Predominantly Non-Solar Origin of Nitrogen in Lunar Soils

How to cite:

For guidance on citations see FAQs.

© 2016 The Author(s)

https://creativecommons.org/licenses/by/4.0/

Version: Version of Record

Link(s) to article on publisher’s website:
http://dx.doi.org/doi:10.1016/j.gca.2016.08.006

Copyright and Moral Rights for the articles on this site are retained by the individual authors and/or other copyright owners. For more information on Open Research Online’s data policy on reuse of materials please consult the policies page.

oro.open.ac.uk
Predominantly non-solar origin of nitrogen in lunar soils

J. Mortimer a, A.B. Verchovsky a, M. Anand a,b,*

a Planetary and Space Sciences, Department of Physical Sciences, The Open University, Milton Keynes MK7 6AA, UK
b Department of Earth Sciences, The Natural History Museum, London SW7 5BD, UK

Received 26 January 2016; accepted in revised form 4 August 2016; Available online 11 August 2016

Abstract

Simultaneous static-mode mass spectrometric measurements of nitrogen, carbon, helium, neon, and argon, extracted from the same aliquot of sample by high-resolution stepped combustion, have been made for a suite of five lunar soils. Noble gas isotope ratios show that the majority of noble gases are derived from a solar wind source; for example, at peak release temperatures of 500–600 °C, $^{21}\text{Ne}/^{22}\text{Ne} = 0.0313 \pm 0.0007$ to $0.0333 \pm 0.0007$, and $^{20}\text{Ne}/^{22}\text{Ne} = 11.48 \pm 0.05$ to 12.43 $\pm 0.07$, with values at the lowest temperature steps less fractionated during implantation from, and therefore even closer to, solar values ($^{21}\text{Ne}/^{22}\text{Ne}_{\text{SW}} = 0.03361 \pm 0.00018$ and $^{20}\text{Ne}/^{22}\text{Ne}_{\text{SW}} = 14.001 \pm 0.042$ (Pepin et al., 2012)). Despite the co-release of nitrogen and solar wind argon, measured nitrogen isotopic signatures at each temperature step, whilst variable, are significantly more enriched in $^{15}\text{N}$ compared to the measured solar wind nitrogen value from the Genesis mission. Therefore, mixing between a $^{15}\text{N}$-enriched non-solar planetary nitrogen source with solar wind nitrogen is required to explain the measured isotopic values from the stepped combustion analysis of lunar soils. Binary mixing calculations, made under different assumptions about the degree of loss of solar wind $^{36}\text{Ar}$, reveal that the majority (up to 98%) of the nitrogen released is derived from a non-solar source. The range of modelled non-solar end-member nitrogen compositions required to satisfy the measured $\delta^{15}\text{N}$ values varies between samples and temperature steps from $+5\text{‰}$ up to $+300\text{‰}$, or between $+87\text{‰}$ and $+160\text{‰}$ for bulk samples. This range of modelled isotopic compositions for the non-solar source of nitrogen encompasses measured values for several different groups of carbonaceous chondrite, as well as IDPs.

© 2016 The Author(s). Published by Elsevier Ltd. This is an open access article under the CC BY license (http://creativecommons.org/licenses/by/4.0/).

Keywords: Moon; Regoliths; Cosmochemistry; Solar wind; Meteorites

1. INTRODUCTION

For decades, there has been debate about the source(s) of nitrogen in lunar soils; since noble gases in lunar soils are predominantly solar, and nitrogen is both well correlated with these solar components, and likewise, co-located on soil grain surfaces (Wieler et al., 1999), a predominantly solar source for nitrogen was expected. However, stepped heating analyses of lunar regolith have consistently revealed a similar release profile across all Apollo and Luna soils, with low temperature, $^{15}\text{N}$-enriched nitrogen preceding a mid-temperature, $^{15}\text{N}$-depleted component, which is then followed at the highest temperature steps by a second $^{15}\text{N}$-enriched release (e.g. Brilliant et al., 1994; Assonov et al., 2002). This pattern of heavy-light-heavy (‘V-shaped’) nitrogen, with large-scale variations in nitrogen isotopic composition of around 300‰ (Becker and Clayton, 1975; Kerridge, 1975, 1993; Clayton and Thiemens, 1980; Frick et al., 1988; Brilliant et al., 1994; Wieler et al., 1999; Assonov et al., 2002; Füri et al., 2012), coupled with the observation that nitrogen isotopic variations correlate with $^{40}\text{Ar}/^{36}\text{Ar}$ ratios (which are taken as a proxy for soil antiquity, or when in
the past a soil was exposed at the lunar surface) (Füri et al., 2012), led many to suggest that the composition of the solar wind must have changed over time, becoming progressively more enriched in $^{15}$N (Kerridge, 1975; Becker and Clayton, 1975; Clayton and Thiemens, 1980).

