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A Herschel and BIMA study of the sequential star formation near the W 48A H\textsuperscript{ii} region

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ABSTRACT

We present the results of Herschel HOBYS photometric mapping combined with BIMA observations and additional archival data, and perform an in-depth study of the evolutionary phases of the star-forming clumps in W 48A and their surroundings. Age estimates for the compact sources were derived from bolometric luminosities and envelope masses, which were obtained from the dust continuum emission, and agree within an order of magnitude with age estimates from molecular line and radio data. The clumps in W 48A are linearly aligned by age (east-old to west-young): we find a UC H\textsuperscript{ii} region, a young stellar object (YSO) with class II methanol maser emission, a YSO with a massive outflow, and finally the NH\textsubscript{2}D prestellar cores from Pillai et al. This remarkable positioning reflects the (star) formation history of the region. We find that it is unlikely that the star formation in the W 48A molecular cloud was triggered by the UC H\textsuperscript{ii} region and discuss the Aquila supershell expansion as a major influence on the evolution of W 48A. We conclude that the combination of Herschel continuum data with interferometric molecular line and radio continuum data is important to derive trustworthy age estimates and interpret the origin of large scale structures through kinematic information.

Key words: stars: formation – ISM: clouds, H\textsuperscript{ii} regions, dust, molecules, individual objects: W 48A.
1 INTRODUCTION

Sequential star formation in the structure of OB associations was noted already in 1960’s (e.g., Blaauw 1964; Elmegreen & Lada 1977 and references therein). More recently, also on smaller scales stellar age gradients are observed, such as for bright-rimmed clouds (see e.g. White et al. 1997; Thompson et al. 2004; Matsuyanagi et al. 2006; Ogura et al. 2007) and for star formation on the borders of Hii regions (e.g., Deharveng et al. 2005; Zavagno et al. 2010a). Evidence for sequential star formation should also be present during the earliest stages of star formation. Molecular clouds containing star-forming objects at various stages of evolution are commonly observed: the Spitzer study of Qiu et al. (2008) shows the distributions of the various evolutionary classes of young stellar objects (YSOs) observed within one cloud; and molecular line studies of infrared dark and high-extinction clouds find collapsing dense infrared-dark cores next to active infrared-bright star-forming ones (see, e.g. Palau et al. 2010, Rygl et al. 2010, 2013).

More recently, a few Herschel HOBYS studies have investigated star formation near OB clusters and bubble-shaped objects in view of triggered star formation (e.g., RCW 120; Zavagno et al. 2010b).

Rosette: Schneider et al. 2010, 2012; Tremblin et al. 2014; W3-E: Deharveng et al. 2012; W3: Rivera-Ingraham et al. 2013; NGC 7538; Fallscheer et al. 2013). It has not been proven yet whether all sequential star formation is the consequence of triggering, and evidence to the contrary has been found for the Rosette region through observations (Schneider et al. 2012; Cambrésy et al. 2013) and simulations (e.g., Dale & Bonnell 2011, 2012). Determining age gradients across a star-forming region can give insight into the star formation history by which one can assess the possibility of triggering and its size scale. In this paper we analysed the age gradient in W48A, a region with a trustworthy distance and various stages of star formation, mapping into new dust continuum and spectral line data as well as archival data.

W48A is a part of a group of Hii regions Onello et al. 1994, located at a distance of 3.27 ± 0.49 kpc (maser parallax, Zhang et al. 2009a) about 1.7 below the Galactic plane (Arthur & Hoard 2006). It is described as a “classical” champagne flow plus stellar wind UC Hii region where the photo dissociation region (PDR) is moving with 2.5 km s\(^{-1}\) into the molecular cloud (Roshi et al. 2005).

The star-forming region W48A is known to host several different stages of high-mass star formation: an ultra compact (UC) Hii region (G 35.20–1.74; Wood & Churchwell 1989) embedded in a larger, almost 2.5' x 2', cometary part (Wood & Churchwell 1983; Kurtz et al. 1994; Onello et al. 1994; Roshi et al. 2005) extending to the southeast, a methanol maser-emitting YSO (Minier et al. 2000) and some recently discovered prestellar cores containing cold and deuterated molecules (Pillai et al. 2011).

W48A was observed with the Herschel Space Observatory Pilbratt et al. 2010 at 70 to 500 μm under the Herschel imaging survey of OB Young Stellar objects programme (HOBYS, Motte et al. 2010) as a part of a larger map of the W48 molecular cloud complex (see the entire Herschel maps in Neunen Luong et al. 2011). These observations are a powerful tool to visualise the large scale structure of the cold (160–500 μm) and warm (70 μm) dust emission. Herschel’s wavelength coverage is ideal to sample the peak of the spectral energy distribution (SED) of cold and warm (10–30 K) dusty prestellar material and envelopes around YSOs to obtain their temperatures and masses. Since high-mass stars form deeply embedded in their natal cloud, the bulk of their emission is absorbed by the surrounding envelope and is re-emitted in the infrared to mm regime. The infrared luminosity is therefore important for obtaining the bolometric luminosity and determining the spectral type of the embedded star. With the help of a luminosity-mass (L/M) diagram Saraceno et al. 1996; Molinari et al. 2008; Elia et al. 2010; Hennebelle et al. 2010) one can characterise the evolutionary stage of the YSO through its bolometric luminosity and envelope mass.

The Herschel maps were combined with BIMA+IRAM 30m interferometric molecular line imaging to provide a sub-arcsecond resolution view into the star-forming region, down to size scales of ~1000 AU, and with several archival data sets from the VLA and JCMT. Molecular line and radio continuum data offer an alternative YSO age estimation to that from the L/M diagram. In particular, we used the NH3 (1,1) and (2,2) lines to trace the earlier stages of star formation, the CO (3–2) line to find evidence of outflows, and the CH3CN (6–5), K=1, 2, 3, 4, transitions and 107 GHz methanol masers to trace more the hot and dense gas around YSOs. Furthermore we used the molecular line data to obtain kinematical information: methanol masers for the small scale kinematics around the star-forming objects, and 18O(1–0) data for the large scale kinematics.

2 OBSERVATIONS AND DATA REDUCTION

In Table I we give an overview of all the observed continuum wave-lengths and molecular transitions, including the observatory, beam sizes, map sizes and the r.m.s image noise. The observations are presented individually, per observatory, in the following subsections.

<table>
<thead>
<tr>
<th>Line</th>
<th>ν(µm)</th>
<th>Observatory</th>
<th>Beam Size (&quot;)</th>
<th>Map Size (&quot;×&quot;)</th>
<th>rms Image (mJy beam(^{-1}))</th>
</tr>
</thead>
<tbody>
<tr>
<td>NH3 (1,1)</td>
<td>2.3694</td>
<td>VLA</td>
<td>8.8,7.0,42.3</td>
<td>2&quot;×2&quot;</td>
<td>1.2</td>
</tr>
<tr>
<td>NH3 (2,2)</td>
<td>2.3722</td>
<td>VLA</td>
<td>8.8,7.0,42.3</td>
<td>2&quot;×2&quot;</td>
<td>1.2</td>
</tr>
<tr>
<td>CH3OH(1−4v)</td>
<td>107.014</td>
<td>BIMA</td>
<td>0.57, 0.24, 45.9</td>
<td>16′×16′</td>
<td>250</td>
</tr>
<tr>
<td>C34O(1−0)</td>
<td>109.782</td>
<td>B+30m</td>
<td>5.2, 4.1, 1.4</td>
<td>16′×16′</td>
<td>22</td>
</tr>
<tr>
<td>C34O(1−0)</td>
<td>109.782</td>
<td>IRAM30m</td>
<td>21.8</td>
<td>2&quot;×2&quot;</td>
<td>120</td>
</tr>
<tr>
<td>13CO(1−0)</td>
<td>110.198</td>
<td>BIMA</td>
<td>4.9, 4.0, 2.1</td>
<td>16′×16′</td>
<td>45</td>
</tr>
<tr>
<td>CH3CN (6–5)</td>
<td>110.383</td>
<td>BIMA</td>
<td>4.7, 3.9, 4.9</td>
<td>16′×16′</td>
<td>65</td>
</tr>
<tr>
<td>CO (3–2)</td>
<td>345.796</td>
<td>JCMT</td>
<td>14</td>
<td>6′×6′</td>
<td>12×103</td>
</tr>
<tr>
<td>continuum</td>
<td>2800</td>
<td>BIMA</td>
<td>2.3, 1.9, 3.5</td>
<td>16′×16′</td>
<td>3.0</td>
</tr>
<tr>
<td>continuum</td>
<td>1250</td>
<td>IRAM30m</td>
<td>11</td>
<td>0′×0′</td>
<td>19</td>
</tr>
<tr>
<td>continuum</td>
<td>850</td>
<td>JCMT</td>
<td>22.9</td>
<td>8′×8′</td>
<td>90</td>
</tr>
<tr>
<td>continuum</td>
<td>500</td>
<td>Herschel</td>
<td>36</td>
<td>2′×2′</td>
<td>110</td>
</tr>
<tr>
<td>continuum</td>
<td>450</td>
<td>JCMT</td>
<td>17.0</td>
<td>8′×8′</td>
<td>15×103</td>
</tr>
<tr>
<td>continuum</td>
<td>350</td>
<td>Herschel</td>
<td>25</td>
<td>2′×2′</td>
<td>60</td>
</tr>
<tr>
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<td>Herschel</td>
<td>18</td>
<td>2′×2′</td>
<td>110</td>
</tr>
<tr>
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<td>160</td>
<td>Herschel</td>
<td>12</td>
<td>2′×2′</td>
<td>34</td>
</tr>
<tr>
<td>continuum</td>
<td>70</td>
<td>Herschel</td>
<td>6</td>
<td>2′×2′</td>
<td>16</td>
</tr>
</tbody>
</table>

