et al., 2008, 2004).

However, most of the above behaviors are primarily characterized in terms of column O$_3$ measurements, due to the limited extent of vertical profile measurements for O$_3$ (Clancy et al., 2017; Gröller et al., 2018; Lebonnois et al., 2006; Olsen et al., 2020; Piccialli et al., 2021) or H$_2$O (Fedorova et al., 2018; Maltagliati et al., 2013). This is a significant limitation given the importance of vertical profile dependences in characterizing these photochemical and water vapor saturation (cloud microphysical) processes in any detail. Surprisingly, our understanding of such a fundamental relationship as the anti-correlation of Mars atmospheric O$_3$ and H$_2$O is obscured by uncertainties associated with column O$_3$ and H$_2$O measurement comparisons. Current model-data comparisons for these column data sets have implied significant (50%-100%) excesses in observed atmospheric O$_3$, while reproducing observed water column abundances more closely (Lefèvre et al., 2008). Heterogenous reduction of HO$_2$ on Mars water ice clouds has been proposed to account for such elevated O$_3$ abundances relative to homogeneous (gas-phase only) photochemical predictions (Lefèvre et al., 2008). Subsequent Mars O$_3$ observations have not yet indicated a heterogenous solution but the model underestimations of O$_3$ columns remain unresolved (Clancy et al., 2016; Daerden et al., 2019). It is possible, perhaps likely, that inaccurate model simulations of water vapor (and O$_3$) profiles contribute significantly to this issue (Clancy et al., 2017). Column O$_3$ measurements are heavily weighted by O$_3$ abundances in the lower scale height (~10 km), a region untouched by profile measurements to date. Column H$_2$O measurements do not provide significant guidance to temperature-sensitive, process-complicated simulations of cloud microphysics, necessary to simulate water vapor profiles accurately (e.g., Navarro et al., 2014).

Consequently, ExoMars Trace Gas Orbiter (TGO) sensitive and accurate determinations for O$_3$ profiles through UV (260 nm Hartley band) solar occultations (SOs) (this article), and water vapor profiles through near-IR SOs (NOMAD- Aoki et al., 2019; Vandaele et al., 2019; ACS- Fedorova et al., 2020), represent a major observational advance in the study of Mars photochemistry. On their own, Nadir and Occultation for Mars Discovery (NOMAD) O$_3$ profile retrievals yield a unique vertical dimension to our definition of Mars photochemistry. Coincident with NOMAD water profile retrievals, they enable an intrinsically diagnostic study of Mars O$_3$ (O + O$_3$) and HO$_2$ chemistry. Remaining contradictions on the presence of atmospheric methane (Korablev et al., 2019) and recent ExoMars detections of atmospheric HCI (Aoki et al., 2019; Korablev et al., 2021) lead to renewed importance in establishing a detailed understanding of Mars photochemistry.
Here, we present retrieved NOMAD O$_3$ profiles, with an emphasis on generating a global view of Mars photochemistry (O$_3$) in the vertical dimension. Preliminary comparisons of these NOMAD O$_3$ profile retrievals are provided with published NOMAD water profile retrievals (over $L_S = 163°$–345° in MY34, Aoki et al., 2019) and modeled water profiles (Holmes et al., 2020) extending over the remaining $L_S$ range of presented O$_3$ profiles. An initial comparison of modeled O$_3$ profiles indicates the degree to which existing models, without assimilated NOMAD water profile retrievals, correspond to the NOMAD O$_3$ profile measurements. Following studies of the combined NOMAD water and O$_3$ profile data sets will employ data assimilation of NOMAD water profile measurements to fully realize anticipated advances in Mars photochemistry associated with NOMAD O$_3$ and H$_2$O profile data.

Vertical profile measurements by the NOMAD spectrometer suite, aboard the ESA/Roscosmos ExoMars TGO mission have now been taken for over 1.5 Mars Years (MY), allowing the first glimpse into the vertical distribution of O$_3$ in the atmosphere of Mars. In this article, we present O$_3$ vertical profiles from April 2018 to November 2020, covering over 1.5 MY of O$_3$ measurements. This period includes the MY34 global dust storm, as well as the occurrence of a regional storm (Guzewich et al., 2019). The companion article to this study (Khayat et al., 2021) studies the high-altitude enhancement observed in more detail. The details of the NOMAD observations and the retrievals method are described in Sections 2 and 3 respectively, and the observational results are discussed in Section 4.

2. Observations

2.1. ExoMars and NOMAD-UVIS

The Ultraviolet and Visible Spectrometer (UVIS) is one of three channels that comprise the NOMAD spectrometer suite, aboard ExoMars TGO. The other two channels of NOMAD are infrared (IR) spectrometers: the SO channel and the Limb Nadir and Occultation (LNO) channel. A full description of the SO and LNO channels can be found in Vandaele et al. (2015). UVIS covers the spectral range of 200–650 nm with a spectral resolution of <2 nm and is capable of (mutually exclusive) observations in both SO and nadir viewing geometries using a single spectrometer with the aim to provide a comprehensive understanding of the climatology and vertical distribution of ozone, dust aerosols and ice aerosols in the Martian atmosphere. At the time of writing, the UVIS nadir ozone retrievals are not yet published. The occultation channel provides the capability to not only measure these species in the vertical and spatial dimension but to also measure changes on seasonal and diurnal timescales. A full description of the NOMAD-UVIS channel can be found in Patel et al. (2017).

TGO is in a near-circular 74° inclined orbit with a periapsis of 380 km and apoapsis of 420 km and an orbital period of approximately 2 h. NOMAD occultations come in four flavors: two different pointing modes and two different spacecraft maneuvers. When the $\beta$-angle, the angle between the orbital plane and the direction to the Sun, is $<\!20°$ TGO will actively track the Sun during an occultation, keeping the instrument boresight fixed on the Sun center. Within each planning period (approximately 1 month) around two of these sun-tracking occultations will be led using the UVIS boresight (i.e., the UVIS boresight is held centered on the Sun center); the rest are led by the NOMAD SO IR channel. At $\beta$-angles $>\!20°$, spacecraft angular momentum constraints prohibit suntracking and TGO rotates around the instrument pointing vector. This has minimal effect on UVIS-led occultations, however, on SO-led occultations, the UVIS occultation boresight (2 arcmin field of view) traces an arc around the solar disc at a distance from the solar center of approximately 0.3 solar radii. In general, there is no discernible impact on the calculated transmissions, since the UVIS boresight remains at the same radial distance from the center of the Sun, and the top of atmosphere reference transmission is defined over a range of altitudes using a linear regression method (see Section 2.3).