In contrast, solar noble gases as measured in lunar soils display very minor variations in comparison to nitrogen isotopic compositions (e.g. Kerridge, 1989, 1993), inconsistent with a secular variation hypothesis for nitrogen. Further, N/$^{36}$Ar ratios measured in lunar soils (Frick et al., 1988; Kerridge, 1993; Wieler et al., 1999; Assonov et al., 2002) are around ten times greater than the solar elemental abundance ratio of 37 (Anders and Grevesse, 1989). Whilst $^{15}$N-depleted (‘light’) nitrogen is correlated with the release of solar wind D-depleted hydrogen from the outer rims of lunar soil grains, $^{15}$N-enriched (‘heavy’) nitrogen correlates with the release of D-enriched hydrogen from a planetary source (Hashizume et al., 2000). Thus, the observed trend of higher average $\delta^{15}$N values with decreasing soil antiquity may not be caused by a secular variation in solar wind composition, but instead may be recording a change in the relative contribution of a planetary nitrogen source to the lunar surface (Hashizume et al., 2002). Therefore, current evidence suggests that nitrogen in lunar soils represents mixing between a ‘heavy’ planetary nitrogen component and a solar wind component of ‘light’ composition (Geiss and Bochsler, 1982; Wieler et al., 1999; Hashizume et al., 2000, 2002; Assonov et al., 2002; Marty et al., 2003; Ozima et al., 2005; Füri et al., 2012). This has been supported by recent direct measurements of solar wind nitrogen implanted into the Genesis Solar Wind Concentrator target, which gave a modern-day $\delta^{15}$N value for solar wind of $-407 \pm 7$‰ (Marty et al., 2011), comparable to the composition of protosolar nitrogen ($-383 \pm 8$‰ (Marty et al., 2011)), and to that of the Jovian atmosphere (Owen et al., 2001).

In light of this, Füri et al. (2012) calculated that most of the nitrogen released from Luna 24 soil grains is from a non-solar source with a planetary composition of between $+100$‰ and $+160$‰, assuming binary mixing with $^{15}$N-depleted solar wind, with the highest proportional contribution from a non-solar source being found in recently exposed Apollo soils (with $^{40}$Ar/$^{36}$Ar ratios $\leq 1.5$).

Characterising the source(s) of volatiles being delivered to the lunar surface is essential for a broader consideration of how volatiles can be transported around the Solar System and added to the surfaces and atmospheres of other rocky Solar System bodies after initial planetary formation. Indeed, the volatile elements included in this study form the majority of most terrestrial planet atmospheres. Thus, to begin to understand planetary atmospheric evolution, any likely exogenous volatile source(s) must be identified and their contributions taken into account. Further, volatiles present at the surface of the Moon and other airless bodies, such as larger asteroids, may prove vital to the further human exploration of the Solar System, providing resources to support life beyond Earth. From this perspective, knowledge of which volatiles are present must be coupled with information about where they originated, thereby enabling resource reservoir estimation and replenishment rate consideration.

To this end, a diverse set of Apollo lunar soils have here been analysed by stepped combustion, with the released gases simultaneously analysed for nitrogen, carbon, and noble gases (He, Ne, and Ar) in order to better constrain the contribution of non-solar nitrogen by taking into account multiple isotope systems, and their variations over discrete temperature increments.

2. SAMPLES

Five Apollo lunar soils, representing both extremely mature (e.g. $I_1/FeO = 94$) and immature (e.g. $I_1/FeO = 5.7$) samples, and covering a range of collection locations, from Apollo 12, 14, 15, 16, and 17, and including both shaded soils and samples collected at the surface of the regolith in unshaded locations, were chosen for stepped combustion analyses. Note that sample 12070 was analysed twice. (Table 1). Soil maturity is expressed as the ratio of observed ferromagnetic intensity of the sample ($I_1$) divided by its total Fe content ($FeO$); ferromagnetic intensity increases as a sample is exposed at the lunar surface via the production of non-charged Fe microparticles, termed nanophase Fe or np-Fe$^0$, resulting from reduction or sputtering processes acting upon lunar regolith materials to remove oxygen from ferrous Fe. Thus, the larger the ratio of $I_1/FeO$, the more mature the soil.

3. METHODS

For stepped combustion analyses, samples were weighed out and transferred into 4 mm × 4 mm clean platinum foil buckets in a Class 100 cleanroom (Pt foil cleaning methods are as described in Abernethy et al., 2013). Tweezers and spatulas used to transfer the sample fines from their respective vials to the platinum buckets were cleaned before use, and wiped with acetone using lint-free cloths between usages with separate samples.

The ‘Finesse’ mass spectrometric instrument used in this study is a custom-built mass spectrometer system, consisting of three dedicated static-mode mass spectrometers (one for carbon, one for nitrogen and argon, and a quadrupole for helium and neon), all linked via high vacuum lines to a common sample inlet and gas clean-up sections, and combustion furnace (e.g. Mikhail et al., 2014, and references therein).

3.1. Stepped combustion procedure

Samples were combusted in oxygen, supplied from CuO, in a double-walled quartz-ceramic furnace for 30 min at each temperature step, followed by 15 min for oxygen resorption, before transfer of the released gases to the clean-up section. For Run 1 of soil 12070, twelve combustion steps were used, in 100 °C steps from 300 to 1400 °C. However, for Run 2 of soil 12070, fifteen combustion steps were employed to acquire higher-resolution data, heating in 100 °C steps from 200 to 700 °C, then in 50 °C steps from...
Table 1
List of five lunar soil samples chosen for stepped combustion analysis (*Morris, 1978).