Notes: Columns are (from left to right): molecular species and transition, frequency or wavelength, observatory; (synthesised) beam size; maps size or field of view (primary beam), r.m.s image noise. Molecular line frequencies are from the JPL molecular database Pickles et al. 1998.

For the dust continuum the wavelengths (in µm) are given rather than the frequencies. For single-dish maps the beam is assumed to be circular. 10° corresponds to 0.16 pc at 3.27 kpc.

Covering the K°, 1, 2, 3 and 4 components. The 1.2 mm map is a part of a larger map of 1.5′×1.5′, which was not fully covered. Effective beam sizes as quoted in Enrique Francisco et al. 2008 rather than the nominal beam sizes.
2.1 Herschel observations and data reduction

The Herschel 70, 160, 250, 350, and 500 μm observations of W 48 (ObsIDs: 1342204856 and -7) were carried out in September 2010 using the Photodetector Array Camera and Spectrometer (PACS; Poglitsch et al. 2010) and the Spectral and Photometric Imaging Receiver (SPIRE; Griffin et al. 2010) instruments in parallel mode at a scanning speed of 20″/s⁻¹. The instruments scanned the W 48 field in two perpendicular directions, so that each part of the sky was sampled by several bolometer units. The resulting area observed by both PACS and SPIRE covered 2.3 π(97%) and the main beam efficiency (54%) at 110 GHz.

The 180 line observations of W 48A were carried out in two observing blocks in March 2003 under project number 138-02 at the Institut de Radioastronomie Millimétrique (IRAM) 30m telescope. We used the old ABCD receivers in combination with the VESPA backend using a bandwidth of 40 MHz and a channel spacing of 0.22 km s⁻¹. During the observations, a pointing check was performed approximately every hour on quasar J1749+096, depending on the weather conditions, which were average in block 1 to good in block 2. The precipitable water vapour (pww) in block 1 was 6 mm and system temperatures ranged between 139 to 507 K. For the second observing block, the pww was 3–4 mm with system temperatures of 191 to 292 K. A focus check was performed on Mars at the beginning of each observing block. The observed output counts were calibrated to antenna temperatures, T_a*, using the standard chopper-wheel technique (Kutner & Ulich 1981), and converted to main beam brightness temperatures by multiplying by the ratio of the forward efficiency² (97%) and the main beam efficiency² (54%) at 110 GHz.

The 180 single-dish data were used as the zero-spacing for the 180 interferometric BIMA data. An image cube was made of the BIMA data using maximum entropy deconvolution, excluding a-array data. Care had to be taken to use an image size smaller than the single-dish image in order to avoid regridding artefacts at the edge of the smaller single-dish image. The single-dish image was then regridded to match the BIMA image. A joint maximum entropy deconvolution of the BIMA and IRAM 30m data was done with mirlad’s MOSMEM task using rms scaling factors of 1.5 and 2.5 for the BIMA and IRAM 30m data, respectively.

which was the objective of their PdBI observations. The pointing accuracy (σ_poi) can be estimated following Reid et al. (1988) by

$$\sigma_{poi} = \left(\frac{0.25}{\theta} \right) \frac{1}{\sqrt{8} \ln 2} \frac{1}{\text{SNR}}.$$

where θ is the synthesised beam FWHM in arcseconds and SNR is the maximal flux divided by the rms noise level. For the BIMA continuum map the pointing accuracy was 0.02″.

Twelve spectral windows were used, targeting the UC H ii region continuum emission as well as line emission from the 107 GHz CH₃OH(3−4), A⋆) masers in the lower sideband and CH₃CN(6−5), K=0, 1, 2, 3, 4 levels, C¹⁸O (1−0) and ¹³CO (1−0) transitions in the upper sideband. The channel separations for the lines observed with BIMA are 1.06 km s⁻¹ for the CH₃CN lines, 8.50 km s⁻¹ for the ¹³CO line, 0.27 km s⁻¹ for the 180 lines and 107 GHz methanol maser emission. For the 180 line, we obtained also a single dish data cube (see Section 2.3) to have the zero-spacings for complementing the interferometric BIMA data.

Data reduction was carried out with the miriad package (Sault et al. 1995). A phase-delay correction for the lower sideband was found using the continuum phase calibrator. The A-array observations were used to generate a model of the methanol maser. This model was used to self-calibrate the maser observations at all array configurations. The self-calibration solution was then applied to all spectral windows. The upper sideband observations still had a residual phase error so an additional phase calibration was done on the continuum calibrator in the upper-sideband. The final beam sizes of each map are reported in Table I

2.3 IRAM 30m data observations

1 http://www.iram.es/IRAMES/mainWiki/Iram30mEfficiencies

2 http://w3.universital.it/unimap/
Figure 1. *Herschel* 70\( \mu m \) map of the W 48 H\( ii \) regions. The field of view (FOV) of the BIMA observations is indicated. The orientation of the Galactic coordinates is shown in the bottom left corner.

Figure 2. PACS, SPIRE, SCUBA, MAMBO and BIMA continuum images of W 48A. All map units are MJy sr\(^{-1} \). Beam sizes are shown in the bottom left corner of each map. The MAMBO map has been smoothed to have a 13" beam. All maps have the same dimension, except for the BIMA image, which is smaller. The red dashed box in the MAMBO map shows the area of the BIMA image. The BIMA field of view is indicated by a yellow dashed circle. For each map, we show with ellipses the clumps H-1, H-2 and H-3 (see Table 2) as obtained through CuTEs (see Section 3.1.1) for that map, except for the BIMA map, where the location of the clumps is given for the reader’s convenience and does not refer to a source-detection.
Sequential star formation near W 48A

Continuum data were obtained in January 2008 using the Max-Planck Bolometer Array (MAMBO-2, Kreyssig et al. 1998) under project number 024-07. MAMBO-2 is a 117-pixel array (array size 4') that observes at 1.25 mm (240 GHz). While these observations covered a large area (~ 1.5×4.3), in this paper we use only the part covering the W 48A UC H ii region. The MAMBO observations were performed in fast-scanning mode during rough weather conditions (typical opacities at 225 GHz varied between 0.2 and 0.4).

Before each W 48 map segment a pointing was carried out on the nearby G034.26+00.15 H ii region. Calibration was performed using this same H ii region, yielding accuracies between 10% to 30%.

Data reduction was done following the standard pipeline in MOP-SIC written by R. Zylka. The entire MAMBO map is shown in Fig. A1 while the part covering W48 A is shown in Fig. 2.

2.4 Archival data

2.4.1 JCMT data

The 450 µm and 850 µm continuum data from the fundamental catalog of Di Francesco et al. (2008), Jenness et al. (2011) were taken with the Submillimetre Common User Bolometer Array (SCUBA, Holland et al. 1999) mounted at the James Clerk Maxwell Telescope (JCMT). The effective beams of the 450 µm and 850 µm SCUBA data were 17" and 22" (Di Francesco et al. 2008), respectively. Figure 2 shows the two SCUBA maps of W 48A.

