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Calibration of Herschel SPIRE FTS observations at different spectral resolutions

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ABSTRACT

The SPIRE Fourier Transform Spectrometer on-board the Herschel Space Observatory had two standard spectral resolution modes for science observations: high resolution (HR) and low resolution (LR), which could also be performed in sequence (H+LR). A comparison of the HR and LR resolution spectra taken in this sequential mode revealed a systematic discrepancy in the continuum level. Analysing the data at different stages during standard pipeline processing demonstrates that the telescope and instrument emission affect HR and H+LR observations in a systematically different way. The origin of this difference is found to lie in the variation of both the telescope and instrument response functions, while it is triggered by fast variation of the instrument temperatures. As it is not possible to trace the evolution of the response functions using housekeeping data from the instrument subsystems, the calibration cannot be corrected analytically. Therefore, an empirical correction for LR spectra has been developed, which removes the systematic noise introduced by the variation of the response functions.

Key words: instrumentation: spectrographs – methods: data analysis – space vehicles: instruments – techniques: spectroscopic.

1 INTRODUCTION

The Spectral and Photometric REceiver (SPIRE; Griffin et al. 2010) was one of three focal plane instruments on-board the ESA Herschel Space Observatory (Herschel; Pilbratt et al. 2010). The instrument consisted of an imaging photometric camera and an imaging Fourier Transform Spectrometer (FTS). Both sub-instruments used bolometric detectors operating at ~300 mK (Turner et al. 2001) with feedhorn focal plane optics giving sparse spatial sampling over an extended field of view (Dohlen et al. 2000). The FTS had two broad-band intensity beam splitters in a Mach–Zehnder configuration (Ade, Hamilton & Naylor 1999; Swinyard et al. 2003) and two bolometer arrays with partially overlapping bands: the SPIRE Long Wavelength spectrometer array (SLW; 447-1018 GHz) and SPIRE Short Wavelength spectrometer array (SSW; 944-1568 GHz), with 19 and 37 detectors. Detailed information about the standard procedure for the calibration of FTS data and the calibration accuracy can be found in Swinyard et al. (2010), Swinyard et al. (2014), and Hopwood et al. (2015). All products in this paper were produced by the standard data processing pipeline (Fulton et al. 2016) in the Herschel Interactive Processing Environment (HIPE; Ott 2010) version 13 and calibration tree spire_cal_13_2. The data were downloaded from the Herschel Science Archive (Leon et al. 2009).

The FTS operated by splitting incoming radiation into two beams. An optical path difference (OPD) was introduced between the beams by scanning an internal mirror, so that when recombined, an interference pattern (known as an interferogram) is formed. The inverse Fourier transform of the interferogram gives the spectrum,
The difference between the HR and the LR spectra shown in Fig. 1. The HR spectra from detectors SLWC3 and SSWD4, respectively; in black and red, the HR ones.

4, 5, and 6, where the most plausible explanation is discussed in detail. Since there is no possible analytical solution to the problem, an empirical correction has been developed (Section 7). Its effectiveness is discussed in Section 8 through the comparison of the calibration uncertainty before and after the correction. The main findings of this work are summarized in Section 9.

2 COMPARISON OF UNCALIBRATED SPECTRA

The total uncalibrated signal is the sum of three components: the instrument emission, the telescope emission, and the source emission. The first two contributions can be expressed in terms of voltage density as

\[ V_{\text{inst}}(\nu) = M_{\text{inst}}(\nu)R_{\text{inst}}(\nu) \quad [\text{V GHz}^{-1}] \]  \hspace{1cm} (1)

and

\[ V_{\text{tel}}(\nu) = M_{\text{tel}}(\nu)R_{\text{tel}}(\nu) \quad [\text{V GHz}^{-1}] \]  \hspace{1cm} (2)

where \( M_{\text{inst}}(\nu) \) is the instrument model based on a blackbody emitter at temperature values provided by the instrument thermometry, \( R_{\text{inst}}(\nu) \) is the instrument relative spectral response function (RSRF), \( M_{\text{tel}}(\nu) \) is the telescope model, which is based on two blackbody emitters (at the primary and secondary mirror temperatures) and the telescope mirror emissivity, and \( R_{\text{tel}}(\nu) \) is the telescope RSRF.

The instrument and telescope contributions to the signal are removed at two key steps of the calibration pipeline (as described in Fulton et al. 2016). First, the instrument emission of \( M_{\text{inst}}(\nu)R_{\text{inst}}(\nu) \) is subtracted. Secondly, the remaining signal is divided by \( R_{\text{tel}}(\nu) \) before \( M_{\text{tel}}(\nu) \) is subtracted.