<table>
<thead>
<tr>
<th>Apollo Sample</th>
<th>Weight (mg) (Combustion)</th>
<th>Maturity ([FeO]/[MgO])</th>
<th>Description</th>
</tr>
</thead>
<tbody>
<tr>
<td>12070</td>
<td>Run 1: 2.063 Run 2: 3.211</td>
<td>Sub-mature (47°)</td>
<td>Collected as part of the contingency sample, taken from a location in front of the Lunar Module (LM)</td>
</tr>
<tr>
<td>14141</td>
<td>7.637</td>
<td>Immature (5.7°)</td>
<td>Coarse soil, collected from near the rim of Cone Crater</td>
</tr>
<tr>
<td>15040</td>
<td>5.662</td>
<td>Mature (94°)</td>
<td>Collected at Station 8, from the surface at the top of a trench</td>
</tr>
<tr>
<td>69921</td>
<td>5.377</td>
<td>Mature (90°)</td>
<td>Skim soil, collected from top 1 cm of regolith in shaded area next to a small (0.5 m) boulder at Station 9</td>
</tr>
<tr>
<td>72501</td>
<td>4.710</td>
<td>Mature (81°)</td>
<td>Collected from landslide at Station 2, at the base of the slope of the South Massif, close to Nansen Crater</td>
</tr>
</tbody>
</table>

750 to 850 °C, followed by 100 °C steps from 900 to 1400 °C. All other soil samples were combusted in seventeen temperature steps, in increments of 100 °C from 200 to 600 °C, then in 50 °C steps from 650 to 850 °C, followed by 100 °C steps from 900 to 1400 °C. During stepped combustion analyses, C, N, He, Ne, and Ar data were collected.

Gas fractions released during stepped combustion analyses were processed in the same way for each sample. First, the released gases were cryogenically separated using liquid nitrogen cooled traps, one of which was filled with a molecular sieve. Argon and neon were purified using Ti–Al getters, and nitrogen was purified using a CuO furnace to ensure no CO was present. Carbon yields (recorded as ng) were measured as parts per thousand (‰) (Table A; Fig. 1), where most of the N in the 900 °C step was lost. It must be noted that the combustion steps releasing the lightest nitrogen component are not coincident with the maximum N abundances, but instead are associated with some of the lower released abundances (relative to combustion steps at temperatures either side of 850–900 °C). Given a solar wind N isotopic signature of around −40‰ (Marty et al., 2011), this indicates that the main contributor to N abundances in lunar soils must be from a non-solar source with a much more 15N-enriched (heavier) isotopic composition, mixing in varying proportions with the less abundant light N component most likely derived from the solar wind (see Discussion section).

In contrast to N releases, hundreds of ng of C are released from the first few temperature steps, up to 500–600 °C. This is most likely associated with the release of terrestrial contaminants from the samples. However, above this temperature, two further releases of C are observed in all soils; one at roughly 600–900 °C, and a second at temperatures above 1000 °C (Table B; Fig. 2). Although large uncertainties are associated with some of these C releases, C isotopic composition varies considerably; in some samples (14141, 15040, and 72501) there is an apparent relative depletion in 13C associated with the 600–900 °C release.
Fig. 1. Step plots showing nitrogen release profiles for all soils analysed by stepped combustion. Abundances, measured on left-hand axis, are displayed as a histogram, and isotopic compositions, measured on the right-hand axis, are displayed as a line graph. Errors in many cases are smaller than the symbol size. (A = 12070 Run 1; B = 12070 Run 2; C = 14141; D = 15040; E = 69921; F = 72501).
peak of C, but in general, despite variations in isotopic composition, $\delta^{13}$C values for C released above around 600 °C are positive (i.e. $^{13}$C-enriched). The exception to this observation is 14141, which is the least mature soil in our sample set. Whilst this soil shows a similar isotope profile to the other, more mature, soils, in this case, $\delta^{13}$C values remain between $-35\%e$ and $-20\%e$, lower than for the other samples.
4.2. Noble gases (Argon)

Again, as previously noted by others (e.g. Hohenberg et al., 1970; Brilliant et al., 1994; Assonov et al., 2002), 40Ar, 36Ar, and N are co-released in all samples (Figs. 3 and 4), albeit with an additional release peak of 40Ar at lower temperatures (typically between 400 °C and 600 °C) that is not observed in releases of 36Ar, which gives rise to the elevated 40Ar/36Ar ratios at these lower temperatures (Table C). Assonov et al. (2002) explain this low temperature 40Ar release as being attributable to terrestrial atmospheric contamination and/or re-implantation of 40Ar liberated to the lunar atmosphere during regolith reworking events (this re-implanted 40Ar being much more loosely fixed at the surfaces of grains which have experienced little or no reworking than the bulk of the argon contained in the soils, which, after soil reworking, is immersed deeper within grains; thus re-implanted 40Ar is liberated at lower temperatures during analysis). As noted in previous studies (e.g. Assonov et al., 2002), atomic N/36Ar ratios above ~650 °C are relatively constant, at around ten times the solar ratio (as estimated by Anders and Grevesse, 1989).

4.3. Noble gases (Helium and Neon)

Typical release patterns for 4He and 20Ne from lunar soils analysed in this study are shown in Fig. 5. The majority of the Ne released by combustion appears to be from a solar source, which is fractionated from the pure solar values (taken from Pepin et al., 2012) upon implantation into the lunar surface (Fig. 6). This dominant solar Ne component is also mixed with a much less abundant cosmogenic Ne component in some samples, although even the highest temperature steps still appear to be a mixture of solar and cosmogenic Ne (Table F; Fig. 6). The 4He/20Ne ratios, which are comparable to those observed in other lunar soils, suggest that the majority of He is of solar wind origin as well.