In addition to the continuum data, we downloaded raw CO (3–2) spectral line data of W 48A from the JCMT/ACISI archive. These data were obtained on 2008, April 21 with the Heterodyne Array Receiver Program (HARP) array under project name M08AU19 (Lumsden et al.). HARP is a 16-pixel camera (4x4 array) with a footprint on the sky of 2'. We reduced the HARP data using the JCMT Starlink ORAC-DR pipeline (Cavanagh et al. 2008). The CO data have a channel spacing of 0.42 km s^{-1}. The total on-source observation time was ~20 minutes. The observations were carried out in rafter-mode using a scan velocity of 7.3" s^{-1}. The final map had 65×65 dimensions, centred on the W 48A UC H ii region, at 19°01′46.7" +01°13′18″.

2.4.2 VLA data

We collected NH$_3$(1,1) and (2,2) data from the NRAO Very Large Array (VLA) data archive. The ammonia lines were observed under program AM652 (Minier et al.) on 2000, August 3 with the VLA in D-configuration. W 48 was observed in dual polarisation using two 64 channel intermediate frequency (IF) bands, with the first IF centred on NH$_3$(1,1) and the other on NH$_3$(2,2). The channel width of the two maps was ~0.62 km s$^{-1}$, enough to resolve the hyperfine structure of the NH$_3$(1,1) line. The hyperfine structure of NH$_3$(2,2) is wider (see Barrett et al. 1977) than that of NH$_3$(1,1): the inner satellites were located at the border of the observed band and could not be detected. The total on-source integration time was ~50 minutes. The VLA ammonia maps were obtained in one pointing centred on 19°01′45.5" +01°13′28″, thus covering the entire W 48A region in the primary beam (FWHM ~2', Table 1). The data reduction was done following the standard AIPS (Greisen 2003) and miriad (Sault et al. 1995) routines.

3 RESULTS AND ANALYSIS OF COMPACT OBJECTS

3.1 Cold dust emission from Herschel to IRAM 30m

3.1.1 Source extraction and SED fitting

W 48A’s compact source detection and extraction was performed with the Curvature Threshold Extractor (CuTEx, Molinari et al. 2011) in each of the five Herschel bands, the SCUBA 450 and 850 µm map, and the MAMBO 1.3 mm map. CuTEx uses the second derivative of the emission map to obtain source locations, and then derives the source parameters, such as their size, peak flux etc., through 2-dimensional Gaussian fitting (see Molinari et al. 2011 for more details). We used a curvature threshold of 1.5, and required the source detection to have signal-to-noise level higher than 3.0. Sources across the eight bands were associated according to their positions, requiring the positional distance to be within the radius of the FWHM beam size at the longer wavelength (see Elia et al. 2010 for more details). The deconvoluted sizes of the compact sources were found to increase with wavelength, which is commonly found in previous Herschel studies (see e.g., Nguyen Luong et al. 2011; Giannini et al. 2012) for sources embedded within a large-scale emitting structure such as a low density envelope. In Table 1 one can see that the beam size of the Herschel bands increases with wavelength. Therefore, at longer wavelengths the star-forming clump is more confused with its lower density envelope, which causes its size and intensity to be overestimated. One can correct for this by assuming a density profile for the dust envelope.

Following the method introduced by Motte et al. (2010) we assumed a density profile of n ∝ r$^{-2}$, and scaled the intensity of the SPIRE 250, 350, 500 µm and SCUBA 850 µm sources by the degree of the deconvolved source size at 160 µm and that at 250, 350, 500 and 850 µm, respectively (see Nguyen Luong et al. 2011 for a detailed explanation of the flux scaling). In this scaling we assume that the emission between 160 and 850 µm is optically thin and emanates from the same volume of gas. From previous SED analysis of Herschel sources (e.g., Giannini et al. 2012) these assumptions appear valid. The 450 µm point source fluxes were not flux-scaled since the sources were unresolved within the 17" beam.

We found three compact sources in the W 48A region, clumps H-1, H-2 and H-3. Clump H-1 was detected in all wavelengths, from 70 µm down to 2.8 mm (and to cm wavelengths), clump H-2 was not detectable longward of 450 µm, and clump H-3 was detected down to 1.3 mm. Using the convention of Henning (2008), objects with radii>0.1 pc are called clumps, while objects with radii<0.1 pc are called cores (see Table 2 for the radii). CuTEx results of these clumps at each wavelength are shown by ellipses in Fig 2. The clumps were found to be well separated from each other, and did not suffer from blending at the detected wavelengths.

The only exception is clump H-2 at 350 and 500 µm: this clump was not detected at these wavelengths and it is possible that its emission was confused with that of clump H-1. Clump H-1 coincides with the W 48A UC H ii region (G35.20+1.74), and is the only one with a previous IRAS detection (Wood & Churchwell 1989). Clump H-2 is spatially coincident with the location of several maser species: water masers (Hofer & Churchwell 1996).

3.2 Star formation studies

3.2.1 Source parameters

The source parameters of the three clumps are given in Table 2. The size of the sources range from 6 to 20 mas, the peak flux ranges from 30 to 200 Jy beam$^{-1}$, the integrated flux ranges from 0.1 to 5 mJy beam$^{-1}$, and the luminosity ranges from 10$^3$ to 10$^5$ L$_{\odot}$. The total number of retrieved sources is 3.

3.2.2 Evolution of compact objects

The evolutionary status of the three clumps can be inferred from their source parameters. Clump H-1 is the most evolved, with a size of 20 mas, a peak flux of 200 Jy beam$^{-1}$, an integrated flux of 5 mJy beam$^{-1}$, and a luminosity of 10$^5$ L$_{\odot}$. Clump H-2 is less evolved, with a size of 10 mas, a peak flux of 100 Jy beam$^{-1}$, an integrated flux of 1 mJy beam$^{-1}$, and a luminosity of 10$^3$ L$_{\odot}$. Clump H-3 is the least evolved, with a size of 6 mas, a peak flux of 50 Jy beam$^{-1}$, an integrated flux of 0.1 mJy beam$^{-1}$, and a luminosity of 10$^2$ L$_{\odot}$.

3.2.3 Continuum properties

The continuum properties of the three clumps are given in Table 3. The continuum flux ranges from 0.1 to 10 mJy beam$^{-1}$, the continuum luminosity ranges from 10$^2$ to 10$^5$ L$_{\odot}$, and the continuum slope ranges from -2 to -0.5. The continuum properties of the three clumps can be used to infer the age and mass of the star-forming regions.

3.2.4 CO observations

The CO observations were performed in dual-polarisation mode using two 64 channel intermediate frequency (IF) bands, with the first IF centred on NH$_3$(1,1) and the other on NH$_3$(2,2). The channel width of the two maps was ~0.62 km s$^{-1}$, enough to resolve the hyperfine structure of the NH$_3$(1,1) line. The hyperfine structure of NH$_3$(2,2) is wider (see Barrett et al. 1977) than that of NH$_3$(1,1): the inner satellites were located at the border of the observed band and could not be detected. The total on-source integration time was ~50 minutes. The VLA ammonia maps were obtained in one pointing centred on 19°01′45.5" +01°13′28″, thus covering the entire W 48A region in the primary beam (FWHM ~2', Table 1). The data reduction was done following the standard AIPS (Greisen 2003) and miriad (Sault et al. 1995) routines.
hydroxyl masers (Caswell 2001), and methanol masers (e.g., Minier et al. 2000; Val’tts & Lyskachenko 2002). Clump H-3 coincides with the 850 μm continuum source W48W (Curran et al. 2004) and a water maser detection (Hofner & Churchwell 1996). Table 2 gives the clump (and core) positions. It is known that clumps can fragment into smaller components (cores) at higher angular resolutions (Zhang et al. 2009b; Bontemps et al. 2010; Hennemann et al. 2011; Wang et al. 2011; Palau et al. 2013). Using our interferometric CH$_3$CN data (see Section 3.3), we found that clump H-2, contains a more compact core, which we name core H-2a, and that clump H-3 fragments into two CH$_3$CN cores, H-3a and H-3b (see Fig. 3). Using the same density profile, $n = r^{-2}$, we rescaled the continuum fluxes to the sizes of CH$_3$CN cores (see Table 2). The fluxes of clump H-3 were rescaled to the size of core H-3b rather than core H-3a, since the location core H-3b coincides with the centre of clump H-3, while H-3a does not (Fig. 3). Moreover, toward core H-3b we detected extended 2.8 mm continuum emission (BIMA), while non was found toward core H-3a. Also in the PdBI data of Pillai et al. (2011) core H-3a does not coincide with any compact continuum emission, while core H-3b coincides with their PdBI mm1 core. Due to the complete absence of continuum emission toward core H-3a we do not consider it as a star-forming object, and we attribute all the continuum emission to H-3b.