\( R_{\text{tel}}(\nu) \) and \( R_{\text{inst}}(\nu) \) depend on resolution mode. The estimation of the response functions is performed through a procedure that aims to minimize the residual noise in observations of the SPIRE dark sky field. This is a dark region of sky centred on RA:17h40m12s and Dec:+69d00m00s (J2000) and selected on the grounds that the region has low cirrus, it is visible at all times and contains no SPIRE-bright sources. Before \textit{HiPE} 9, LR calibration was extrapolated from the LR portion of CR and HR dark sky observations. This method, however, was not optimal for LR observations, as large systematic residuals were left in the calibrated spectra. For the later versions of \textit{HiPE}, the calibration of LR spectra is carried out using a new set of response functions, based on LR dark sky observations. The

We present an analysis of the discrepancies seen in the spectra as a function of operating mode. Figs 1 and 2 show that the problem primarily concerns spectra from SLW, so we focus on the centre SLW detector (SLWC3). The basic characteristics of the signal distortion are identified in Sections 2 and 3. Several hypotheses concerning the nature of the problem are presented in Sections
HR RSRFs, instead, are calculated by considering HR dark sky observations.

The existence of a difference between the response functions for HR and LR data implies that the amount of signal removed at these stages of the data calibration changes with the resolution mode, raising the question as to whether the discrepancies shown in Fig. 2 are intrinsic to raw data, or are introduced during processing.

The answer comes from the comparison between uncalibrated H+LR(H) and H+LR(L) spectra. Their difference before the telescope and instrument correction is marginal. It also appears that HR uncalibrated spectra are consistent with both H+LR(H) and H+LR(L) spectra, while the LR ones are systematically different from all the others (see upper panel of Fig. 3). Therefore, the root problem is not with H+LR spectra, rather it is the systematic difference between LR spectra and the ones obtained at any other resolution mode, from now on indicated as HR/H+LR. Looking at the difference between LR and H+LR(L) (black line in Fig. 3), the double bump already evident in Fig. 2 is clearly discernible. Significant discrepancies between LR and H+LR data concern not only the spectra, but also the interferograms (see lower panel of Fig. 3), which are themselves constructed from the raw data, known as the Spectrometer Detector Timelines (see Fulton et al. 2016), suggesting that the differences observed are present in the raw signal.

It can be inferred that, while the $R_{\text{inst}}(v)$ function implemented in the pipeline is efficient in the calibration of LR spectra, it introduces a systematic bias in the calibration of the H+LR(L) spectra. It should be noted, however, that such bias does not exclusively affect H+LR(L), as its characteristic shape can also be recognized in LR spectra, albeit with lower amplitudes.

3 ANALYSIS OF DARK SKY OBSERVATIONS

A fundamental issue that must be addressed when investigating the origin of the discrepancy between HR/H+LR and LR spectra: is the difference between these spectra a constant or does it depend on some parameters, such as time, instrument temperature, or telescope temperature? The differences calculated from several pairs of observations (as it will be shown in Section 3.2) are not identical to each other. The observed variations can be due to random noise, or to other causes which have nothing to do with the resolution modes in which the observations were taken (see Swinyard et al. 2014). Besides these contributions, there are variations specifically relating to the resolution mode of the data, and the root cause of this discrepancy has been investigated.

In the following, we will calculate the differences between spectra for specific pairs of observations; these differences will be referred to as $\delta_{\text{HR-LR}}(\text{ObsID}_{\text{HR,HR-LR}})$, where ObsID stands for the identification numbers of the observations. The systematic difference between HR/H+LR and LR spectra, defined as the ideal difference between spectra observed simultaneously and free from noise and resolution-independent sources of radiation, will be referred to as $\delta_{\text{HR-LR}}$.

The difference $\delta_{\text{HR-LR}}$ was estimated for a set of dark sky observations performed at different resolution modes. Given the low number of H+LR observations and the similarity between H+LR and HR data, we focused on the comparison between HR and LR dark sky observations. Note that HR observations can be calibrated as if they were taken in the LR mode, by truncating HR interferograms at the same OPD for LR. While true in principle, this may not be valid in practice, as the heat dissipated during the HR scan (which goes as OPD$^2$) is much greater and results in a different instrument environment. Most of the LR dark sky observations acquired by the SPIRE FTS are concentrated within the time span between the Operational Days (OD) 1079 and 1433. Before OD 1079, dark sky observations were generally taken in the CR mode, which is analogous to the HR mode (see the SPIRE Handbook 2016). For the present analysis, we focused on a sample of dark sky observations between OD 1079 and OD 1325. The identification numbers of these observations are listed in Table 1.