5. DISCUSSION: SOLAR VS NON-SOLAR NITROGEN

Although spallation processes generated by incoming cosmic rays lead to a significant enrichment in 15N progressively over time as a sample resides at or near to the lunar surface, stepped combustion analyses of lunar basalts have demonstrated that such cosmogenic N is only present in very low abundances of between 0.20 ppm and 0.25 ppm (Mortimer et al., 2015). Lunar soils contain much greater abundances of total N, ranging from 18.56 ppm N in immature soil 14141, up to 108.04 ppm N in mature soil 69921, and so the effect of a cosmogenic N component on measured isotopic compositions of lunar soils will be negligible. Similarly, stepped combustion analyses of lunar basalts reveal a minor indigenous lunar nitrogen component (at +0.93 ± 9.3‰, and present in abundances of between 0.13 ppm and 0.90 ppm (Mortimer et al., 2015) and additional data that is contained in the Appendix to this present study (Table 1)), or <0.1 ppm to 0.5 ppm (Füri et al., 2015); this is volume-correlated, and as with any cosmogenic N, is several orders of magnitude less abundant than the total N released from lunar soils, which is mostly surface-correlated (e.g. Hashizume et al., 2000). Therefore, for the purposes of the following calculations, we assume that inputs from both cosmogenic and indigenous lunar N are so low in relation to the total N release from lunar soils that they have essentially no impact on weighted average (across all temperature steps) N isotopic compositions for each soil sample.

Assuming instead that the total N released from lunar soils represents a binary mixture of solar and non-solar nitrogen components, and taking into account the isotopic composition of the solar wind as measured by Genesis (at ~40‰; Marty et al., 2004), it is possible to carry out similar calculations to those of Füri et al. (2012) using data from this study, working backwards from the measured (mixed) N values to derive estimates for both the percentage contribution from a non-solar source of N to the lunar surface, and the likely isotopic composition of this non-solar extra-lunar contributor.

Weighted average nitrogen isotopic values(δ15Nrel) for the soils in this study vary between +1.0‰ and +65.7‰, with average atomic 36Ar/N ratios (normalised to the solar 36Ar/N value of 0.02714, taken from Anders and Grevesse, 1989) varying between 0.12 and 0.17 times the solar value (Table 2). The relative proportions of solar and non-solar N are calculated using the following equation:

\[ \delta^{15}N_{rel} = \frac{\delta^{15}N_{SW} + f_p \delta^{15}N_p}{1 + f_p} \]

where \( \delta^{15}N_{SW} \) and \( \delta^{15}N_p \) are the end-member isotopic compositions of solar wind nitrogen and non-solar ('planetary') nitrogen, respectively. \( f_p \) is the mixing proportion of non-solar to solar nitrogen, and is defined thus:

\[ f_p = \frac{N_p}{N_{SW}} \]

Clearly, with 36Ar/Nrel values varying between 1/9th and 1/6th of the solar ratio, either 36ArSW has been lost as a result of impact and gardening processes at the lunar surface (with NSW being preferentially retained as a result of implantation into grains and its greater reactivity with common elements within soil grains), or a significant proportion of non-solar nitrogen (Np) has been added to that derived from a solar source (negligible amounts of 36Ar are added from the non-solar source, as evidenced by the majority of other noble gases (i.e. Ne) being derived from solar and not planetary sources). Whilst the preferential loss of solar noble gases over solar nitrogen is certainly an important factor, average δ15Nrel values measured in lunar soils by stepped combustion methods display considerable enrichment in 15N relative to the solar end-member value, which strongly suggests that the majority of the N measured in lunar soils comes from a non-solar source. Therefore, assuming that no loss of 36ArSW has taken place, and that the difference between measured and solar 36Ar/N values is caused by the addition of non-solar N alone, should give an upper limit to the proportion of non-solar N present and provide the lower limit of the isotopic composition of the non-solar end-member.
Fig. 3. Release patterns for nitrogen (grey histogram, measured on the left-hand axis) and $^{36}$Ar (dotted black lines, measured on the right-hand axis) abundances, showing the co-release of $^{36}$Ar and N above $\sim 400$ °C. (A = 12070 Run 1; B = 12070 Run 2; C = 14141; D = 15040; E = 69921; F = 72501).
Thus, to calculate the mixing proportion ($f_P$) of non-solar to solar N ($N_{P}/N_{SW}$), the following equation is used:

$$
\frac{(^{36}\text{Ar}/N)_{tr}}{(^{36}\text{Ar}/N)_{SW}} = \frac{1 - f_R}{1 + f_P}
$$

where $(^{36}\text{Ar}/N)_{tr}$ is the average measured ratio for the soil sample, $(^{36}\text{Ar}/N)_{SW}$ is the solar ratio taken from Anders and Grevesse (1989), and $f_R$ is the fraction of solar gas lost.

To assess the likelihood of preferential loss of solar wind $^{36}\text{Ar}$, a solar wind $\text{Ar}/\text{Xe}$ atomic ratio of 23,360, calculated

Fig. 4. Release patterns of $^{40}\text{Ar}$ (black dotted lines, measured on left-hand axis) and $^{36}\text{Ar}$ (solid black lines, measured on right-hand axis), showing co-release of both isotopes at temperatures over approximately 650 °C. (A = 12070 Run 1; B = 12070 Run 2; C = 14141; D = 15040; E = 69921; F = 72501).
from data contained in Vogel et al. (2011) using direct measurements of the solar wind made with the Genesis Cz-Si targets, can be compared with Ar/Xe atomic ratios calculated from the bulk 36Ar abundances in this study and bulk Xe abundances from published literature for the same samples. Taking 132Xe abundances from Hintenberger et al. (1971) and Heymann et al. (1974), Ar/Xe = 23,400 for soil sample 12070 (Run 2), and Ar/Xe = 20708.66 for soil 72501, respectively. With Ar/Xe ratios so close to the directly-measured solar wind value, it seems that very little, if any, loss of solar wind 36Ar from lunar soils has occurred, relative to Xe which has the highest retention among the noble gases. Assuming no loss of 36ArSW means that a range of Np end-member isotopic compositions are required for these samples, from +87‰ to 12070 (Run 2), up to +140‰ for 72501 (Fig. 7). However, this range of isotopic compositions results in a much tighter range of mixing proportions between Np and NSW, with the non-solar nitrogen source contributing 77.15% of the total N in soil 14141, up to 87.25% of the total N in soil 15040 (Table 2). It is interesting to note that, although most sub-mature and mature soils have a similar proportion contribution from a non-solar source (at approximately 86–87% of the total N in the soils), the sample with the lowest maturity, 14141, also has the lowest contribution from non-solar sources of N, at 77.15%. Therefore, it seems that N contributions from non-solar sources are increasingly dominant as a soil matures at the lunar surface. Furthermore, despite slight differences in the mixing proportions of solar and non-solar nitrogen sources between samples, it is clear that in all lunar soils, solar N only accounts for a small portion of the total N present (<23% solar N). This conclusion is in agreement with those of Füri et al. (2012), who state that between 60% and 83% of the N present in low-antiquity 40Ar/36Ar ratios ≤ 1.5 Apollo soils comes from a non-solar source. All of the Apollo soils used in this study have 40Ar/36Ar ratios ≤ 1.5, except 14141, which has a 40Ar/36Ar ratio of 10.92. The range of Np isotopic compositions suggested here is also in close agreement with values (+90‰ to +130‰) proposed by Füri et al. (2012) for Luna 24 samples.