2.2. Data Set

Data acquired by the UVIS spectrometer from the start of the TGO science phase in April 2018 to November 2020 were used in this study, covering a period of just over one full Martian year (MY 34, $L_S = 163°$ to MY 35, $L_S = 320°$). This data set comprises 5,990 solar occultations (3,051 egresses, 2,939 ingresses) and corresponds to the Level 1.0 data made available through the NOMAD PI institute. The derived retrieval data data set can be accessed freely through the Open Research Data Online repository (Patel, 2021).
The geographical distribution of the occultation measurements is shown Figure 1. Due to TGO’s orbital characteristics and the terminator geometry for S0s, the majority of occultations are to be found clustered symmetrically around ±63° latitude bands (Figure 1) in the northern and southern hemispheres, with minimally slanted (vertical) atmospheric profiles. Relatively few occultations penetrate the equatorial latitudes; those that do tend to exhibit very long ground tracks reflecting a highly slanted atmospheric profile.

The latitudinal progression of the occultations over the duration of the current data set is shown in Figure 2 and is controlled by the changing β-angle of the spacecraft orbit. The β-angle varies in a cyclic pattern, moving between low and high β-angle phases. At low β-angles an orbit will typically consist of both an ingress and egress occultation, alternating between high-latitudes in the northern and southern hemispheres. As the β-angle increases, the occultations tend toward equatorial latitudes until the orbital geometry (β-angles >60°) prohibits occultation pointing and results in a period of no occultation observations.

Each occultation occurs over a period of between approximately 1–7 min, with equatorial occultations generally having longer durations as a result of the longer slant paths. The total integration time of each measurement was 45 or 75 ms. In the initial phase of the science mission (\(L_s = 163°–345°\)), the UVIS sampling frequency was set to 3 s providing a spectral resolution of <2 nm. Since the altitude resolution is also dependent on the occultation duration, the 3 s sampling frequency initially gave vertical resolutions that varied between 0.7 and 4.4 km. This somewhat coarse altitude resolution for occultations with short durations was a result of the 3 s measurement rhythm which is set by the CCD readout time and the data transfer time from UVIS to the NOMAD data processing unit. The CCD read time is fixed and cannot be changed but the data transfer time can be reduced by sending less data. After \(L_s = 345°\), UVIS was modified...
to spectrally bin the CCD array from 1,024 pixels to 128 pixels, a factor of eight reduction in the data volume, allowing an increase in the sampling frequency to 1 s, a factor of three improvement, and a corresponding increase to the altitude resolution in the range 0.2–1.4 km. The spectral features of ozone, water ice aerosols, and dust aerosols are all broad band features, as such an increase in the spectral resolution to 4 nm had no impact on the retrievals.

A great benefit of the SO technique is that it does not rely on laboratory radiometric calibration, and it is a self-calibrating measurement. During a SO, the line of sight (LoS) between the instrument and the Sun is maintained as the LoS passes through the atmosphere, from a very high altitude (200–250 km) down to the surface (or vice versa). Relative transmission is then calculated using a reference spectrum obtained from high up in the atmosphere where no atmospheric contribution is expected (typically \( \sim 110 \) km), resulting in a consistently self-calibrated transmission profile through the atmosphere for each occultation. Additionally, no correction for UV airglow is required, given the significantly higher signal intensity from the line-of-sight Sun illumination with respect to atmospheric emission at these wavelengths.

## 2.3. Atmospheric Transmittance

This transmission is defined as the spectral radiance observed through the atmosphere at a given tangent altitude, \( z \), above the surface, \( S(\lambda, z) \), taken as a proportion of the total (above atmosphere) solar radiance, \( S_0(\lambda) \):

\[
T(\lambda, z) = \frac{S(\lambda, z)}{S_0(\lambda)}
\]  

(1)

\( S_0(\lambda) \) is determined using the algorithm defined by Trompet et al. (2016), as discussed in Vandaele et al. (2018) as the standard occultation procedure for NOMAD transmissions. This method uses an iterative algorithm to define a mean value for spectra above 120 km, where it is assumed there is no absorption due to Martian atmospheric gases, thus defining a unique reference for each occultation.

The spectral transmission obtained from the data can be related to the number density of all-absorbing agents in the LoS through the atmosphere through the Beer-Lambert law:

\[
S(\lambda, z) = S_0(\lambda) e^{-\Sigma_i \sigma_i N_i(z)}
\]  

(2)

where, \( \sigma_i \) is the cross-section of the \( i \)th absorbing chemical species and \( N_i = \int n_i \cdot dx \) is the integrated total number density per species along the LoS column. Given the small field of view, the signal contribution from forward scattering was modeled to be approximately two orders of magnitude lower than the direct solar signal and is considered negligible. A single transmission spectrum yields a transmission ratio for each wavelength in the detector range of 200–650 nm at a spectral resolution varying between approximately 1.5–3.5 nm (depending on observation configuration). An example of the corresponding set of transmission spectra obtained during a single occultation is shown in Figure 3, representing the observed transmission through different layers of the Martian atmosphere.

## 3. Retrieval Method

### 3.1. Spectral Inversion

The transmission spectra observed by the instrument at successive tangent altitudes above the surface are converted into number densities of different species by a spectral inversion process. These number density profiles, or slant densities, are a measure of the total integrated number of a given species in the line-of-sight column of each observation.
3.1.1. The Line-of-Sight Points

A schematic of the location of the LoS between the TGO spacecraft and the Sun for each observation is shown in Figure 4. The tangent height point, denoted by $X_{tan}$, is defined as the point along the LoS that has the shortest distance from the LoS to the center of the planet.

The geometry parameters recorded in each observation specify the height of TGO above the center of the planet $R_{sat}$, together with the planetocentric latitude $\theta_{sat}$ and longitude $\phi_{sat}$ of the TGO position. They also give the height $H_{sat}$ of the tangent point above the MGM1025 areoid (Lemoine et al., 2001), defining the term “altitude” used throughout this article.