From the comparison of the uncalibrated spectra (see Fig. 4), it is easy to recognize the double bump that characterizes the discrepancy between HR and LR data. However, it is also evident that there is a large spread in the plotted curves caused, as previously mentioned, by the significant differences in the instrument emission

<table>
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</table>

Figure 3. Upper panel: an example of the differences between the uncalibrated spectra obtained at different resolution modes. Lower panel: the differences seen in the interferograms. Uncertainties on the plotted quantities are below 1 per cent.

Table 1. Dark sky observations used for the comparison of spectra at different resolution modes. Column 1 reports the operational day. Columns 2 and 3 give the identification number of the LR and HR observations, respectively.
between LR and HR/H+LR spectra can be regarded as constant in time. In the standard calibration procedure (see Section 2), the total signal is regarded as the sum of three components, one depending on the target of the observation, the second on telescope temperature, and the third one on instrument temperature. Since the three components are variable, it is hard to understand how any of these can contribute a spurious, resolution-dependent signal that is approximately constant in time. To check whether the problem can be resolved with an empirical modification to the calibration, we consider a third noise component that only depends on resolution modes and is added to the noise contributed by the telescope and the instrument.

3.1 Three-parameter model for the calibration of the data

The three-parameter calibration model assumes that the measured voltage density $V_i$ for a dark sky observation $i$ can be expressed as

$$V_i(\nu) = M_{Te}(\nu)R_{Te}(\nu) + M_{Inst}(\nu)R_{Inst}(\nu) + f(\nu) \quad [V \, GHz^{-1}]$$

(3)

where $R_{Te}(\nu)$ and $R_{Inst}(\nu)$ are new estimates of the telescope and instrument response functions, and $f(\nu)$ is a resolution-dependent calibration parameter. For each of the observations in Table 1, the variables $V_i(\nu)$, $M_{Te}(\nu)$, and $M_{Inst}(\nu)$ can be determined. Given the 19 LR (17 HR) observations at our disposal, for each frequency a system of 19 (17) equations with three unknowns can be built to calculate the best-fitting parameters for LR (HR) data. Since the systems are overdetermined, a least-square fitting algorithm was used to simultaneously estimate $R_{Te}(\nu)$, $R_{Inst}(\nu)$, and $f(\nu)$. The results are plotted in Fig. 6.

The differences between the best-fitting parameters for LR and HR data are not limited to $f(\nu)$. Similarly to the standard pipeline, we find that the RSRFs at different resolutions are inconsistent. In Fig. 7, the differences between the contributions at different resolution modes provided by each of the three components on the right-hand side of equation (3) are shown.

This result has two important consequences. (i) We regarded $\delta_{HR-LR}$ as a constant, but this assumption is correct only to a first approximation; the discrepancies in the response functions at different resolutions imply that the amplitude of the discrepancy must depend on the instrument and telescope temperatures, which vary as

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1 The modest fluctuations among the $\delta_{HR-LR}$ curves obtained at different ODs indicates that, to a first approximation, the discrepancy between LR and the quasi-simultaneous HR observation. On the one hand, in order to isolate the systematic part of the discrepancy from possible contamination introduced by the instrument and telescope corrections, it would seem reasonable to look at uncorrected spectra. On the other hand, the differences in the instrument emission between consecutive observations are such that, without instrument correction, $\delta_{HR-LR}$ cannot be properly evaluated. The similarity between the spectra obtained after the instrument correction (see Fig. 5) suggests that this is the optimal stage of calibration to estimate $\delta_{HR-LR}$.

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Footnotes:

1. The important contribution of the instrument emission to the difference between LR and HR spectra might appear in contradiction with the negligible difference seen between H+LR(H) and H+LR(L), as reported in the previous section. The two cases, however, are different, because of the order in which the observations are taken. LR observations are almost always taken before HR ones, while, in H+LR observations, the LR scans always follow the HR ones. This point will be further considered later on.

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Figure 4. Difference between quasi-simultaneous LR and HR spectra. The relatively large spread between the lines is caused by the uncorrected instrument emission. Uncertainties on the plotted quantities are below 1 per cent.

Figure 5. Difference between quasi-simultaneous LR and HR spectra, after applying the instrument correction discussed in Section 2. Uncertainties on the plotted quantities are below 1 per cent.