If, however, we assume that some degree of 36ArSW loss is also occurring alongside addition of non-solar N, theoretical mixing lines can be constructed between a solar N end-member (at ~400‰) and a non-solar (Np) end-member composition of +140‰ (the value suggested by negligible loss of 36ArSW from the most 15N-enriched soil, 72501) (Fig. 7). In this case, 12070 (Run 2) has experienced the largest loss of 36ArSW, at ~40% loss (fR = 0.42) (Table 2).

Soil maturity (I/FeO) is roughly positively correlated (R² = 0.58) with weighted average δ15N values (Fig. 8), indicating that, as a soil undergoes progressive processing at the lunar surface, this seems to be driving the variations observed in δ15N values. If, then, the potential loss of 36ArSW from all soil samples is also taken into account, and a non-solar end-member isotopic composition of +160‰ is assumed (as proposed by Füri et al., 2012), the range of relative contributions from a non-solar N source remains largely the same (71.6%, up to 83% non-solar N) (Table 2), confirming that non-solar N is the major contributor to the total N present in lunar soils, and that solar N contributes only around a quarter of the total N abundance. A non-solar composition of +160‰ requires losses of 36ArSW of between 18% and 48%, which again fits well with the calculations of Füri et al. (2012) (Fig. 9).

Although the binary mixing calculations based on the bulk measurements of Apollo soils already discussed in this paper are in excellent agreement with conclusions reached for Luna samples (Füri et al., 2012), both in terms of relative proportions of NSW and Np, and the range of isotopic compositions suggested for the non-solar end-member, the stepped heating data in this study can go further to consider how the relative proportions of NSW and Np change with the well-documented variations in nitrogen isotopic compositions as a sample is heated in discrete temperature steps. It must be noted, however, that such calculations can only be made for temperature steps that represent a two-component mixture of nitrogen. Therefore, temperature...
Fig. 6. Neon isotopic compositions for stepped combustion analyses of lunar soils. Most of the neon released is from a solar source, fractionated from true solar values during implantation. In some soils, this fractionated solar neon component is mixed with a minor cosmogenic neon component (grey dotted lines indicate mixing trend). Solar and cosmogenic end-member compositions are represented by dark circles and light grey ellipses (respectively). Instrument errors are 5%. (A = 12070 Run 1; B = 12070 Run 2; C = 14141; D = 15040; E = 69921; F = 72501).
steps which may contain a third nitrogen component (whether terrestrial atmospheric N or cosmogenic N) and steps that released no measurable abundance of 36Ar are not included in the following binary mixing calculations. Therefore, assuming that measured nitrogen isotopic compositions for temperature steps between 500°C to 1200°C (inclusive) represent binary mixing of N SW and N P, and assuming that the most 15N-enriched steps within this range have experienced no loss of 36ArSW, the variable isotopic compositions for each temperature step within a sample can be reconciled with an NP end-member composition that is more enriched in 15N than bulk data would suggest, with separate samples indicating NP end-member compositions between +210‰ and +300‰ (Fig. 10 and Table G). These NP end-member isotopic compositions are more 15N-enriched than measured values for most carbonaceous chondrite groups (Pearson et al., 2006), but are within the range of values measured in bulk IDPs, which also display ‘hotspots’ reaching up to +1300‰ (Floss et al., 2006), as well as being within the range of values measured in the chondrite Isheyovo, up to +4900 ± 300‰ (Briani et al., 2009), and the stony-iron meteorite Bencubbin, which has a maximum δ15N of +1033‰ (Franchi et al., 1986). Within an individual sample, the proportion of NP to N SW changes with each temperature step, with the largest contributions from N SW observed at 850°C to 900°C, coincident with the least 15N-enriched measured isotopic compositions. However, even at these temperature steps where N SW contributions are greatest, in most samples, the majority (between 50% and 61%) of the nitrogen released is derived from the non-solar NP source (Fig. 11). The one exception to this is 14141, which is the least mature of the soil samples analysed in this study; here, at 900°C, a non-solar NP component with an end-member composition of +260‰ contributes only 49% of the nitrogen released in that temperature step, with the remaining 51% of the released nitrogen being derived from a solar source. It therefore seems that stepped combustion data confirm the conclusions reached based on bulk data that the majority of nitrogen in lunar soils comes from a non-solar source in all but the most immature samples.