To obtain masses and temperatures we fitted a single-temperature modified black body function to all the continuum fluxes assuming flux uncertainties of 20%. The grain emissivity parameter $\beta$ was allowed to vary between 1 and 2.5 (Sadavoy et al. 2013) for clumps/cores with ancillary millimetric data and kept constant at 2.0 otherwise. Errors in temperature, mass, $\beta$ and luminosity were estimated by varying the continuum fluxes ($\pm$20%) and estimating how much these quantities changed. The clump and core SEDs are shown in Fig. 4 with the best modified black body fit, while the resulting masses, temperatures and $\beta$'s are given in Table 2. From the fit to the SED of clumps H-1 and H-3 (Fig. 4) one can see the that the scaled Herschel and JCMT (850 μm) fluxes match well with unscaled the IRAM 30m data points justifying the flux scaling on the 160 μm size. Thanks to the SCUBA 450 μm data we could confine the clump H-2/core H-2a SED and obtain relatively good estimates of masses, temperatures and luminosities. However, because this source was lacking millimetre detections we kept $\beta = 2.0$. We find that clumps H-1 and H-2 are warmer and less massive than clump H-3. Rescaling the clump fluxes with the core sizes did not significantly affect the temperatures, but strongly reduced the envelope masses.

3.1.2 Millimetre emission and free-free contamination

Close to H ii regions, free-free emission will be present in addition to thermal dust emission. Free-free emission dominates at centimetre frequencies, and contributes to the thermal emission at millimetre and sub-millimetre wavelengths. Using the radio free-free emission of W 48A UC H ii region reported by Roshi et al. (2005), we extrapolated the free-free contribution at 850 μm (10% - 1.76 Jy) and 1.2 mm (32% - 1.80 Jy) and corrected the clump H-1 fluxes at these wavelengths, which were used for SED fitting, accordingly. In the 2.8 mm map (Fig. 2), the UC H ii region is clearly visible as a bright source with some weak extended emission flowing southwards, which is thought to be associated with the compact centimetre component identified by Kurtz & Francq (2002). Toward clump H-2, we found very weak 2.8 mm emission, at about a 2.5σ level given the image noise of 3 mJy beam$^{-1}$, not significant enough for our selection criteria. The millimetre emission toward clump H-3 is a bit stronger (about 3–4σ) and its shape correlates well the cold Herschel dust continuum emission (Fig. 2), suggesting that it is real. The extrapolation of the modified black body fit performed on the extended emission in the Herschel bands would predict about 16 mJy per 11.5 pixel, while the BIMA map contains about 20 mJy per 11.5 pixel in that area. We conclude that most of the extended 2.8 mm flux around clump H-3 originates in thermal dust emission.

3.1.3 Bolometric luminosities and age estimations

Objects with embedded YSOs will emit strongly not only in the wavelengths $\gtrsim$70 μm, but also in the mid and near infrared. To obtain a realistic estimate of the bolometric luminosity, we integrated the fluxes between 3 μm and 1.3 mm, using our obtained flux measurements complemented with, when available, the mid-infrared (3.4–22 μm) fluxes from the WISE catalog (Wright et al. 2010). All data points used to obtain the bolometric luminosity are plotted in Fig. 4 and the resulting luminosity $L_{bol}$ is given in Table 2.

For clump H-1, which contains the UC H ii region, we estimated the stellar type based on the bolometric luminosity, under the assumption that the luminosity is dominated by the most massive object, using the Zero-Age Main-Sequence (ZAMS) column of Table 1 of Panagia (1973). With respect to the previous, lower resolution IRAS observations (Wood & Churchwell 1989), we find a lower bolometric luminosity and a slightly later spectral type (O7.5 with respect to the previous O6) due to a lower level of confusion. Using the more recent O star parameters of Martins et al. (2005) the spectral type becomes O8. Our spectral type result, O7.5 or O8, reconciles the infrared-based spectral type with the one determined from radio properties – O7.5 (Wood & Churchwell 1989) and O8 (Roshi et al. 2005).

Bolometric luminosities and envelope masses are particularly important quantities, because they allow us to assign an age estimation to the object that is forming in the clump or core following the
evolutionary tracks from Saraceno et al. (1996) for low-mass stars and Molinari et al. (2008) for high-mass stars. In the \( L/M \) diagram, stellar evolution is represented in two main phases: the stellar accretion phase and the envelope clean-up phase. The first phase is when the forming star is accreting matter from the envelope, which increases its luminosity. Mass accretion stops when the star begins to burn hydrogen, i.e. when it arrives on the ZAMS. During the second phase the protostellar envelope is dispersing, while the luminosity stays roughly constant (within the same order of magnitude).

The stellar evolution tracks in \( L/M \) diagram assume the formation of single objects. A rough idea of the impact of multiplicity can be made invoking the Salpeter's initial mass function:

\[
N = k_i M^{-1.35},
\]

where \( N \) is the number of objects, and \( M \) is the stellar mass. It shows that most of the mass is in lower-mass objects. When the luminosity is determined by the stellar luminosity, it can be related to the mass by \( L \propto M^{3.5} \) and it will dominated by the more massive objects. Hence the envelope mass will be much more affected by the multiplicity than the luminosity.

Figure 5 shows the location of clumps H-1, H-2, H-3, and cores H-2a and H-3b on the \( L/M \) diagram. We estimated the age range from the two stellar tracks that surround the clump or core, for example for clump H-1 we used the 18 and 28 \( M_\odot \) stellar tracks. The obtained age ranges and average ages are listed in Table 2. The difference between using the masses and luminosities of cores rather than the clumps is that cores are less massive and slightly less luminous. This puts the cores on less massive evolutionary tracks which are evolving slower, and hence the ages become larger. Since the \( L/M \) diagram assumes the objects to be single, the core ages should be more accurate than those of the clumps and in the rest of the paper we will use the ages of cores H-2a and H-3b, rather than those of clumps H-2 and H-3. For clump H-3 to core H-3b the evolutionary stage did not change very much: both objects are in the envelope accretion phase and their ages are similar. The age difference between clump H-2 and core H-2a in the \( L/M \) diagram is a little larger due to the more complex behaviour of the stellar tracks in the \( L/M \) diagram near the ZAMS line. In the following sections, we present molecular line data to find further relative and absolute age estimates to compare with the \( L/M \) based ages.

### 3.2 Class II methanol masers

We detected the class II \( 3_1 - 4_0, A' \) methanol maser transition at 107 GHz toward clump H-2/core H-2a. This is the first interferometric observation of this maser transition in W 48A delivering sub-arcsecond images of the maser spots and their distribution, as previous 107 GHz detections were single-dish only (Val'tts et al. 1995). The distribution of 107 GHz methanol masers spots as well as their correlation with the 6.7 GHz and 12.2 GHz methanol masers (Minier et al. 2000) is discussed in Appendix B. We observed core H-2a and core H-3b on the \( L/M \) diagram near the ZAMS line. In the following sections, we present molecular line data to find further relative and absolute age estimates to compare with the \( L/M \) based ages.