Figure 6. The best-fitting RSRFs (left-hand panel) and the $f(\nu)$ parameter (right-hand panel) estimated from the 19 LR and 17 HR dark sky observations in Table 1. The HR spectra utilized for calculating the parameters have been produced in LR mode using the specific option in the data reduction pipeline.
SPIRE spectra at different resolutions

Figure 7. Average differences between the correction curves for LR and HR data. The instrument (black line) and telescope (red line) correction curves are calculated by multiplying the RSRF plotted in Fig. 6 by the instrument and telescope models.

Figure 8. The temporal evolution of $\delta_{\text{HR-LR}}(\text{ObsID}_{\text{HR}}, \text{ObsID}_{\text{LR}})$ at three different frequencies: 524 GHz (black dots), 712 GHz (red squares), and 884 GHz (green dots).

Table 2. Dark sky observations carried out in OD 1291 (Col. 1), with their resolution modes (Col. 2) and number of repetitions (Col. 3).

<table>
<thead>
<tr>
<th>Obsid</th>
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<th>Reps</th>
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<tr>
<td>1342256091</td>
<td>HR</td>
<td>70</td>
</tr>
</tbody>
</table>

Figure 9. Upper panel: the variation of the instrument temperature, measured by one of the dedicated sensors, during the observations in OD 1291. Lower panel: the spectra of the observations from 1342256085 to 1342256089, after subtracting the spectrum of observation 1342256091. Uncertainties on the plotted values are below 1 per cent.

3.2 Dark sky observations in OD 1291

The different behaviour of the spectral response functions at LR and HR seems to suggest a different sensitivity of the mechanics or the optics to the resolution mode, which is an unlikely scenario. The origin of the problem can be further investigated through the analysis of a set of dark sky observations carried out in OD 1291 and performed in the LR, HR, and H+LR resolution modes (see Table 2). The difference between instrument-corrected spectra of all the observations from 1342256085 to 1342256089 and the one of obsid 1342256091 is plotted in the lower panel of Fig. 9 from which several conclusions can be drawn.
(i) Depending on the temporal sequence in which they are observed, HR and LR spectra are sometimes almost identical (see the green and the blue line). This is consistent with the negligible differences seen between uncalibrated H+LR(H) and H+LR(L) (see Fig. 3).

(ii) The differences among HR spectra can be similar in shape and comparable in amplitude to $\delta_{HR,LR}$ (see e.g. the green line, showing the difference between obsids 1342256087 and obsid 1342256091).

(iii) LR spectra show important variations from one observation to another.

All these arguments point towards excluding the resolution mode as the direct cause of the problem. The differences with respect to the 1342256091 spectrum are all characterized by the typical double bump of $\delta_{HR,LR}$, but the amplitude of the effect varies considerably. Inspecting the housekeeping products for this set of observations reveals most of the parameters show only mild variation with time – above all the telescope. However, fast variability can be detected in the instrument temperature. Looking at its behaviour during the data acquisition (upper panel of Fig. 9), it is hard to find an unequivocal pattern. (i) Most of the spectra seem to gradually tend towards a decreasing difference with respect to 1342256091, which could suggest a time dependence of the amplitude of the effect; however, the transition from obsid 1342256085 (black line) to 1342256086 (red line) seems to go in the opposite direction. (ii) Obsid 1342256086 (red line) is characterized by a positive temperature gradient and a larger departure from obsid 1342256091, while the LR spectra in obsids 1342256088 (blue line) and 1342256089 (brown line) by negative temperature gradients and proportionally smaller discrepancies from obsid 1342256091. A correlation between gradients and amplitude of the effect, however, could apply only to LR data, because HR observations seem to contradict it. (iii) The most evident distinction between observation 1342256091 (violet line) and all others is the number of repetitions and the duration of the data acquisition. While the 70 repetitions that comprise 1342256091 allow the instrument temperature to reach an approximately constant value, observations from 1342256086 to 1342256089 show strong temperature gradients; this does not apply to obsid 1342256085 (black line), which is short, but does not show important instrument temperature variations.

In summary, the discrepancies among the spectra are not connected to the instrument temperature in a straightforward way, while the number of repetitions of the observations seems to have a significant influence on the measured flux densities.