It is useful to define a cartesian coordinate system with origin at the center of the planet, the positive $x$ axis cutting the 0° meridian, and the positive $z$ axis passing through the north pole. Then, using the geometry of the ellipsoid, it is possible to calculate the cartesian coordinates $x_{sat} = [x_{sat}, y_{sat}, z_{sat}]$ of the point on the
ellipsoid corresponding to the position \((\theta_{sat}, \phi_{sat})\). This is done by first calculating the corresponding radial distance \(r_{sat}\) of the point from the center of the ellipsoid:

\[
\begin{align*}
    r_{sat} &= \sqrt{\left(\frac{\cos(\theta_{sat})}{a}\right)^2 + \left(\frac{\sin(\theta_{sat})}{c}\right)^2}, \\
    x_{sat} &= r_{sat}\cos(\theta_{sat})\cos(\phi_{sat}), \\
    y_{sat} &= r_{sat}\cos(\theta_{sat})\sin(\phi_{sat}), \\
    z_{sat} &= r_{sat}\sin(\theta_{sat}).
\end{align*}
\]

(3)

Since \(x_{sat}\) is collinear with the center of the planet and the location of the satellite \(X_{sat}\), a simple rescaling then gives the cartesian coordinates of the satellite location:

\[
X_{sat} = \frac{R_{sat}}{r_{sat}} x_{sat}.
\]

(4)

Similarly, it is possible to calculate the cartesian coordinates \(X_{tan}\) of the tangent height point, by rescaling the cartesian coordinates of the point \(x_{tan}\) on the ellipsoid corresponding to the position \((\theta_{tan}, \phi_{tan})\). The procedure is as described above for the case of \(X_{sat}\), except that the final rescaling is given as follows:

\[
X_{tan} = \frac{r_{tan}}{r_{tan}} x_{tan} + \frac{H_{tan}}{r_{tan}} X_{tan}.
\]

(5)

Once the cartesian coordinates of \(X_{sat}\) and \(X_{tan}\) are known, the LoS points are readily calculated as points that lie along the line joining \(X_{sat}\) and \(X_{tan}\). It is noted that the effect of refraction in low-pressure atmospheres has been found to be small by previous studies (e.g., Bellucci et al., 2009; Bétrémieux & Kaltenegger, 2015; Hubbard et al., 2001) and therefore refraction effects are not considered here.

3.1.2. GCM and Re-Analysis

To perform the spectral fitting (described in the next section), initial estimates of aerosol opacity and ozone abundance along the LoS points are determined. These initial estimates are not strictly needed to perform the initial opacity retrieval, but are used for subsequent mixing ratio calculations, and were therefore used as initial estimates to optimize the fitting process since they were available. These values were extracted from outputs of the Open University (OU) modeling group Mars GCM coupled to the Analysis Correction assimilation scheme that has a strong heritage in data assimilation on Mars (Holmes et al., 2020, 2018, 2019; Lewis et al., 2007; Montabone et al., 2006, 2014; Streeter et al., 2020). The GCM is comprised of physical parameterisations (Forget et al., 1999) and a photochemical module (Lefèvre et al., 2008, 2004) shared with the Laboratoire de Météorologie Dynamique (LMD) modeling group, alongside a spectral dynamical core and semi-Lagrangian advection scheme (Newman et al., 2002). It includes the latest sub-models to provide the most realistic modeling of the turbulent processes in the planetary boundary layer (Colaïtis et al., 2013), and a water-ice cloud microphysics package (Navarro et al., 2014) that includes radiatively active water ice clouds and supersaturation. The OU modeling group GCM has been developed in a collaboration between the OU, the LMD, the University of Oxford, and the Instituto de Astrofísica de Andalucia. Initial \(O_3\) values originate from a model run that assimilated \(O_3\) column data measured by the SPICAM instrument during MY 27 (Holmes et al., 2018). Although this model run is for a different year, it produces a good initial estimate of ozone for the beginning of the model run assuming that the seasonal \(O_3\) distribution is repeatable from year to year. For temperature, dust and aerosol fields, a different model run was used, that assimilated contemporaneous temperature profiles and column dust data measured by the Mars Climate Sounder (MCS) instrument (Streeter et al., 2020). The basis of the dust distribution for this period is presented in Montabone et al. (2020).

3.1.3. Cross-Sections

The three primary absorption features acting on the observed atmospheric transmissions in the observational wavelength region between 200 and 650 nm are those due to carbon dioxide (CO\(_2\)), aerosols, and \(O_3\). Only the cross-section of \(O_3\) is of relevance in this study since wavelengths close to CO\(_2\) absorption are...
not considered, and the O$_3$ cross-section values at 218 K of Malicet et al. (1995) are used here. The spectral absorption characteristics of aerosols exhibit a continuum attenuation in the transmission spectra so an increase in aerosol abundance can in principle be modeled as a fit to any part of the spectra well separated from the reserved O$_3$ band. Due to the spectral shape changes at longer wavelengths due to aerosol particle size these wavelengths are avoided for modeling the aerosol content in the spectra.

3.1.4. Spectral Fitting

A least squares ($\chi^2$) procedure involving minimizing the sum of the squares of the residual differences between the measured transmissions and the best fit model across a range of wavelengths, is performed utilizing the Levenberg-Marquardt algorithm to obtain the best fit solution. Modeled values for aerosol opacity and ozone abundance are extracted for each LoS from the assimilated data set described earlier as a priori values for the iterative fitting process, purely to reduce the number of iterations required.

The first iteration of the fitting is performed in the region between 320 and 360 nm (i.e., outside the O$_3$ absorption band) for the aerosol abundance determination (Clancy et al., 2016), with an optional additional region at 220–230 nm to adjust for CO$_2$ Rayleigh scattering (Ityaksov et al., 2008) across the O$_3$ band when required. Due to the reduced performance of the detector in this shorter wavelength fit region, in instances where the measured transmission in this region results in negative retrieved values of CO$_2$ Rayleigh scattering, an aerosol re-fit constrained solely to the 320–360 nm region is performed around the initial obtained aerosol value and the re-fitted aerosol fit is extrapolated over the O$_3$ band. All parameters in the model were initially fixed except for the CO$_2$ Rayleigh scattering contribution and aerosol abundances along with the Angstrom coefficient, which were treated as free parameters over each iteration to determine an aerosol abundance, after CO$_2$ Rayleigh scattering, on convergence. The fitting procedure for O$_3$ abundance then follows, with updated CO$_2$ Rayleigh scattering and aerosol values fixed in the model, within the wavelength range 240–320 nm on the positive gradient of the O$_3$ Hartley band. This wavelength range was used since it permits an accurate representation of O$_3$ abundance based on a measure of the depth of this feature in the spectrum, whilst avoiding the CO$_2$ Rayleigh scattering fit region and the less well-characterized transmission values below 220 nm, where there is greater noise in the spectra and straylight contamination represented by a growth in relative errors.