Table 2
Comparing mixing proportions of solar and non-solar N end-members in 3 different scenarios (no loss of 36ArSW and variable NP compositions; variable loss of 36ArSW and NP = +140‰; variable loss of 36ArSW and NP = +160‰).

<table>
<thead>
<tr>
<th></th>
<th>12070 Run 1</th>
<th>12070 Run 2</th>
<th>14141</th>
<th>15040</th>
<th>69921</th>
<th>72501</th>
</tr>
</thead>
<tbody>
<tr>
<td><strong>Assuming 0% loss of 36ArSW</strong></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>Weighted average isotopic composition (δ15N, ‰)</td>
<td>32.3</td>
<td>23.3</td>
<td>1.0</td>
<td>36.2</td>
<td>34.6</td>
<td>65.7</td>
</tr>
<tr>
<td>Non-solar end-member isotopic composition (δ15N, ‰)</td>
<td>120</td>
<td>87</td>
<td>120</td>
<td>100</td>
<td>100</td>
<td>140</td>
</tr>
<tr>
<td>Contribution of non-solar nitrogen (%)</td>
<td>83.15</td>
<td>86.90</td>
<td>77.15</td>
<td>87.25</td>
<td>86.90</td>
<td>86.25</td>
</tr>
<tr>
<td>Contribution of solar nitrogen (%)</td>
<td>16.85</td>
<td>13.10</td>
<td>22.85</td>
<td>12.75</td>
<td>13.10</td>
<td>13.75</td>
</tr>
<tr>
<td>Loss of solar wind 36Ar (%)</td>
<td>0</td>
<td>0</td>
<td>0</td>
<td>0</td>
<td>0</td>
<td>0</td>
</tr>
<tr>
<td>fR</td>
<td>0.1685</td>
<td>0.1130</td>
<td>0.2285</td>
<td>0.1275</td>
<td>0.1310</td>
<td>0.1375</td>
</tr>
<tr>
<td>(1 − fR)/(1 + fP)</td>
<td>0.1685</td>
<td>0.1130</td>
<td>0.2285</td>
<td>0.1275</td>
<td>0.1310</td>
<td>0.1375</td>
</tr>
<tr>
<td>(36Ar/N)HN/(36Ar/N)SW</td>
<td>0.1685</td>
<td>0.1270</td>
<td>0.2319</td>
<td>0.1206</td>
<td>0.1282</td>
<td>0.1325</td>
</tr>
</tbody>
</table>

**Assuming NP = +140‰**

<table>
<thead>
<tr>
<th></th>
<th>12070 Run 1</th>
<th>12070 Run 2</th>
<th>14141</th>
<th>15040</th>
<th>69921</th>
<th>72501</th>
</tr>
</thead>
<tbody>
<tr>
<td>Weighted average isotopic composition (δ15N, ‰)</td>
<td>32.3</td>
<td>23.3</td>
<td>1.0</td>
<td>36.2</td>
<td>34.6</td>
<td>65.7</td>
</tr>
<tr>
<td>Non-solar end-member isotopic composition (δ15N, ‰)</td>
<td>140</td>
<td>140</td>
<td>140</td>
<td>140</td>
<td>140</td>
<td>140</td>
</tr>
<tr>
<td>Contribution of non-solar nitrogen (%)</td>
<td>80.05</td>
<td>78.40</td>
<td>74.26</td>
<td>80.80</td>
<td>80.15</td>
<td>86.20</td>
</tr>
<tr>
<td>Contribution of solar nitrogen (%)</td>
<td>19.95</td>
<td>21.60</td>
<td>25.74</td>
<td>19.20</td>
<td>19.85</td>
<td>13.80</td>
</tr>
<tr>
<td>Loss of solar wind 36Ar (%)</td>
<td>15</td>
<td>42</td>
<td>10</td>
<td>37</td>
<td>35</td>
<td>4</td>
</tr>
<tr>
<td>fR</td>
<td>0.15</td>
<td>0.42</td>
<td>0.10</td>
<td>0.37</td>
<td>0.35</td>
<td>0.04</td>
</tr>
<tr>
<td>fP</td>
<td>4.013</td>
<td>3.630</td>
<td>2.885</td>
<td>4.208</td>
<td>4.038</td>
<td>6.246</td>
</tr>
<tr>
<td>(1 − fR)/(1 + fP)</td>
<td>0.1696</td>
<td>0.1253</td>
<td>0.2317</td>
<td>0.1210</td>
<td>0.1290</td>
<td>0.1325</td>
</tr>
<tr>
<td>(36Ar/N)HN/(36Ar/N)SW</td>
<td>0.1685</td>
<td>0.1270</td>
<td>0.2319</td>
<td>0.1206</td>
<td>0.1282</td>
<td>0.1325</td>
</tr>
</tbody>
</table>