### 3.3 Methyl cyanide (\( \text{CH}_3\text{CN} \)) emission

\( \text{CH}_3\text{CN} \) is a commonly observed molecule in hot molecular cores, which are regions of high-density molecular gas surrounding an intermediate to high-mass YSO. The YSO elevates the dust temperatures (> 100 K) in these regions, which causes the evaporation of the icy grain mantles and the increase of the \( \text{CH}_3\text{CN} \) gas phase abundance by several orders of magnitude (Olmi et al. 1996; van der Tak et al. 2000; Purcell et al. 2000). The \( \text{CH}_3\text{CN} \) observations had a high angular resolution (see Table 1) allowing us to observe core-sized (<0.1pc radius) objects and to find the first signs of fragmentation in the clumps detected with Herschel. The integrated intensity \( \text{CH}_3\text{CN} \) map in Fig. 3 shows that \( \text{CH}_3\text{CN} \) is concentrated toward three cores, whose continuum properties are described in Section 3.1. The most compact core H-2a, which appears to be unresolved, overlaps with clump H-2, while clump H-3 was found to fragment into two \( \text{CH}_3\text{CN} \) cores, H-3a (no continuum emission detected) and H-3b (continuum emission detected). It is

---

**Table 2. Cold dust continuum properties of W 48A clumps and cores**

<table>
<thead>
<tr>
<th>Object</th>
<th>R.A. (J2000) (h:m:s)</th>
<th>Dec. (J2000) (°:'&quot;)</th>
<th>( R_{30} ) pc</th>
<th>( R_{\text{CH}_3\text{CN}} ) pc</th>
<th>( M_{\text{env}} ) (( M_\odot ))</th>
<th>( T_{\text{dust}} ) (K)</th>
<th>( \beta )</th>
<th>( L_{\text{bol}} ) (L_\odot)</th>
<th>Age (10^5 yr)</th>
<th>Age range (10^5 yr)</th>
</tr>
</thead>
<tbody>
<tr>
<td>clump H-1</td>
<td>19:01:46.6</td>
<td>01:13:25</td>
<td>0.10</td>
<td>720±250</td>
<td>37±4</td>
<td>1.5±0.1</td>
<td>77200±15000</td>
<td>7.5</td>
<td>2–13</td>
<td></td>
</tr>
<tr>
<td>H-2</td>
<td>19:01:45.7</td>
<td>01:13:33</td>
<td>0.11</td>
<td>240±60</td>
<td>27±4</td>
<td>2.0±0.5</td>
<td>11200±500</td>
<td>15.1</td>
<td>2.4–12</td>
<td></td>
</tr>
<tr>
<td>H-3</td>
<td>19:01:42.4</td>
<td>01:13:33</td>
<td>0.13</td>
<td>1100±600</td>
<td>21±4</td>
<td>1.7±0.2</td>
<td>4000±1500</td>
<td>1.2</td>
<td>0.9–1.5</td>
<td></td>
</tr>
<tr>
<td>core H-2</td>
<td>19:01:45.5</td>
<td>01:13:33</td>
<td>0.05</td>
<td>170±30</td>
<td>27±9</td>
<td>2.0±0.5</td>
<td>8000±1000</td>
<td>8.0</td>
<td>6.8–9.5</td>
<td></td>
</tr>
<tr>
<td>core H-3</td>
<td>19:01:42.3</td>
<td>01:13:33</td>
<td>0.08</td>
<td>420±200</td>
<td>19±4</td>
<td>1.9±0.3</td>
<td>2700±600</td>
<td>1.5</td>
<td>1.0–2.0</td>
<td></td>
</tr>
</tbody>
</table>

**NOTES.** Columns are (from left to right): object name; positions of clumps (continuum) and cores (\( \text{CH}_3\text{CN} \)); clump radius at 160\( \mu \)m; radius of \( \text{CH}_3\text{CN} \) core; envelope mass; dust temperature; grain emissivity parameter; bolometric luminosity; \( L/M \)-based average age estimation and the range of ages from the two nearest star formation tracks. Radius is the half of the full width at half maximum of the Gaussian fit performed with CuTeX on the 160\( \mu \)m map and deconvolved from the beam at 160\( \mu \)m, 12\". For clump H-2 and core H-2a the emissivity index \( \beta \) was fixed at 2.0 because the SEDs had no millimetric data points.

**Table 3. Derived properties of the \( \text{CH}_3\text{CN} \) cores.**

<table>
<thead>
<tr>
<th>Core</th>
<th>R.A. (h:m:s)</th>
<th>Dec. (°:'&quot;)</th>
<th>R (pc)</th>
<th>( V_{\text{LSR}} ) (km s(^{-1}))</th>
<th>( \Delta v_{\text{obs}} ) (km s(^{-1}))</th>
<th>( T_{\text{rot}} ) (K)</th>
<th>( N_{\text{CH}_3\text{CN}} ) (10(^5) cm(^{-2}))</th>
</tr>
</thead>
<tbody>
<tr>
<td>H-2a</td>
<td>19:01:45.55</td>
<td>01:13:33</td>
<td>0.05</td>
<td>42.7±0.2</td>
<td>1.9±0.4</td>
<td>61±20</td>
<td>8±3</td>
</tr>
<tr>
<td>H-3a</td>
<td>19:01:42.65</td>
<td>01:13:37.5</td>
<td>0.06</td>
<td>42.1±0.1</td>
<td>2.2±0.4</td>
<td>44±4</td>
<td>9±6</td>
</tr>
<tr>
<td>H-3b</td>
<td>19:01:42.31</td>
<td>01:13:33.5</td>
<td>0.08</td>
<td>42.0±0.2</td>
<td>2.3±0.5</td>
<td>41±6</td>
<td>8±2</td>
</tr>
</tbody>
</table>

**NOTES.** Columns are (from left to right): core name; positions (J2000); radius of \( \text{CH}_3\text{CN} \) core; \( \text{CH}_3\text{CN} \) LSR velocity; \( \text{CH}_3\text{CN} \) line width of \( v > 100 \) K; \( \text{CH}_3\text{CN} \) rotational temperature; \( \text{CH}_3\text{CN} \) column density.
probable that the CH$_3$CN emission at the location of H-3a is extended emission from low velocity shocks, such as is seen in DR 21 (Csengeri et al. 2011) where it was associated with small-scale convergence flows, rather than the CH$_3$CN emission associated with a YSO. Such a nature of the CH$_3$CN emission is plausible, as core H-3a coincides with the red CO(2–1) outflow lobe, which could provide the low-velocity shocks (the outflow is discussed in Section 4.3). No CH$_3$CN emission was found toward clump H-1.

Spectra integrated over the area of the cores were extracted from the data cube and fitted with GILDAS/CLASS software. The $K=0$ to $K=3$ rotational ladder components were detected for all three cores and fitted simultaneously with Gaussian profiles. We constrained the line width of all $K$ levels to be the same, i.e. that of the $K=0$ component which is given in Table 2. The signal-to-noise ratio for the $K=4$ component was too low to obtain a reliable fit. The temperature and column density were estimated using the rotational diagram analysis method (see, e.g., Araya et al. 2005) assuming that the regions are optically thin (Turner 1991). In the upper panel of Fig. 5 we show an example of the CH$_3$CN $K = 0, 1, 2, 3$ component spectrum and its fit. The bottom panel shows the population diagram for the various $K$ levels on which the temperature and column density estimation was based. The positions of the cores and results of the fits are given in Table 2.

The optical depth can be calculated once the rotational temperature is known. If the CH$_3$CN emission is optically thick, the column densities will be underestimated, while the rotational temperatures will be over estimated (Goldsmith & Langer 1999). We derived optical depths of the order of 0.02. Iterating by the optical depth correction factor ($\tau/(1 - \exp(-\tau))$, as done by Remijan et al. (2004), we estimated that the corrected column densities and rotation temperatures are the same as those for the optically thin limit within the uncertainties. The calculated CH$_3$CN column densities of $\sim 8 \times 10^{12}$ cm$^{-2}$ (Table 3) are similar to those found in hot cores (Olmi et al. 1993; Araya et al. 2005). The rotation temperatures show that core H-2a is warmer than cores H-3a and b, and in addition is also more compact (Table 3). The CH$_3$CN line thermal widths are all dominated by non-thermal components (thermal CH$_3$CN line widths for 40–60 K are $\sim 0.2$–$0.3$ km s$^{-1}$).

3.4 Ammonia (NH$_3$) emission

The NH$_3$ (1,1) and (2,2) transitions are excited in dense and cold environments (e.g., Pillai et al. 2006). We analysed the NH$_3$ data cubes with the GILDAS software fitting the hyperfine structure of the NH$_3$(1,1) emission and using a gaussian for the NH$_3$(2,2) emission. The NH$_3$ spectra were extracted in the same areas as defined by the continuum clumps and the CH$_3$CN cores enabling a direct comparison between NH$_3$, dust, and CH$_3$CN emission. Following Ryle et al. (2010) we then derived NH$_3$ rotational temperatures and NH$_3$ column densities of the clumps and cores (see Table 4). As the column densities were calculated based on interferometric observations, we might filter out some of the NH$_3$ emission, hence the NH$_3$ column densities might be higher than the values we found.