This optimal wavelength range over which to perform the O$_3$ fit coincides with the magnesium II h and k doublet solar emission at 280 nm, which is often used as a diagnostic of solar activity (e.g., Snow et al., 2019). As the determination of the Hartley band depth in any given spectrum may be influenced by the residual presence of this feature as the LoS moves across the solar disk, the values at 280 nm are discarded and an interpolation is performed either side of this wavelength.

Examples of the fitting procedure results on two example occultations are shown in Figure 5, where the overall composite fit has been combined to construct the final modeled transmission. The O$_3$ model sometimes deviates from the observation below 255 nm primarily due to known straylight within the system causing a non-real excess in the recorded transmission values below this threshold. In general, the results <220 nm are not considered here due to the loss in responsivity of the CCD at these wavelengths (Patel et al., 2017). A lower altitude cut-off is applied to all profiles at the 1% transmission level at a wavelength of 255 nm, below which the signal on the CCD detector becomes comparable to the noise. Using the total abundance extracted from the O$_3$ fit of the modeled spectra coupled with the cross-section values the total integrated number density (slant density) is then calculated by applying the Beer-Lambert law defined in Section 2.3.

3.2. Vertical Inversion

The slant density of an absorbing species gives a measure of the total integrated number of molecules per unit area along the LoS of each measurement. To convert this slant density profile into an associated set of local densities, or a true vertical profile, a vertical inversion of the slant densities is performed.

To achieve this inversion a standard onion peeling process is used (Auvinen et al., 2002; Quémerais et al., 2006; Rodgers, 2000), a procedure that requires the observed two-dimensional atmospheric slice to be split into concentric layers. This structure, illustrated in Figure 4, is provided by the observational data, each layer being defined by the tangent altitude of the LoS in successive measurements above the same areoid.
This results in a single measured transmission spectrum corresponding to each atmospheric layer. A spherically symmetric atmosphere is assumed in the plane defined by the observed lines of sight. It should be noted that $O_3$ is unlikely to be distributed in a spherically symmetric manner along the line of sight along the terminator (see e.g., Piccialli et al., 2021 for a discussion of this issue), and therefore this assumption will have some associated (unavoidable) error, that cannot be quantified in a meaningful way at this time. The separate path lengths that every measurement's LoS takes through each of the atmospheric layers are then calculated geometrically:

$$S = \sqrt{z^2 - r_i^2}$$  \hspace{1cm} (6)

In combination with an interpolation between the layers, an Abel integral provides the coefficients of the individual elements of the weighting function matrix, $E_A$, which defines all combinations of path lengths per layer in the two-dimensional observational grid:

$$A_{ij} = \int_{z_{min,i}}^{z_{max,i}} f(z) \frac{z}{\sqrt{z^2 - r_i^2}} \, dz$$  \hspace{1cm} (7)

where $f(z)$ is the interpolated pressure as a function of altitude $z$, $z_{min,i}$ and $z_{max,i}$ are the altitude limits covered by the $i$th layer. All (virtual) path lengths constructed via this method that are external to the atmosphere will not contribute to the total observed abundance in any given atmospheric line-of-sight and therefore are given a zero weighting in the final matrix.

With the weighted path lengths encoded in the matrix $A$ the local densities, $D$, forming the vertical profile can be calculated from the known slant densities, $N$, and their associated uncertainties, $\epsilon$, by a matrix inversion of the form:

$$D = N \cdot A^{-1} + \epsilon$$  \hspace{1cm} (8)

This then provides a final vertical profile from the observed slant densities. An example of a calculated $O_3$ vertical profile obtained through this method and its associated slant density profile are shown in Figure 6. Measured $O_3$ number densities range from $\sim 10^{10}$ in the lower altitudes down to $\sim 10^7$ below which the noise floor inherent in the retrieved slant densities is reached, demonstrated here by the shaded envelope in Figure 6. Since in the present study only upper limits can be defined above $\sim 60$ km, the results presented here will focus on the detections confirmed below this altitude.
3.2.1. Uncertainties on the Vertical Profile

The main source of uncertainty in the occultation data arises from instrumental noise—each spectrum has an associated instrumental error value, with the median value of this error within each occultation being less than 1%.

To determine the associated uncertainty in the retrieval, the 1−σ experimental uncertainty from the random instrumental noise associated with every observed transmission, ε, is used in order to calculate a weighted deviation (W) between the model, as evaluated from the Beer-Lambert equation (Section 2.3), the fitted model data \( M(\lambda) \), and the observation data \( x(\lambda) \):

\[
W = \frac{M(\lambda) - x(\lambda)}{\epsilon}
\]  

The square of these weighted residuals are the \( \chi^2 \) values that, through successive iterations, are minimized in order to arrive at a final best fit solution to the modeled data, the associated cost function can be expressed as:

\[
V = A^T C^{-1} A
\]

The covariance matrix, \( C \), in this process describes the covariance of every combination of pairs of measurements; the diagonal elements, therefore, represent the self-covariance (the variance) of each observation with itself, the individual errors on the obtained slant density measurements are thus obtained from the square-root of the diagonal of the covariance matrix. An averaging kernel, \( K \), is constructed from the weighting function and covariance matrices, whose \( i \)th row describes the weighting influence of each atmospheric layer on the \( i \)th LoS derived number density:

\[
K = \left( V + \lambda_s L^T L \right)^{-1} (A^T C^{-1})
\]

where the matrix \( L \), whose contribution is governed by the smoothness coefficient \( \lambda_s \), is given by:
This kernel represents the quality of the retrieval per retrieved number density value in the final vertical profile.

The uncertainty in the number densities of the final vertical profile is derived in an analogous way to the number densities themselves, as shown below. In an iterative procedure, the respective errors of each observed slant density are combined in quadrature with the sum of the previously calculated errors on all higher layers, weighted by the proportional path length that those layers represent in the current line of sight column:

$$ L = \frac{1}{h^2} \left[ \begin{array}{cccccc}
-1 & 1 & 0 & 0 & \cdots & 0 \\
1 & -2 & 1 & 0 & \cdots & 0 \\
0 & 1 & -2 & 1 & \cdots & 0 \\
\vdots & \vdots & \vdots & \vdots & \ddots & \vdots \\
0 & \cdots & 0 & 1 & -2 & 1 \\
0 & \cdots & 0 & 0 & 1 & -1 \\
\end{array} \right] $$

(12)

This quantity is then scaled based on the path length of the layer in question (the lowest layer) in the current column to obtain a final error on the number density in that layer of the vertical profile. Median errors on the calculated number density are found to be up to 7% between an altitude range of 20–50 km.