4 A LINK TO THE TELESCOPE MODEL CORRECTION?

The hypothesis of a relationship between number of repetitions of an observation and the amplitude of the discrepancy with respect to an average HR spectrum is particularly interesting when confronted with the result of an independent study about the calibration of the SPIRE FTS data, namely the telescope model correction (see Hopwood et al. 2014). This study, which takes into account HR spectra, demonstrates the significant improvement of the calibration results after multiplying the telescope model by a time-dependent factor; it is hypothesized that the correction is required because of an extra emission caused by the buildup of dust on the surface of the telescope. It is also shown that the correction factor changes according to the number of repetitions of the observation: observations with $<20$ repetitions require a higher correction factor than those with $>20$ repetitions (see figs 5 and 6 in Hopwood et al. 2014). The authors hypothesize that this difference is caused by a higher-than-average instrument temperature (most of the short dark sky observations were taken at the end of an FTS observing cycle, when the temperatures are generally higher). However, provided that the duration of an observation correlates with the number of repetitions, it might also be that the dependence of the correction factor on the number of repetitions is an indirect consequence of a dependence on the duration of the observation. Since LR observations are systematically shorter than the HR ones, it follows that the discrepancy between HR and LR data may be one aspect of a more general problem that has to do with the duration of an observation, rather than its resolution mode.

From the discussion above, one might wonder if the origin of $\delta_{HR,LR}$ is the telescope model itself. The multiplication of the original model by a correction factor can have similar effects as a change of the telescope RSRF, which would be compatible with the results reported in Section 3.1, but only up to a point. Hypothesizing the existence of a correction factor that varies with either the duration or the number of repetitions of an observation, the difference between two spectra could be expressed, as a function of frequency, as $(E_{\text{corr}} - E_{\text{inst}})M_{\text{tel}}(\nu)R_{\text{tel}}(\nu)$, where $E_{\text{corr}}$ and $E_{\text{inst}}$ would be the correction functions to apply for the given observations. This difference would be proportional to $M_{\text{tel}}(\nu)R_{\text{tel}}(\nu)$, whose shape is known, and is incompatible with the double bump characterizing $\delta_{HR,LR}$. Inverting the problem, it could instead be hypothesized that the detected difference between the correction factors is the consequence of observation-duration-dependent RSRFs.

5 RSRF CALCULATION FROM SCANS OF THE SAME OBSERVATION

The differences among the quasi-simultaneous dark sky observations in OD 1291 demonstrate that variation of the RSRFs can occur on very short time-scales -- of the order of minutes. A way to investigate such short-term variations is to analyse on a scan-by-scan basis, rather than using the standard pipeline products that (for each detector) are averaged over all scans per observation. In the following, we will focus on the flux densities at a frequency of 524 GHz, around which the discrepancy between LR and HR/H+LR data is most pronounced.

Changes in telescope emission during an observation is minimal and therefore subtraction of a single telescope model per observation is not responsible for any significant deviation of scans from the expected behaviour. In Fig. 10, the telescope-corrected signal at 524 GHz from observations 1342256089 (H+LR) and 1342256091 (HR) is plotted versus the instrument model $M_{\text{inst}}$. According to equation (3), the data points should fall along a straight line, whose slope should provide the best instrument response function for the set of data, $R_{\text{inst}}$(524 GHz). For comparison, the plot also reports the telescope-corrected signal at 524 GHz (black dots) and fitted slope (black line) for the LR dark sky (henceforth, LR$_{\text{dark}}$) observations in Table 1.

The HR scans of obsid 1342256089 (orange squares) are the first in order of time and during these 10 scans, the instrument temperature increases almost monotonically. The data points for the first three scans are consistent with the average behaviour of LR observations. As the temperature further increases, the decrease in $V_{\nu}(\nu) - M_{\text{tel}}R_{\text{tel}}$ becomes much steeper than expected. This indicates a change of $R_{\text{inst}}$(524 GHz) (represented by the slope of
The evidence presented so far indicates that slow changes of the instrument temperature cause variation of the voltage density $V_i (v)$ (which can be described with the standard formula $M_{\text{Inst}} (v) R_{\text{Inst}} (\nu, LR) (v)$). In contrast, the execution of scans in the HR mode initially causes a fast increase of the instrument temperature, which produces a change in the instrument response function and leads to an augmented deviation of the measured flux density when compared to the standard LR observation. After a number of HR scans, the instrument temperature stabilizes before gradually decreasing. The instrument response function tends to a value slightly steeper than $R_{\text{Inst}, LR} (v)$, approximately preserving the flux density difference introduced by the previous scans. Since the absolute value of the response function for these scans is marginally higher than $R_{\text{Inst}, LR} (v)$, the instrument emission will eventually return to a level that is compatible with $M_{\text{Inst}} (v) R_{\text{Inst}, LR} (v)$, although this happens on time-scales much longer than the duration of an average observation. As in a hysteresis cycle, the instrument contribution to $V_i (v)$ depends on the evolution of the instrument temperature, i.e. the way it changed in previous observations.