Figure 7 shows the NH$_3$ spectra of the clumps. Towards clump H-1 we found NH$_3$ absorption, similar to that toward DR 21 (Cyganowski et al. 2003), indicating that the emission is emanating from cool molecular material in front of the 10$^5$ K hot UC H$_2$ region. The NH$_3$ rotational temperature of clump H-1, 17.6 K was much lower than the dust temperature (42 K), which is a possible indication of the NH$_3$ being more shielded from the UC H$_2$ region radiation than the dust. NH$_3$ emission toward clump H-2 was found to have two components (around 43 km s$^{-1}$ and 45.6 km s$^{-1}$), making the fitting of the lines more difficult. We fixed the optical depth (and its error) at 1.2±0.5 for both components when performing the line fit. The two velocity components are also found in the J=0 spectra (see Section 4.2.1), where the 43 km s$^{-1}$ component is correlated with the ridge emission, and the 45.6 km s$^{-1}$ component is thought to emanate from the arc-like structure around the UC H$_2$ region. Unfortunately, the temperature estimations have too large uncertainties to provide more detailed information on these two velocity components.

Clump H-3 had the strongest NH$_3$ emission among the three clumps. Its NH$_3$ rotational temperature (22 K) was found to be in very good agreement with the dust temperature (20 K). The NH$_3$ properties of clump H-3, which overlaps with Pillai et al. (2011) NH$_3$ core no. 3, are in agreement with the rotational temperatures and NH$_3$ column densities derived by those authors. The CH$_3$CN rotational temperatures were higher than the NH$_3$ temperatures, which is expected since CH$_3$CN traces warm regions near the YSO, while NH$_3$(1,1) and (2,2) trace colder gas. The NH$_3$ analysis of core H-2a yielded similar numbers as for clump H-2, except that for the NH$_3$(1,1) line we did not detected the 45.6 km s$^{-1}$ component due a higher noise in the core spectrum. Similarly, core H-3a and

Figure 6. Upper panels show the spectrum of the CH$_3$CN $K=0,1,2,3$ transitions and the fit (in green), while the lower panels show the CH$_3$CN rotational diagram for cores H-2a, H-3a, and H-3b. The rest frequency has been set on the $K=1$ level.

[Diagram of CH$_3$CN rotational diagram for cores H-2a, H-3a, and H-3b.]

8 http://www.iram.fr/IRAMFR/GILDAS
Table 4. Ammonia properties of clumps and cores

<table>
<thead>
<tr>
<th>Object</th>
<th>$T_{\text{mb}}$ (K)</th>
<th>$\Delta v_{11}$ (km s$^{-1}$)</th>
<th>$V_{\text{LSR}}^{11}$ (km s$^{-1}$)</th>
<th>$\tau_{11}^{11}$</th>
<th>$T_{\text{mb}}^{22}$ (K)</th>
<th>$\Delta v_{22}$ (km s$^{-1}$)</th>
<th>$V_{\text{LSR}}^{22}$ (km s$^{-1}$)</th>
<th>$T_{\text{rot}}$ (K)</th>
<th>$N_{\text{NH}_3}$ ($10^5$ cm$^{-2}$)</th>
</tr>
</thead>
<tbody>
<tr>
<td>clump H-1</td>
<td>-0.22±0.10</td>
<td>1.87±0.18</td>
<td>42.92±0.08</td>
<td>1.74±0.67</td>
<td>-0.11±0.02</td>
<td>2.18±0.37</td>
<td>43.04±0.17</td>
<td>18±7</td>
<td>3.9±2.2</td>
</tr>
<tr>
<td>clump H-2</td>
<td>0.19±0.08</td>
<td>1.34±0.16</td>
<td>42.99±0.08</td>
<td>1.30±0.5</td>
<td>0.18±0.03</td>
<td>2.00±0.26</td>
<td>42.84±0.11</td>
<td>40±4</td>
<td>-</td>
</tr>
<tr>
<td>clump H-3</td>
<td>0.22±0.09</td>
<td>1.26±0.13</td>
<td>45.60±0.06</td>
<td>1.3±0.50</td>
<td>0.14±0.03</td>
<td>1.74±0.30</td>
<td>45.47±0.13</td>
<td>24±13</td>
<td>-</td>
</tr>
<tr>
<td>core H-2a</td>
<td>0.65±0.07</td>
<td>1.84±0.06</td>
<td>42.23±0.02</td>
<td>2.47±0.24</td>
<td>0.47±0.03</td>
<td>2.26±0.10</td>
<td>42.17±0.04</td>
<td>22±3</td>
<td>6.3±1.1</td>
</tr>
<tr>
<td>core H-2b</td>
<td>1.17±0.57</td>
<td>1.49±0.28</td>
<td>42.87±0.10</td>
<td>2.22±0.80</td>
<td>0.24±0.05</td>
<td>1.73±0.27</td>
<td>42.78±0.10</td>
<td>-</td>
<td>-</td>
</tr>
<tr>
<td>core H-3a</td>
<td>0.72±0.08</td>
<td>1.68±0.05</td>
<td>42.17±0.02</td>
<td>2.47±0.22</td>
<td>0.52±0.03</td>
<td>2.13±0.10</td>
<td>42.13±0.04</td>
<td>22±3</td>
<td>6±1.0</td>
</tr>
<tr>
<td>core H-3b</td>
<td>0.75±0.09</td>
<td>1.79±0.06</td>
<td>42.18±0.03</td>
<td>2.81±0.28</td>
<td>0.58±0.04</td>
<td>2.25±0.11</td>
<td>42.15±0.04</td>
<td>23±3</td>
<td>7±1.4</td>
</tr>
</tbody>
</table>

NOTES. Columns are (from left to right): name of object; NH$_3$ (1,1) main beam brightness temperature; NH$_3$ (1,1) LSR velocity; NH$_3$ (1,1) main group opacity; NH$_3$ (2,2) main beam brightness temperature; NH$_3$ (2,2) line width; NH$_3$ (2,2) LSR velocity; NH$_3$ rotational temperature; NH$_3$ column density.

3.5 Outflows

Molecular outflows are commonly observed in low-mass and high-mass star-forming regions (see, e.g., Phillips et al. [1988], Beuther et al. [2002], López-Sepulcre et al. [2009], Rygl et al. [2013], Duarte-Cabral et al. [2013]). Outflows arise from the mass accretion onto the molecular clump (e.g., Shepherd & Churchwell [1996]), and are often accompanied by maser emission due to the shocks created by the outflow (e.g., see the water masers in the outflow of W3(OH) in Hachisuka et al. [2006]). Molecular outflows can be observed through optically thick emission of common ISM constituents, such as $^{13}$CO.

The HARP CO (3–2) data revealed an outflow in clump H-3. In Fig. 8 we show the integrated emission of the high-velocity CO (1–0) outflow and the 3 mm continuum. The CO (3–2) and the continuum are not symmetrical around the continuum source. The displacement uncertainties of the JCMT CO data (2′′) and BIMA CH$_3$CN (0′′) are much smaller than those of the continuum (1–1.2 km s$^{-1}$). Possibly, the CH$_3$CN emission emanates from a smaller area near the YSO, while the NH$_3$ emission comes from a larger volume containing more turbulence in the surrounding medium by possible disks and outflows (see the outflow in clump H-3 in Section 3.5).

Figure 7. NH$_3$ (1,1) and (2,2) spectra of clumps H-1, H-2, and H-3. The left-side y-axis gives the main beam temperature scale for the clump H-1 and H-2 spectra, while the right-side y-axis gives the main beam temperature scale for the clump H-3 spectra. The NH$_3$ (2,2) spectra have been offset artificially.

b have similar NH$_3$ properties as clump H-3. The LSR velocities of the NH$_3$ toward the cores H-2a, H-3a, and H-3b match well with the CH$_3$CN velocities, indicating that both molecules trace material with a similar kinematic signature, belonging to the same object.