4. Results and Discussion

4.1. O$_3$ Vertical Distribution

Figure 7 shows the seasonal variation of the O$_3$ vertical profiles from the NOMAD-UVIS retrievals for observations made in the northern hemisphere and the southern hemisphere. Generally, retrieved NOMAD O$_3$ profiles extend to 50–70 km altitudes, based on NOMAD O$_3$ sensitivity limits and Mars upper-level O$_3$ abundances (previous section, e.g., Figure 6). It should also be noted that these plots show the vertical profile over varying latitudes—in some cases, moving from the polar regions to equator in a small $L_S$ period. Thus, care must be taken in interpreting these plots, as latitudinal gradients can appear mapped into temporal ($L_S$) changes.

Given the strong photochemical anticorrelation of O$_3$ with H$_2$O, we also present the water profiles simultaneously retrieved with the IR channel of NOMAD (Aoki et al., 2019), providing the first direct comparison of contemporaneous and co-located O$_3$ and H$_2$O vertical profiles. Unfortunately, at the time of writing NOMAD water profiles retrievals are only publicly available up to $L_S = 345^\circ$ in MY34 (February 2019), and therefore the water profiles shown in Figure 7 only contain NOMAD water retrievals to $L_S = 345^\circ$.

To provide a comparative water vapor data set for the remaining period through to $L_S = 320^\circ$ in MY35 (November 2020), we performed an assimilation of MCS v5.3.2 temperature profiles and column dust retrievals (Kleinböhl et al., 2017) using the OU modeling group Mars GCM described in Section 3.1.1. An ensemble of assimilation runs was conducted over the whole time period, in which the parameters of albedo and thermal inertia of perennial surface ice were varied to select an optimum water vapor simulation. Navarro et al. (2014) adopted this approach in a forward GCM study, to model Mars atmospheric water column variations. The selected assimilation run was chosen to optimize water vapor comparisons with respect to NOMAD water vapor retrievals toward the end of the period for which they are available (late 2018, when the MY34 dust storm effects have largely abated). Comparison with as yet unpublished NOMAD water profile observations for this period show the modeled water to be qualitatively consistent with the observed water profiles.
Figure 7. (top) Northern hemisphere altitude profiles for (a) NOMAD retrievals of O$_3$; (b) modeled O$_3$ based on initial conditions derived from SPICAM observations and evolved through time considering assimilated temperature and dust observations from MCS; (c) retrieved and modeled H$_2$O abundances (H$_2$O is modeled after $L_s = 345^\circ$). NOMAD retrievals of H$_2$O are from Aoki et al. (2019). (bottom) The equivalent as above, but for the southern hemisphere. The UVIS O$_3$ retrieval profiles cover the period $L_s = 163^\circ$ in MY34 to $L_s = 320^\circ$ in MY35 and consist of 5,990 individual profiles and 677,473 individual spectra. The wide and narrow shaded periods represent the timing of the MY34 global dust storm and regional storm respectively. Altitude values are with reference to the Mars areoid, and the latitude that each profile is taken at is indicated by the top plot.
The bottom panels of Figure 7 present this NOMAD-model composite description of water profiles corresponding (in position and \(L_s\)) to retrieved NOMAD \(O_3\) profiles in the top panels. It should be noted that the average difference in time between the model output time and each specific observation time is 12 min. This provides a partial comparison of NOMAD water vapor and \(O_3\) profile measurements, for the published dust storm period of NOMAD observations (\(L_s = 190°–345°\) in MY34). It also allows a predictive (model) description of the very distinct behavior of Mars water vapor centered on the aphelion climate of Mars (\(L_s \sim 0°–180°\)) in MY35.

NOMAD-retrieved \(O_3\) profiles are presented in the top panels of Figure 7 and compared to simulated \(O_3\) profiles in the middle panels. These simulated \(O_3\) profiles are strictly based on the MCS temperature and dust assimilation modeling described above. Consequently, they correspond to modeled, not observed, atmospheric water vapor distributions for the entire presented \(L_s\) range in Figure 7. These simulated water vapor profiles are constrained by assimilated MCS temperature and dust profiles, as described above, such that they more closely correspond to contemporaneous water vapor saturation conditions, including the MY34 dust storm, high-latitudes, and the aphelion period.

Nevertheless, the model-NOMAD comparisons for \(O_3\) profiles in the top two panels of Figure 7 still retain significant uncertainties associated with GCM simulations of Mars water vapor (Daerden et al., 2019; Navarro et al., 2014). Direct assimilation of NOMAD retrieved \(H_2O\) profiles should allow much more quantitatively direct (and photochemically diagnostic) comparisons of model and observed \(O_3\) profiles. For current purposes, these model/data \(O_3\) comparisons allow a preliminary assessment of the detailed seasonal and spatial variations of the vertical distribution of Mars \(O_3\) from NOMAD, in the context of a best-effort approach for modeling Mars atmospheric water profile behavior based on assimilated dust and temperature profiles.

The effect of the MY34 global dust storm (approximately \(L_s = 190°–250°\)) on the \(O_3\) retrievals, through reduced limb transmission, is most noticeable in the increasing lower altitude cut-off due to the sharp increases in dust loading (e.g., Liuzzi et al., 2020; Montabone et al., 2020; Smith, 2019). In the equatorial and mid-latitudes, retrievals are only possible down to \(\sim 30\) km, allowing only mid-to-upper level regions of atmospheric \(O_3\) to be detected. At higher latitudes, where lower dust content is present, \(O_3\) retrievals are possible to altitudes below 10–20 km. Elevated water abundances at high altitudes (>30–50 km) are observed in both northern and southern hemispheres during the MY34 (2018) global dust storm after \(L_s = 190°–200°\) (Aoki et al., 2019; Fedorova et al., 2020), similar to behavior observed and modeled during the MY28 (2007) global dust storm (Heavens et al., 2018). This leads to correspondingly reduced \(O_3\) abundances above 20–30 km between \(L_s = 200°–300°\) in both hemispheres, in both MYs, associated with the strong photochemical anti-correlation between \(O_3\) and water. As noted previously, the high-altitude enhancements of \(O_3\) observed in the southern hemisphere (and to a lesser extent the northern hemisphere) from spring equinox are covered in detail in the companion article to this study (Khayat et al., 2021) and are not discussed here. A comparison of the difference in MY34/35 \(O_3\) distribution during the global dust storm period is given in Section 4.3; however, care must be taken when comparing periods in Figure 7, since the latitude of the observations vary significantly and will be different between Mars Years.