### 5.1 Extending the scan-by-scan analysis to a large sample of HR observations

To verify that fast instrument temperature variations are responsible for the problem under discussion, a scan-by-scan analysis was also applied to a sample of 21 HR dark sky observations, from operational day 466 to 1389. They range from 30 to 160 repetitions in length, where one repetition corresponds to two scans. Since the strongest temperature variations concern the first 20–30 scans of each observation, our analysis is limited to the first 30 and the last 10 scans only. The results are shown in Figs 11 and 12, where the telescope-corrected signal is plotted versus the instrument model. The green and the orange dotted lines show the typical relationship between signal and instrument model for LR and HR spectra, respectively. Two distinct kinds of behavior are evident. 

(i) For the 11 cases shown in Fig. 11, the signal evolves with the changing temperature according to a clearly recognizable pattern. To start with, the signal follows the typical evolution expected for LR spectra (green line). After a number of scans, the signal rapidly moves towards a lower state, which indicates that the contribution of the instrument emission to the signal is higher than expected for the measured temperature. When the instrument temperature starts to decrease, the signal is generally moving along the orange line.

Table 3 summarizes the main findings. For all scans under consideration (Col. 1), we report the resolution mode (Col. 2), the behaviour of $M_{\text{Inst}}$ during the scans (Col. 3), and the inferred $R'_{\text{Inst}}$ (Col. 4). These results indicate a link between $R'_{\text{Inst}}$ (which has a direct influence on the measured flux density at 524 GHz) and the variation of $M_{\text{Inst}}$. The resolution mode does not cause systematic variations of the instrument response function.

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<td>1342256091</td>
<td>HR</td>
<td>increasing</td>
<td>$(-3.2 \pm 0.5)e+12$</td>
</tr>
<tr>
<td>1342256091</td>
<td>HR</td>
<td>decreasing</td>
<td>$(-2.00 \pm 0.15)e+12$</td>
</tr>
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</table>
The evolution of the telescope-corrected signal at 524 GHz for samples of scans of HR observations. The signal follows a kind of U-shaped pattern (a model of this pattern is shown as a magenta line in the top-right box), moving clockwise from a typical LR behaviour (green dotted line) to a typical HR one (orange dotted line).

Figure 11. Telescope-corrected signal at 524 GHz plotted versus the instrument contribution to the signal follows the typical LR line, temperature variations are slow (as during or after LR scans), the trigger by fast variations of the instrument temperature. When between observations performed in different resolution modes are..
combinations of housekeeping parameters. The situation is complicated further by the fact that the response functions are empirically calculated in a way that tends to minimize the residual noise, which means that part of the spurious signal may be removed by the standard calibration.

The correction of the spectra can be achieved empirically, by exploiting the peculiar shape of the systematic noise introduced by the variation of the response functions (see Figs 3 and 13). In OPD space (lower panel of Fig. 3), the discrepancy between spectra at different resolutions corresponds to the narrow and heavily modulated part of the interferogram that occurs around zero path difference. Developing a correction schema for the interferogram domain is very challenging since each point in the interferogram represents an integral over frequency. Also, the amplitude of the discrepancy cannot be precisely estimated without a preliminary subtraction of both the telescope and instrument contributions to the total signal, which can only be achieved for those LR spectra that can be compared with quasi-simultaneous HR spectra. In the spectral domain, instead, the discrepancy is characterized by a double bump whose shape is approximately constant, and can be precisely modelled.

An algorithm for the a posteriori correction of LR spectra, both from LR and H+LR(L) observations, has been developed based on the cross-correlation of an archetype of $\delta_{HR-LR}$ with the spectrum to correct. The $\delta_{HR-LR}$ archetype (from now on, $\sigma_{HR-LR}(v)$) has been calculated by selecting a set of quasi-simultaneous HR and LR spectra, whose differences are clearly affected by the LR calibration problem. These differences have been scaled in order to have the same standard deviation, and then averaged to provide a robust model.