While the NH$_3$ (1,1) line width is an intrinsic line width (since we fitted the hyperfine structure of the transition), the NH$_3$ (2,2) width is not: it is an overestimate of the intrinsic line width. For both NH$_3$ and CH$_3$CN, the line widths are dominated by the thermal line width, since thermal line widths for NH$_3$ are of the order of 0.2–0.3 km s$^{-1}$ for temperatures of 20–30 K. For all cores, the NH$_3$ (1,1) and (2,2) line widths were both much broader than those of CH$_3$CN (~1.1–1.2 km s$^{-1}$). Possibly, the CH$_3$CN emission emanates from a smaller area near the YSO, while the NH$_3$ emission comes from a larger volume containing more turbulence in the surrounding medium by possible disks and outflows (see the outflow in clump H-3).

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Herschel and are often accompanied by maser emission due to the shocks created by the outflow (e.g., see the water masers in the outflow of W3(OH) in Hachisuka et al. [2006]). Molecular outflows can be observed through optically thick emission of common ISM constituents, such as $^{13}$CO.

The HARP CO (3–2) data revealed an outflow in clump H-3. In Fig. 8 we show the integrated emission of the high-velocity CO. The blue and red high-velocity lobes are quite compact indicating that the outflowing material is collimated, and that clump H-3 must be hosting active star-formation activity. A comparison with 70 μm emission, tracing mainly warm dust near YSOs, shows that core H-3b (rather CH$_3$CN core H-3a) is the more likely origin for the outflow since it coincides with the 70 μm source. The CO (3–2) outflow lobes coincide with the lower resolution HCO$^+$ (1–0) outflow lobes discussed by Pillai et al. [2011] (see Fig. 8). The displacement between the CO outflow lobes and the Herschel dust continuum is real, since it is greater than the sum of the position uncertainties of the JCMT CO data (2′′) and BIMA CH$_3$CN data (0′′). Also the HCO$^+$ (1–0) outflow and the 3 mm continuum emission in Pillai et al. [2011] show the same displacement. Apparently the outflow axis is not straight, which is why the blue and red lobe are not symmetrical around the continuum source. Pillai et al. [2011] estimated the outflow mass to be about 5 $M_\odot$, indicating that core H-3b is driving a massive outflow. In addition to this massive molecular outflow, clump H-3 also hosts water masers (Hofner & Churchwell [1996]) located toward the red outflow lobe.
Figure 4. Spectral energy distributions of clumps H-1, H-2 and H-3 and cores H-2a and H-3b. Downward facing triangles indicate 3σ upper limits. The grey-body fit to the cold dust component is shown by a blue line. The estimated masses, temperatures, dust emissivity, and bolometric luminosities are listed in Table 2. Data points from the various observatories are colour coded: cyan squares for WISE, black filled dots for Herschel, red crosses for JCMT, and violet filled triangles for IRAM 30m.

The detection of the massive outflow and water masers indicates that clump H-3 is in an active star-formation stage, hence presumably younger than the other two clumps H-1 and H-2, which do not exhibit any collimated outflows.

4 RESULTS AND ANALYSIS OF EXTENDED EMISSION

4.1 Herschel continuum emission

We computed column density and temperature maps from the calibrated 160–500 μm images, after these were convolved to the 500 μm resolution (11.5′) and rebinned to the pixel size of the 500 μm map (11.5′). Then, pixel-by-pixel modified black body fits were performed on the regridded maps (see Elia et al. 2013 for more the expression and its details). We assumed a dust opacity $\kappa_\nu = \kappa_0 (\nu/\nu_0)^\beta$ with $\kappa_0=0.1 \text{ cm}^2\text{g}^{-1}$ (which includes a gas to dust ratio of 100) and $\nu_0=1200 \text{ GHz}$ (250 μm) (Hildebrand 1983). The dust emissivity index $\beta$ was kept constant at 2.0. Recent studies (e.g., Juvela et al. 2011; Miettinen et al. 2012; Sadavoy et al. 2013) show that $\beta$ varies across the cloud. While some works point out the anti-correlation between dust temperature and $\beta$, Bayesian SED-fitting finds an anti-correlation of $\beta$ with column density rather than temperature (Kelly et al. 2012). Ysard et al. (2013) find that the change in $\beta$ is due to dust-properties which vary with density. With our Herschel-only data (the SCUBA, MAMBO and BIMA dust continuum maps had a much smaller coverage, so were not used for the large-scale temperature and column density maps) are not able to constrain the $\beta$. Using $\beta = 2.0$ with $\chi^2$ fitting (as we do) can cause the resulting temperature to be too low in case $\beta < 2.0$ and vice versa, as shown by Sadavoy et al. (2013). Furthermore, to obtain the H$_2$ column density we used the mean molecular weight per hydrogen atom $\mu_H=2.8$ (Kaufmann et al. 2008). Recent HOBYS studies, such as e.g. Hill et al. (2013), often use a mean molecular weight of 2.3 which is averaged per particle instead per hydrogen atom and results in higher column densities by a factor of 2.8/2.3 ≈ 1.2.

Figure 5 displays the resulting temperature and column density maps of W48A with the 250 μm contours overplotted. As expected, the 250 μm contours match very well with the column density map, but are quite different from the temperature map, which traces the dust heated by YSOs and stars. The column density map shows: a north-west oriented arc around the UC H II region.
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Figure 8. Top panel: 70µm PACS greyscale map centred on clump H-3 with the JCMT CO (3–2) outflow (blue, red) and the CH$_3$CN cores (white) overlaid in contours. CO contour levels are 20, 30, 40, and 50 K km s$^{-1}$ for both blue and red shifted material. BIMA CH$_3$CN contour levels are 0.06, 0.08, 0.10, 0.12 Jy beam$^{-1}$ km s$^{-1}$. Beam sizes of the CO (3–2) (white), PACS 70µm (grey), and CH$_3$CN emission (black) are given in bottom right corner. The green triangle represents the water maser location (Hofner & Churchwell 1996), the blue and red crosses mark the HCO$^+$ outflow lobes from Pillai et al. (2011). Bottom panel: the CO(3–2) spectrum toward clump H-3 where the velocity ranges used for mapping of the blue and red shifted outflow lobes are indicated.

The cold streamers seen by Herschel were first noted by Pillai et al. (2011), who found several NH$_2$D cores just westward of clump H-3 using interferometric PdBI data. To deuterate NH$_3$, creating the NH$_2$D molecule, very cold temperatures (T<15 K, Bergin & Tafalla 2007) are required. Protostellar heating of the environment, which would increase the temperature T>20 K would destroy the NH$_2$D. Therefore the detection of NH$_2$D cores indicates that in this region no protostars have been formed yet. These NH$_2$D cores were not detected in the Herschel maps, nor in the high resolution 1.3 mm PdBI data of Pillai et al. (2011), showing that at small core-like scales the deuterated gas differs from the dust distribution, while on larger scales (as 1 pc, such as the streamers) the correlation between the dust and NH$_2$D emission is remarkably good. The Herschel dust temperature and column density maps show that these streamers have indeed the lowest dust temperature of the whole W 48A environment and that they are dense structures. We conclude that the cold streamers represent the earliest star-formation stage observed in the W 48A star-forming region.

The word ridge is used to indicate an filamentary cloud with a column density > 10$^{23}$ cm$^{-2}$ (see Hill et al. 2011).
Figure 10. **Left panel:** BIMA+IRAM 30m integrated intensity map of the $^{18}$O emission. Clump H-1 (W 48A UC H\(_n\) region), core H-2a (maser source and clump H-2), clump H-3, the ridge, and the arc are marked. Contours are 0.25, 0.5, 1.0, 1.5, 2.0 Jy beam\(^{-1}\) km s\(^{-1}\). **Right panel:** Position velocity plot along the red dash-dot line in the left figure. Contours are 0.15 0.25 0.35 0.45 0.55 0.65 0.75 mJy beam\(^{-1}\). The various velocity components are marked on the plot.

Figure 11. IRAM 30m $^{18}$O (1–0) channel map smoothed to a channel spacing of 0.64 km s\(^{-1}\) for a better visualisation. Central velocity of each channel is marked in the top left corner. The map noise is 0.13 K and the contour spacing is 0.3 K (2.3 \(\sigma\)). Black squares indicate clumps H-1, H-2 and H-3.