Altitude-resolved NOMAD \(O_3\) abundances below 15 km altitudes represent new measurements and will prove significant in assessing previous column (nadir) measurements of Mars ozone (e.g., Perrier et al., 2006). Such column measurements are often dominated by maximum lower atmospheric \(O_3\) abundances, particularly at high-latitudes, that may contribute disproportionately to existing model-data disagreements for column \(O_3\) comparisons (Clancy et al., 2016). Peak near-surface \(O_3\) densities (and mixing ratios) are generally observed to be present throughout the Mars Year and occur, in part, due to reduced \(HO_3\) production from \(H_2O\) UV photolysis (associated with increasing \(CO_2\) extinction of UV radiation, e.g., Montmessin & Lefèvre, 2013).

This near-surface \(O_3\) layer is particularly evident at southern winter latitudes, when cold aphelion temperatures lead to the aphelion cloud belt (ACB) and much-reduced transport of water vapor into the southern winter hemisphere (Clancy et al., 1996), and thus minimum \(HO_3\) loss rates for \(O_3\) at all altitudes (Daerden et al., 2019; Lefèvre et al., 2004). Figure 7 demonstrates this behavior in, respectively, NOMAD retrievals and assimilation modeling of \(O_3\) and \(H_2O\) abundances in the southern hemisphere lower atmosphere.
The northern hemisphere winter is generally obscured due to the presence of dust from the MY34 global dust storm. Increasing \( O_3 \) abundances below \( \sim 15 \) km altitudes are retrieved between \( L_s = 190^\circ - 250^\circ \), associated with northern winter. But overall, reduced \( O_3 \) and increased \( H_2O \) abundances (both, in this case, from NOMAD retrievals) are present relative to the aphelion, southern winter period.

Figure 8 shows the ratio of the observed to modeled ozone abundance shown in Figure 7 for (a) the northern hemisphere and (b) the southern hemisphere. Lighter (darker) shading indicates stronger (weaker) agreement between observation and model abundances; higher ratio values indicate that modeled values are lower than that observed by NOMAD.
in our understanding of the HO$_3$ processes in the atmosphere of Mars (see e.g., Lefèvre et al., 2021), which could contribute to such a discrepancy in O$_3$ abundance. Model assimilation of the observed NOMAD water profiles, once publicly available, will address this modeled O$_3$ abundance deficiency.

4.2. Latitudinal Distribution

Figure 9 shows the NOMAD ozone profile retrievals, plotted as a function of latitude for $L_S$ periods of 20°. Toward perihelion ($L_S = 251°$) in northern winter, O$_3$ is generally present at low altitudes at mid/high-latitudes. A comparison of the northern autumnal equinox periods for MY34 and MY35 ($L_S = 160°–180°$ and 180°–200°) indicates that observed O$_3$ abundances extend to higher altitudes in the non-global dust storm year (MY35). Mid/high-latitude O$_3$ abundances above ~20 km are reduced during the MY34 global dust storm, due to the elevated water abundance above these altitudes associated with this storm (e.g., Aoki et al., 2019; Fedorova et al., 2020). The period $L_S = 180°–200°$ presents an especially high number of observations in both years, and thus is most appropriate for this comparison. In MY34 during this period, O$_3$ reaches an altitude of 20–30 km in the southern polar regions, at the onset of the global dust storm. In the MY35 equivalent period, O$_3$ reaches an altitude of 30–40 km, highlighting the reduction of O$_3$ abundances in this region in MY34 associated with dust-driven increases in water vapor at higher altitudes. Northern mid-latitudes (~30°–45°N) also exhibit higher O$_3$ abundances in the $L_S = 180°–200°$ period in MY35 with respect to MY34, but limited observations during the MY34 period make it difficult to compare this latitude region further. Following cessation of the regular period of dust activity centered near $L_S = 330°$ in MY34 (e.g., Kass et al., 2016; Smith, 2004), the development of increased O$_3$ abundances to altitudes of ~35 km is observed in the southern polar regions in MY34 at $L_S = 340°–0°$.

The notable low-to-mid latitude increases in 20–50 km altitude O$_3$ over the extended northern summer season ($L_S = 40°–120°$) reflect the aphelion saturation conditions for water vapor above ~20 km altitudes. This orbitally driven cooling of the global Mars atmosphere leads to the formation of the ACB (Clancy et al., 1996; Montmessin et al., 2004; Smith, 2004) and sharp increases in O$_3$ abundances above 20 km (Clancy & Nair, 1996; Lefèvre et al., 2004) at this time. These vertically extended increases in aphelion O$_3$ abundances continue until $L_S \sim 120°$, after which O$_3$ abundances above 20–30 km begin to decrease with increasing atmospheric temperatures and water vapor. The ACB similarly declines over this period (Smith, 2004).

At $L_S = 40°–60°$ in Figure 9, meridional circulation during the transition from equinoctial to solstitial circulation is evident in the latitudinal O$_3$ distribution, where O$_3$ formation along the flow path of O atoms highlights the meridional circulation. The downwelling branch of the main Hadley cell typical of this time of year (e.g., Barnes et al., 2017; Montmessin & Lefèvre, 2013) shapes the O$_3$ distribution in the southern hemisphere and manifests as a latitudinal contour of O$_3$ abundance beginning at a latitude of ~30°S and altitude of ~45 km down to a latitude of ~55°S and altitude of 20 km. There is evidence of a thermally indirect (Ferrel) cell at high northern latitudes at altitude, manifesting as an O$_3$ abundance contour extending from a latitude of ~50°N and 30 km down to a latitude of ~70°N and ~20 km. Progressing toward solstitial circulation, in the $L_S = 60°–80°$ period there is a possible signature of the elevation of the thermally indirect (Ferrel) cell flow path, being elevated by approximately 10–15 km, with the O$_3$ contour highlighting the circulation flow path spanning from a latitude of ~35°N at 45 km down to ~65°N at ~30 km. It should be noted that the meridional circulations highlighted here do not represent the strength or extent of the circulation nor the transport of O$_3$, but serve only to highlight the (possibly weak) circulation paths where the formation of sufficient O$_3$ is permitted and thus made observable.