Given a calibrated spectrum $S(v)$, with frequency-averaged value $\bar{S}$, the corrected spectrum $S'_{\nu}$ can be expressed as

$$S'(v) = S(v) - \sum \left( \frac{(S(v) - \bar{S}) \cdot \sigma_{HR-LR}(v)}{\sigma_{\nu}^2} \right) \cdot \delta_{HR-LR}(v)$$

where $\sigma_{\nu}$ is the standard deviation of $\sigma_{HR-LR}(v)$. This is equivalent to calculating the amplitude of the spurious noise for $S(v)$ and removing it, assuming its shape to be $\sigma_{HR-LR}(v)$.

When a calibrated spectrum has a strong continuum, the amplitude of the spurious noise may be overestimated. This overestimate can be avoided by fitting and subtracting a second-order polynomial to remove the continuum before calculating the noise amplitude. Such a subtraction is only necessary for sources where the continuum level at 1000 GHz is found to be greater than 0.5 Jy.

The LR correction developed is scaled depending on the amplitude of the double bump present. This means that if the double bumps are negligible, then the corrections applied are negligible too, and the data are essentially unchanged. It is only applied to SLW detectors that are the only ones affected by the problem, and efficiently corrects both LR and H+LR(L) observations.

The correction was developed for point-source calibrated spectra, using sparse-mode observations. The standard FTS pipeline also provides extended-source calibrated data, which are either spectra for sparse-mode observations, or spectra and spectral cubes for intermediate and fully sampled mapping observations (see Fulton et al. 2016 for full details of the FTS data reduction pipeline). The LR correction is not applied directly to the extended-source calibrated data, but is propagated by reversing the point-source conversion once the point-source calibrated spectra have been corrected. For LR mapping observations, the conversion to point-source calibrated is an additional step for the SLW array, purely so the double-bump correction can be applied. Fig. 14 shows an example of the extended-source calibrated spectra from the centre detectors for observation 1342243638, before and after correction for the double bump. It should be noted that there is no point-source conversion factor for the vignetted detectors. These detectors are removed before point-source calibration and therefore remain uncorrected in the final extended-source-calibrated products. For sparse observations, the uncorrected spectra from these detectors should only be used with caution. For mapping observations, the spectra for all SLW detectors are present in the pre-processed cube that collates the spectra ready for gridding into a spectral cube. However, the spectra from the vignetted detectors are not included in those used to create the standard pipeline SLW cube products.

8 LR CALIBRATION UNCERTAINTY

Due to the pronounced systematic noise, the derived correction removes from the SLW spectra of LR observations; here, we report the LR calibration uncertainties, before and after the correction has been applied. LR sensitivity is assessed following the method in Hopwood et al. (2015). Briefly, the spectral noise within 50 GHz frequency bins is measured for 34 LR observations of the SPIRE
Figure 15. LR sensitivity for point-source calibrated SLWC3 spectra on the left and extended-source calibrated SLWC3 spectra on the right. Black curves show the sensitivity before the application of the LR correction, which is compared to after LR correction in green. The other centre detector (SSWD4) is not shown as the LR correction does not apply.

Figure 16. LR continuum offset for point-source calibrated SLWC3 spectra on the left and extended-source calibrated SLWC3 spectra on the right. Black curves show the continuum offset before the application of the LR correction, which is compared to after LR correction in green. The other centre detector (SSWD4) is not shown as the LR correction does not apply.