**Assuming NP = +160‰**

<table>
<thead>
<tr>
<th></th>
<th>12070 Run 1</th>
<th>12070 Run 2</th>
<th>14141</th>
<th>15040</th>
<th>69921</th>
<th>72501</th>
</tr>
</thead>
<tbody>
<tr>
<td>Weighted average isotopic composition (δ15N, ‰)</td>
<td>32.3</td>
<td>23.3</td>
<td>1.0</td>
<td>36.2</td>
<td>34.6</td>
<td>65.7</td>
</tr>
<tr>
<td>Non-solar end-member isotopic composition (δ15N, ‰)</td>
<td>160</td>
<td>160</td>
<td>160</td>
<td>160</td>
<td>160</td>
<td>160</td>
</tr>
<tr>
<td>Contribution of non-solar nitrogen (%)</td>
<td>77.2</td>
<td>75.6</td>
<td>71.6</td>
<td>77.9</td>
<td>77.6</td>
<td>83</td>
</tr>
<tr>
<td>Contribution of solar nitrogen (%)</td>
<td>22.8</td>
<td>24.4</td>
<td>28.4</td>
<td>22.1</td>
<td>22.4</td>
<td>17</td>
</tr>
<tr>
<td>Loss of solar wind 36Ar (%)</td>
<td>25</td>
<td>48</td>
<td>18</td>
<td>45</td>
<td>42</td>
<td>20</td>
</tr>
<tr>
<td>fR</td>
<td>0.25</td>
<td>0.48</td>
<td>0.18</td>
<td>0.45</td>
<td>0.42</td>
<td>0.2</td>
</tr>
<tr>
<td>(1 − fR)/(1 + fP)</td>
<td>0.1710</td>
<td>0.1269</td>
<td>0.2329</td>
<td>0.1215</td>
<td>0.1299</td>
<td>0.1360</td>
</tr>
<tr>
<td>(36Ar/N)HN/(36Ar/N)SW</td>
<td>0.1685</td>
<td>0.1270</td>
<td>0.2319</td>
<td>0.1206</td>
<td>0.1282</td>
<td>0.1325</td>
</tr>
</tbody>
</table>
temperature range used in this study (Fig. 3). This pattern of co-release suggests that both species are located in sites with similar energies required for gas release to occur, and consequently, if there has been any partial loss of gas, this would be expected to have happened to both $^{36}\text{Ar}$ and nitrogen equally. Therefore, the range of $N_P$ end-member compositions required to satisfy measured $\delta^{15}\text{N}$ values for each temperature step has been calculated based on an assumption of no loss of $^{36}\text{Ar}$ (Fig. 12, Table H). In this case, the non-solar end-member composi-
than the majority of the N P end-member compositions sug-
ging no loss of $^{36}$Ar$_{SW}$, the maximum cometary contribu-
tion allows a wider range of chondritic groups to be
included as possible contributing sources (e.g. CO, CI,
CM, CV), despite their average
$^{15}$N values being lighter
than that of the N P end-member compositions (N P values between $+5\%$ and $+300\%$) are required to explain the measured
values, assuming varying levels of $^{36}$Ar$_{SW}$ loss, with the
largest contributions from N$_{SW}$ occurring at 850 °C to 900 °C,
coinciding with the most $^{15}$N-depleted releases. However,
even at temperatures where N$_{SW}$ contributions to the total
nitrogen release are greatest, the majority of the nitrogen
being released (between 50% and 61%) is still derived from
a non-solar source in most soils. The exception is 14141,
which is the least mature soil analysed in this study (I$_{FeO}$ = 5.7); in this sample, N$_{SW}$ contributes 51% of
the total nitrogen release at 900 °C.

Alternatively, Ar/Xe ratios measured in soils are very
similar to those measured for the solar wind, suggesting
that there has been no significant loss of $^{36}$Ar$_{SW}$ from soil
samples. Furthermore, the matching release patterns of $^{36}$Ar
and nitrogen across the whole temperature range used
in this study suggest that no preferential loss of $^{36}$Ar$_{SW}$ has
occurred. Therefore, a model of nitrogen mixing based on
this alternative assumption is considered to be most likely.
Following on from this, if the fraction of solar gas lost is
taken to be zero ($f_{R} = 0$), and the equation
\[ \frac{1}{f_{R}} = R \]
the end-member composition of the non-solar source (N$_{P}$)
can be calculated using the following relationship:
\[ \delta^{15}N_{u} = \frac{(\delta^{15}N_{P} + (\delta^{15}N_{SW} - f_{R} - \delta^{15}N_{P}) \times R)}{(1 - f_{R})} \]

With no loss of $^{36}$Ar$_{SW}$, a wider range of non-solar end-
member compositions (N$_{P}$ values between $+5\%$ and
$+300\%$) is required, although this is perhaps expected when
the range of material being delivered to the lunar surface
from different sources is considered. Again, the majority
of the nitrogen measured in lunar soils is derived from
a non-solar source, with the calculated N$_{P}$ end-member
isotopic compositions being consistent with a likely IDP
source for the N$_{P}$ component, in agreement with the con-
cclusions of Füri et al. (2012), and/or consistent with various
carbonaceous chondrite sources of nitrogen. Possible addi-
tions from a cometary source mixing with a chondritic
source cannot contribute more than $\sim 10\%$ to the N$_{P}$ non-
solar component.

6. CONCLUSIONS

Simultaneous measurements of carbon, nitrogen, and
noble gases in lunar soils by stepped combustion analysis
reveal the characteristic release pattern of heavy-light
heavy nitrogen, as well as reconfirming that $^{40}$Ar, $^{36}$Ar,
and N releases are well-correlated (above 650 °C). An
excess of $^{40}$Ar (relative to releases of $^{36}$Ar) at low tempera-
tures is most likely caused by terrestrial atmospheric con-
tamination, or possibly by the re-implantation of liber-
ated $^{40}$Ar from the lunar atmosphere. Neon isotopes
measured by stepped combustion reveal that the majority
of neon in lunar soils comes from a solar source, with a
minor cosmogenic component released at only the highest
temperature steps.