4.3. O$_3$ Distribution Compared to H$_2$O

O$_3$ and H$_2$O are always photochemically anti-correlated in terms of HO$_3$ chemistry. Care must be taken when using the terms “correlated” and “anti-correlated” in studies such as presented here, to avoid misunderstanding regarding the relationship between O$_3$ and H$_2$O. While these two species are fundamentally anticorrelated from a chemical reaction perspective, the variation of these species over altitude is not necessarily constrained in the same manner. For example, it is possible for both O$_3$ and H$_2$O to follow the same decrease in relative abundance over altitude, while remaining photochemically anti-correlated. Competing photochemical processes will dominate at different altitudes (e.g., solar UV at high altitudes enhancing...
Figure 9. Latitudinal distribution of the NOMAD O₃ retrievals over 1 Martian year. Each plot shows the vertical profiles obtained during an $L_s = 20^\circ$ period.
destruction pathways and reducing in effectiveness toward the surface). A common assumption that follows the statement of anti-correlation between the two species is that their profile shapes will also be anti-correlated—this is not necessarily always the case. NOMAD is capable of measuring simultaneously $O_3$ and $H_2O$ vapor profiles to demonstrate this through observations; here we discuss the relationship of the two species regarding their relative vertical distribution.

Figure 10 presents the comparison of $O_3$ and $H_2O$ as a function of altitude and latitude. Figure 10 shows the distribution of NOMAD $O_3$ and $H_2O$ retrievals from $L_S = 163^\circ$ onwards for the remainder of MY34, and also the NOMAD $O_3$ compared to the modeled $O_3$ (following MCS dust and temperature assimilation) for MY35 up to $L_S = 280^\circ$.

At perihelion in MY34 between $L_S = 200^\circ$–$320^\circ$ there is a greater relative density of NOMAD-retrieved $H_2O$, particularly at high altitudes (>30 km) in both hemispheres, due to the ~20 K increase in average atmospheric temperatures associated with the ~40% increase in solar flux (Clancy et al., 1996), and due to the related intense dust activity during the global and regional storms (Liuzzi et al., 2020; Smith, 2019). During this period, $O_3$ is confined to lower altitudes, with the altitude of peak $O_3$ values in general being coincident with peak $H_2O$ values. $O_3$ and $H_2O$ show a similar overall trend in increasing relative abundance with decreasing altitude. Either side of this period (i.e., $L_S = 160^\circ$–$200^\circ$ and $320^\circ$−$0^\circ$), water abundance is generally lower, and distinct increased $O_3$ abundance is observed at high-latitudes in both hemispheres, coincident with latitudes where $H_2O$ abundances are decreased. This anti-correlation in distribution continues into MY35, with the latitudinal anti-correlation immediately evident in the $L_S = 0^\circ$–$40^\circ$ period, and continuing albeit to a lesser degree throughout the rest of the year.

Toward aphelion in MY35 during the northern summer there is an overall lower density of modeled $H_2O$, and a higher density of $O_3$, consistent with the generally colder aphelion season due to decreased solar insolation and dust loading. In MY35, $O_3$ retrievals were possible to lower altitudes than in MY34 due to the absence of storm-related high dust loading. From $L_S = 0^\circ$–$160^\circ$, $H_2O$ is confined below 40 km in both hemispheres, showing dry south polar regions to the minimum altitudes observed by NOMAD (~5–10 km). From $L_S = 40^\circ$–$120^\circ$, $O_3$ density is consistently high across all latitudes up to altitudes of ~50 km, consistent with the generally low $H_2O$ densities in these regions. At $L_S = 120^\circ$–$160^\circ$ $H_2O$ abundances increase across all latitudes, shifting the abundances upwards by ~5 km in altitude. This general increase in $H_2O$ at high altitudes results in a corresponding ozone decrease in $O_3$, showing an anti-correlation with the $H_2O$ distribution across latitudes and above 30 km. With the reappearance of the high-latitude, high altitude $O_3$ enhancement ~40 km, no corresponding $H_2O$ peak is seen below the $O_3$ enhancement, as observed during the vernal equinox.

Figure 11 shows a sample comparison of several coincident NOMAD $O_3$ and $H_2O$ profiles (top row = NOMAD $H_2O$ observations, bottom row = assimilation model profiles). These profiles cover the time period $L_S = 213.3^\circ$ (MY34) to $L_S = 165.8^\circ$ (MY35), at high-latitudes (up to 83.2°S). In MY34 through perihelion, the relative abundance of both $O_3$ and $H_2O$ increases coherently toward the surface with peak values of both species occurring at the lowest altitudes, consistent with the build-up of $O_3$ at lower altitudes attributed to the reduced $H_2O$ photolysis rates arising from the overlying CO$_2$ atmosphere and the increase in efficiency of the three-body ozone forming reaction at lower altitudes (higher pressures; Montmessin & Lefèvre, 2013). In MY35 through aphelion, $H_2O$ profiles show an inflection around 20–30 km, resulting in maximum $H_2O$ abundances occurring at higher altitudes than during perihelion. $O_3$ profiles show a sharp drop at the point of the $H_2O$ maxima. While the photochemical anti-correlation of $O_3$ and $H_2O$ remains valid throughout the Martian year, there is a clear difference in the shape of the $H_2O$ profiles between the warmer, wetter perihelion season and the colder, drier aphelion season. The top row in Figure 11 shows the $O_3$ and $H_2O$ vertical profile shapes at high-latitudes in the southern summer, whereas the bottom row demonstrates the equivalent profiles for similar latitudes located in the southern polar vortex. For the latter, the higher altitude $O_3$ enhancements are created through the descent of O atoms via Hadley circulation (Montmessin & Lefèvre, 2013), altering the shape of the profile with respect to the southern summer period. Direct NOMAD $O_3$-$H_2O$ comparisons in MY35 should confirm these observations in detail. Furthermore, assimilating retrieved NOMAD $H_2O$ profiles are expected to lead to much better $O_3$ absolute abundance simulations, which often appear lower than NOMAD measurements (Figure 8).
Figure 10. Latitudinal distribution comparison of (left column pair) NOMAD $O_3$ retrievals and (right column) NOMAD $H_2O$ retrievals (Aoki et al., 2019), in $L_s$ intervals of 40° during MY34, and (right column pair) as left pair but for MY35, and with assimilated model output for $H_2O$. Retrievals are binned every 5° latitude and 1 km altitude for clarity. The same color scale is used for both species, with different absolute scales.
4.4. Night-to-Day and Day-to-Night Terminator Perspectives

A key feature of the non-Sun-synchronous orbit of TGO is that it allows NOMAD to observe different local times of day on Mars. This benefit is key to NOMAD nadir observations, which enable total column measurements at varying local times across the planet. Whilst occultation observations are also made at differing local times as a function of latitude and season, by definition an occultation observation always occurs at local sunrise/sunset, at the terminator. Interpreting fine-scale diurnal variations in NOMAD occultation data should be done with care, since occultation measurements will always be performed with the observation tangent point at the day/night boundary, and do not represent the diurnal mean ozone profile.