dark sky field. The noise taken for each dark sky observation is scaled to 1σ in 1 h (i.e. the sensitivity) and then the median is found over all the observations for each frequency bin. There is negligible change in the sensitivity after the LR correction has been applied to the centre SLW detector (see Fig. 15), which is expected. However, the correction of large-scale distortion in the continuum does significantly affect the uncertainty associated with continuum measurements. Fig. 16 shows this ‘continuum offset’ (1σ additive uncertainty), before and after the LR correction of dark sky. Again, the method detailed in Hopwood et al. (2015) is followed. The same set of LR dark sky observations is used as for the sensitivity, but in their unaveraged form, i.e. as individual scans. The spectra are smoothed to remove small-scale noise and then the standard deviation over all the scans of the 34 dark sky is taken to provide the continuum offset. Before the correction is applied, the effect of the double bump is clearly evident, and when compared to the HR continuum offset (fig. 32 in Hopwood et al. 2015), the shape of the curve is notably different. This indicates that although HR data can suffer from the same ‘bumpiness’ as LR data (but generally as an inverted double bump), overall this is not a significant issue for HR, as the HR continuum offset is a relatively smooth curve without any correction for this systematic effect. After the bumpiness has been removed from the LR dark sky, the resulting continuum offset shows an improved and more similar form as to that seen for HR. Fig. A1 (in the appendix) shows the continuum offset for all unvignetted SLW detectors, indicating there is a range in how the curves differ when comparing the continuum offset found before and after the LR correction is applied. For some detectors, there is almost no change, while for others there is a pronounced improvement. These figures illustrate that the double bump affects different detectors to differing degrees and the correction is working harder where needed. Similar results are found for the extended-calibrated sensitivity and continuum offset, although the improvement for SLW is somewhat easier to see for the point-source calibrated data, as the difference in beam size across the frequency array has been accounted for. The average results for both FTS calibrations schemes, and for all off-axis detectors, can be found in Table 4. The sensitivity is relatively flat, so the mean sum is used for the average. Similarly to the case of SLWC3 (see Fig. 15), the correction of the double bump in off-axis detectors leads to marginal differences in the sensitivity. The continuum offset is also relatively flat, but at the centre of the SLW and SSW bands, and only once the LR correction has been applied. However, the high systematic noise at the edge of the bands causes the offset to rise sharply even for the corrected data. Therefore, for the average values, the median is taken over 500–900 GHz for SLW and over 1000–1500 GHz for SSW. Note that as there is no LR correction for the vignetted detectors (as discussed in Section 7), only the uncorrected results for these detectors are presented. However, all detectors are included in Table 4. The extended-source calibrated data used do include the correction applied for the feed-horn coupling efficiency, as described in Valchonov et al. (in preparation), and is a comparison of HIPE version 14.0 and HIPE version 14.1. Whereas, for the point-source-calibrated data, the LR correction was already in place for HIPE version 14.0 and so this is compared to HIPE version 13.0 reduced data.

9 SUMMARY

In this paper, we present a thorough analysis of the systematic discrepancy between spectra obtained at different resolution modes. The discrepancy, which was first detected in the fully calibrated H+LR(H) and H+LR(L) spectra, originates in a systematic difference between uncalibrated LR and HR/LR spectra, which the standard calibration pipeline partially corrects by assuming different response functions for LR and HR data.

We tested the hypothesis that the discrepancy, δLR−HR, is constant in time, and therefore whether it can be corrected by adding a constant parameter to the standard two-parameter calibration pipeline. The test showed that, in first approximation, the amplitude of the discrepancy can be regarded as constant; however, a minor part of the discrepancy is certainly affected by the (time-dependent) telescope and instrument temperatures. Given the shape of δLR−HR (which shows an excess of signal around 550 and 900 GHz, and a deficiency of signal around 700 GHz), a correction of the problem cannot be achieved by modifying the telescope and/or the instrument model; a modification of the response functions would be needed.

The analysis of the variations in dark sky observations spectra performed in OD 1291, even on a scan-by-scan level, seems to indicate that the amplitude of the discrepancy is not directly related to the instrument temperature or its increasing/decreasing trend. Rather, the amplitude appears to be affected by the history of the temperature variations before the start of the observation. The extension of the scan-by-scan analysis to a sample of 21 HR observations strongly supports this conclusion.

We hypothesize that the existence of δLR−HR is related to fast temperature changes in the instrument, which cause temporary variations in both the telescope and the instrument response functions. The sensitivity of the telescope response function to the instrument temperature may also explain the important analogies observed between some characteristics of δHR−LR and the telescope emission, such as the dependence on the number of repetitions (and consequently the duration) of an observation.

Given that the housekeeping parameters do not trace the SLW LR calibration problem in all its aspects, an analytical correction
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Fulton T. R. et al., 2010, Proc. SPIE, 7731, 773134

APPENDIX A: LR CONTINUUM OFFSET
The plots of the continuum offset for all unvignetted SLW detectors are shown in Fig. A1. Note that the correction of the spurious double bump in LR spectra has only been applied to these detectors.

is not possible. Instead, an empirical a posteriori correction has been developed, based on the cross-correlation of a calibrated spectrum with the characteristic double bump the problem causes. The strength of the correction applied is dependent only on the amplitude of the bumps present. With this empirical method, it has been possible to significantly reduce the spectral artefacts that appear in the long wavelength spectrometer channel, for both LR spectra and the LR spectra of H+LR observations.

The analysis of the LR calibration uncertainties shows that applying the correction results in an efficient removal of the double bump. This is true for even the most pronounced cases, while for the rare cases where the double-bump feature is negligible, the spectra remains essentially unchanged.

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**Figure A1.** LR continuum offset: for each detector, the offsets for point-source calibrated data (left) and extended-source calibrated data (right), before (black) and after (green) application of the LR correction are shown.

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