However, the release of $^{15}$N-enriched nitrogen from soils
shows that it is not purely solar in origin (given a solar wind
nitrogen signature of $-407\%$), and calculations reveal that
up to $87\%$ of the nitrogen measured in bulk lunar soils
could come from a non-solar/planetary source, with an iso-
topic signature of between $+87\%e$, up to $+160\%e$. Taking
measured nitrogen isotopic compositions and abundances
of $^{36}$Ar and N for each temperature step, more $^{15}$N-
enriched non-solar end-member compositions of between
$+210\%e$ and $+300\%e$ are required to explain the measured
values, assuming varying levels of $^{36}$Ar$_{SW}$ loss, with the
largest contributions from N$_{SW}$ occurring at 850 °C to 900 °C,
coinciding with the most $^{15}$N-depleted releases. However,
even at temperatures where N$_{SW}$ contributions to the total
nitrogen release are greatest, the majority of the nitrogen
being released (between 50% and 61%) is still derived from
a non-solar source in most soils. The exception is 14141,
which is the least mature soil analysed in this study (I$_{FeO}$ = 5.7); in this sample, N$_{SW}$ contributes 51% of
the total nitrogen release at 900 °C.

Alternatively, Ar/Xe ratios measured in soils are very
similar to those measured for the solar wind, suggesting
that there has been no significant loss of $^{36}$Ar$_{SW}$ from soil
samples. Furthermore, the matching release patterns of $^{36}$Ar
and nitrogen across the whole temperature range used
in this study suggest that no preferential loss of $^{36}$Ar$_{SW}$ has
occurred. Therefore, a model of nitrogen mixing based on
this alternative assumption is considered to be most likely.
Following on from this, if the fraction of solar gas lost is
taken to be zero ($f_{R} = 0$), and the equation
\[ \frac{1}{f_{R}} = R \]
the end-member composition of the non-solar source (N$_{P}$)
can be calculated using the following relationship:
\[ \delta^{15}N_{u} = \frac{(\delta^{15}N_{P} + (\delta^{15}N_{SW} - f_{R} - \delta^{15}N_{P}) \times R)}{(1 - f_{R})} \]

With no loss of $^{36}$Ar$_{SW}$, a wider range of non-solar end-
member compositions (N$_{P}$ values between $+5\%$ and
$+300\%$) is required, although this is perhaps expected when
the range of material being delivered to the lunar surface
from different sources is considered. Again, the majority
of the nitrogen measured in lunar soils is derived from
a non-solar source, with the calculated N$_{P}$ end-member
isotopic compositions being consistent with a likely IDP
source for the N$_{P}$ component, in agreement with the con-
cclusions of Füri et al. (2012), and/or consistent with various
carbonaceous chondrite sources of nitrogen. Possible addi-
tions from a cometary source mixing with a chondritic
source cannot contribute more than $\sim 10\%$ to the N$_{P}$ non-
solar component.
Fig. 10. Isotopic composition at each temperature step taken to indicate binary mixing only versus $^{36}\text{Ar}/^{36}\text{Ar}_{\text{SW}}$ ratios. For each sample, the yellow line indicates $N_{\text{SW}}$ end-member composition, and the red circle indicates the $N_{\text{P}}$ end-member composition required, which varies between samples from $+210^{\circ}$ to $+300^{\circ}$. Solid black lines are theoretical mixing lines between $N_{\text{SW}}$ and $N_{\text{P}}$ end-members, assuming different fractions of $^{36}\text{Ar}_{\text{SW}}$ loss. See Table G for data.
What is clear from these results is that there is an apparent decoupling between the sources of volatiles and their release temperatures and profiles following processing of material at the lunar surface. Although the most solar-like Ne compositions are released at the lowest temperatures, additions of non-solar N, mixing with solar N, results in the greatest proportion of solar N being released at higher temperatures of around 900°C. Further, despite 36Ar being purely solar and N being derived from a mixture of solar and non-solar sources, both are well-correlated and co-released across the whole temperature range. Therefore, the sites volatiles are initially delivered to are not necessarily the sites where they are ultimately stored and later released from during analysis.

A further implication of these findings is that, assuming the Moon’s surface has experienced volatile additions from similar sources to the Earth’s atmosphere, chondritic and/or IDP sources appear to play the dominant role in the formation of planetary atmospheres, not comets, at least in the case of nitrogen.

Fig. 11. Variations in the percentage of nitrogen released from lunar soils derived from a non-solar planetary source (N_p) and from a solar source (N_SW) at individual combustion temperature steps, assuming some degree of solar wind loss has occurred. Note that the lowest contribution from N_p (and therefore the greatest contribution from N_SW) occurs at 850°C to 900°C in each soil sample.
Fig. 12. Isotopic composition at each temperature step taken to indicate binary mixing only versus $\left( {^{36}\text{Ar}/N}_u / {^{36}\text{Ar}/N}_{\text{SW}} \right)$ ratios. For each sample, the yellow line indicates $N_{\text{SW}}$ end-member composition, and the red line indicates the $N_P$ end-member composition required, which varies both within and between samples from $+5\%$ to $+300\%$. Solid black lines are theoretical mixing lines between $N_{\text{SW}}$ and $N_P$ end-members, assuming no loss of $^{36}\text{Ar}_{\text{SW}}$. See Table H for data.
ACKNOWLEDGEMENTS

The authors wish to thank CAPTEM for allocation of lunar samples. J.M. thanks the Science and Technology Facilities Council (STFC) and The Open University for PhD studentship funding. This work has been partially funded by the STFC (Grant Numbers ST/I001298/1 and ST/L000776/1-Project N to M.A.).

APPENDIX A. SUPPLEMENTARY DATA

Supplementary data associated with this article can be found, in the online version, at http://dx.doi.org/10.1016/j.gca.2016.08.006.

REFERENCES


APPENDIX A. SUPPLEMENTARY DATA

Supplementary data associated with this article can be found, in the online version, at http://dx.doi.org/10.1016/j.gca.2016.08.006.


*Associate editor:* David L. Shuster