It is valid, however, to draw conclusions on the difference between sunrise and sunset occultations. From the perspective of viewing geometry, the TGO spacecraft is in eclipse behind Mars prior to a sunrise occultation; the NOMAD-Sun LoS, therefore, transects the atmosphere that is emerging from the nightside conditions of Mars for a sunrise occultation. For a sunset occultation, the opposite is true, where TGO points toward the Sun after passing over the dayside of Mars. For a sunset occultation, the NOMAD-Sun LoS, therefore, transects the atmosphere that has been subjected to the dayside conditions of Mars (high...
temperatures, solar UV). Both cases will consist of observations with part-night/part-day optical paths, with the difference being the exposure period of each optical path to day/night conditions.

Figure 12 shows the vertical profiles for a group of observations between $L_S = 42^\circ–54^\circ$, and latitude range 30$^\circ–50^\circ$S. Whilst local time is plotted, this serves primarily to discriminate between sunrise and sunset occultations, to discern the differences between the nightside and dayside atmospheric “context” of the occultations.

In Figure 12a, the sunrise occultations show an $O_3$ number density at approximately $10^9$ cm$^{-3}$ up to an altitude of $\sim 35$ km. This is in stark contrast to the equivalent sunset occultations, where the equivalent number density only reaches an altitude of $\sim 20$ km. Figure 12b shows this striking distinction, with the profile averages for sunset and sunrise during this period at similar values above $\sim 45$ km and below $\sim 15$ km but showing vastly different abundances between these altitudes. The relative shape of the distribution between 20 and 40 km is similar, but the abundance of $O_3$ is an order of magnitude higher at sunrise than at sunset. Figure 12 thus shows a clear bias toward higher $O_3$ abundances at sunrise for altitudes between 15 and 45 km than for equivalent sunset occultations within the same period. This difference could arise from diurnal partitioning between $O_3$ and $O$, which leads to nighttime conversion of $O$ to $O_3$ when $O_3$ photolysis is inactive, resulting in significant $O_3$ increases above 20–30 km altitudes associated with altitude increasing $O$ densities (e.g., Nair et al., 1994). Day-time photolysis of $O_3$ to $O$ leads to reduced $O_3$ densities above 20–30 km. Consequently, the photolysis-driven diurnal partitioning of $O_3$ and $O$ can plausibly create the observed sunrise and sunset $O_3$ profile variations indicated in Figure 12a. However, the occultation atmospheric paths transect day and night illuminated regions at both terminators, and diurnal $O_3$ variations associated with $O_3$ partitioning are not instantaneous at the terminators. Previous studies have investigated diurnal changes in $O_3$ from a model perspective (e.g., Holmes et al., 2018), however a photochemistry model investigation of the diurnal variation of ozone is beyond the scope of this article. Additionally, it should be noted that the measurements shown in Figure 12 are not made at precisely the same local time and $L_S$, and there is a possibility that the water vapor profile abundance has changed over the short period of $\sim 3$ weeks for the observations in Figure 12, and the difference could therefore be driven by differences in the water profile between sunrise and sunset. Comparison to NOMAD water profile observations is unfortunately not possible at the present time, but once available, detailed photochemical modeling is planned to assess these NOMAD sunrise-sunset distinctions which should yield new constraints on Mars photochemistry.
5. Conclusions

We have analyzed the first Martian year of occultation observations from the NOMAD instrument on ExoMars TGO and presented a new climatology of the vertical distribution of $O_3$ for the period $L_s = 163^\circ$ to $L_s = 320^\circ$ in MY34 to MY35. The vertical, latitudinal, and seasonal ($L_s$) variations of these retrieved NOMAD $O_3$ profiles are compared with coincident NOMAD $H_2O$ profiles obtained during the MY34 global dust storm (Aoki et al., 2019), and to OU data assimilation modeling of $O_3$ and $H_2O$ profiles, in which MCS temperature and dust profiles are assimilated to constrain water vapor saturation conditions.

These NOMAD $O_3$ profile retrievals present the first detailed, global description of Mars $O_3$ vertical distributions as a function of season and altitude, and with accurate sensitivity from below 10 km (depending on variable atmospheric aerosol extinction) to above 50 km (depending on variable $O_3$ abundance) altitudes. Key global-scale variations in Mars atmospheric $O_3$ are well characterized in their vertical extents, including: large $ip_l$ ($40^\circ - 120^\circ$) increases in low-to-mid latitude $O_3$ abundances over 20–50 km altitudes, distinct high-latitude ($> \pm 55^\circ$), high altitude (40–55 km) enhancements in $O_3$ abundances during equinoxial seasons (centered on $L_s = 0^\circ$, $180^\circ$), and very large perihelion $O_3$ decreases over the extended atmosphere. The apHELion increases and perihelion decreases are generally consistent with model predictions of global variation in water vapor saturation altitudes (hygropause) as driven by the elliptical Mars orbit and accentuated by perihelion dust storm heating of the extended atmosphere. During perihelion, the relative vertical abundances of $O_3$ and $H_2O$ increases below $\sim 40$ km toward the surface, with both species having peak abundances at similar altitudes close to the surface; during aphelion, the same general trend applies, but with well-defined decreases in $O_3$ abundances present in the profiles at altitudes of 20–30 km, coincident with local peaks in $H_2O$. Between $L_s = 40^\circ$ and $50^\circ$ in MY35, the difference in $O_3$ abundance between 20 and 40 km was found to be on average an order of magnitude higher at the sunrise terminator, compared to the sunset terminator.

In general, modeled $O_3$ densities are either in agreement with or lower than NOMAD observed $O_3$ densities, most likely due to $H_2O$ abundance conditions being inadequately constrained in the model without significant tuning. An in-progress objective of assimilating NOMAD water profile measurements should provide the most diagnostic photochemical analysis of Martian $O_3$ necessary to accurately represent the formation and distribution of $O_3$ on Mars.

Data Availability Statement

Public access to the derived ozone retrieval data data set and the modeling and assimilation data that support the plots within this article are available through the following data citation reference: Patel (2021).

References