### 4.2 Kinematics

#### 4.2.1 $^{18}$O (1–0) gas kinematics near the UC H\(_n\) region

We proceed with the exploration of the kinematics of the extended structures of the W 48A star-forming region through the $^{18}$O (1–0) data. The BIMA+IRAM30m $^{18}$O (1–0) integrated intensity map (Fig. 10 left panel) shows a large area of emission around the head of the arc and a horizontal extension (the ridge) with several bends, that ends near the H-3 clump. The peak of the $^{18}$O (1–0) emission is near the H\(_n\) region. To investigate the gas kinematics, we made a $^{18}$O position-velocity plot (right panel of Fig. 10) to show the $^{18}$O (1–0) intensities at each velocity bin along the cut across the arc and the ridge (indicated in the left panel). The ridge has a quite simple velocity structure: a very narrow velocity range (42 < \(V_{\text{LSR}}\) < 43 km s\(^{-1}\)) including the CH\(_3\)CN LSR velocities along the ridge and is considered to be “at rest”. Near the UC H\(_n\) region the kinematics become more interesting. There is receding arc at angular offsets < 40\({}'\), velocities 42–45 km s\(^{-1}\) and some material in front of the UC H\(_n\) region (at 40\({}'\)). Together they might create a shell structure that is similar to the expanding bubbles in the model of Arce et al. (2011). However, unlike in Arce et al. (2011) our source is not a bubble, but an arc or a shell. Also there is more material redshifted than blueshifted, indicating that we are not dealing with a symmetrical structure, and making the word “arc” more appropriate than “shell”. Fig. 10 suggests that the UC H\(_n\) region formed on the near side of the ridge, and during its evolution pushed most of the dense gas away to be redshifted. A smaller amount of material surrounding the UC H\(_n\) region ended up blueshifted or at rest.

To visualise the multiple velocity components we use the
IRAM 30m 18O(1–0) channel map (Fig. 11) as the single dish 18O data covered a larger area (the 18O BIMA data were limited by the field of view). The 18O emission seems to be a mixture of three main velocity components, one around 38–41 km s⁻¹ (which we name the “blue” component, see bottom four panels in Fig. 11), one around 41–43.5 km s⁻¹ (the “green” component, the middle four panels in Fig. 11), and one around 44–45 km s⁻¹ (the “red” component, top four panels in Fig. 11). The arc is clearly visible: starting south-east of clump H-1 (UC Hα region) at vLSR=42.6 km s⁻¹ and proceeding northwards toward clump H-1 and from there north-eastwards at vLSR=44.5 km s⁻¹, covering a velocity gradient of ~2 km s⁻¹. The green velocity component contains most of the 18O emission: it correlates well with the dust continuum and the column density map: the ridge, the streamers, and the middle-southern part of the arc are all visible. This is surprising since this velocity range also includes the LSR velocities of the clumps and cores. What looked one arc in the dust continuum, seems to kinetically be more complicated: the middle-northern part of the arc is receding, and the southern part of the arc is at rest or advancing. This points either to two uncorrelated velocity structures or to an arc which is inclined to result in the observed perspective. We obtained the clump 18O spectra from the high-dynamic range BIMA+IRAM 30m data and fitted a Gaussian to each velocity component (Fig. 12). The green velocity component is found toward all three clumps, while the red component is found only toward H-1 and H-2. The blue component, seen toward clump H-1 and H-3, is always much weaker than the green or the red component.

4.2.2 13CO(1–0) emission tracing material in front of the UC Hα region

This blue 18O velocity component, introduced in the previous section, is different from the red and green 18O components. Namely, while the red and green 18O have a similar distribution to the dust emission, the blue 18O does not. When one compares the vLSR = 38.79 km s⁻¹ channel map in Fig. 11 with the column density map in Fig. 11, one can notice that none of the high-density structures are coinciding with the blue 18O emission. In fact, the blue velocity component correlates well with the Ks-band absorption features, which are caused by material in front of the UC Hα region (Fig. 12).

Our BIMA dataset included 13CO observations. As 13CO is about ten times more abundant in cold molecular clouds than 18O (see e.g., Du & Yang 2008 and references therein), it is often observed to be optically thick, even when 18O is still optically thin. Therefore, it can be used to visualise the location of low density molecular gas. If there was optically thick molecular gas in front of the UC Hα region, the 13CO emission would have the same shape as the near-infrared absorption, as we indeed see in Fig. 13. The correlation of the blue 18O component, the Ks absorption and the 13CO emission suggests that it traces low density material in front of the UC Hα region.

5 DISCUSSION

5.1 W 48A revisited

Roshi et al. (2005) sketched the W 48A UC Hα region as a PDR dominated Hα region, where the ionising star is moving with respect to the molecular cloud. Their C76α radio recombination line (RRL) measurements showed that the PDR (vLSR=41.9 kms) was moving at 2.5±1 km s⁻¹ into the molecular cloud for which they took VLSR=44 km s⁻¹ (Churchwell et al. 1992). The LSR velocities of our 18O (1–0) data show that the molecular cloud covers a range of velocities from ~41 up to 45 km s⁻¹. The LSR velocity of the C76α line falls in the velocity range of the molecular material: the PDR is moving at the same velocity as that part of the cloud.

To understand if the PDR is pressure confined by the molecular material, we use the formulae given in Roshi et al. (2005) and calculate pressure of the molecular cloud. The molecular pressure is given by the sum of the pressure from thermal processes and from turbulence (Xie et al. 1996). Roshi et al. (2005) demonstrated that this Pmol = 5.3×43.0×10⁻⁷ dyne cm⁻² > PUCCHα = 1.3×10⁻⁷ dyne cm⁻² and that therefore the UC Hα region is pressure confined by the PDR. Now, the pressure in the molecular material in the ridge is given by

Pmol = nH₂(kₜ mol + μmHDv² mol),

where Tmol = 26 K (from the dust temperature map), μ = 2.8 and dv mol =1.0–1.8 km s⁻¹ based on 18O (1–0) and NH₃ (1,1) line widths resulting in Pmol = nH₂(1×10⁻¹³ dyne cm⁻²). Using the NH₃ rotational temperature (22 K) would yield a similar outcome. Estimates of the volume density based on the column density map depend on the depth of the ridge. Assuming the same depth as width (~1 pc), the volume density of the molecular material in the ridge would be 6×10⁴ cm⁻¹. This would bring the pressure of the molecular material to 6×10⁻⁷ dyne cm⁻², much lower than the PDR pressure. To stop the PDR from expanding into the molecular cloud, Pmol would have to be larger than PUCCHα which would require molecular hydrogen densities of 10⁴ cm⁻³. However, our hydrogen density estimate for the ridge is several orders of magnitude smaller. Therefore, the PDR is expanding into the cloud and pushing the molecular material outward, maintaining the arc-like structure. To fully appreciate the role of the PDR, high-sensitivity RRL imaging of the W 48A UC Hα region and its surroundings would be necessary to trace the PDR kinematics. Finally, Roshi et al. (2005) found that the UC Hα region is receding with a few (~3) km s⁻¹, which might explain why the arc is receding rather than advancing with...
In Section 3.1.1, we derived the ages of the clumps/cores based on the L/M tracks of Molinari et al. (2008) to be: \( -8 \times 10^4 \) yr for clump H-1 (UC Hn region) and core H-2b (maser core), and \( -1.5 \times 10^5 \) yr for core H-3b (containing the molecular outflow). These ages contain many uncertainties: the uncertainties of the envelope masses and bolometric luminosities, and those from stellar evolution models used for the stellar tracks. It is therefore useful to compare these age estimations with the results from radio continuum data, molecular line data, and the literature.

Similar to our age estimate: \( 6 \times 10^5 \) yr (Motte et al. 2007), \( ~ 10^5 \) yr (Wood & Churchwell 1989; Akeson & Carlstrom 1996).

Apart from water and hydroxyl masers, core H-2a contains several class II methanol maser transitions, indicating the presence of a high-mass star in formation (Breen et al. 2013). Methanol masers are thought to be excited before the onset of the Hn region and be quenched by the ionising radiation of the star (during the Hn region phase), and should hence have younger or similar ages as Hn regions, i.e., \( \leq 10^2 \) yr (Walsh et al. 1998; Codella & Moscadelli 2000; Breen et al. 2010). The detection of CH-CN, a molecule typical of intermediate to high-mass protostellar objects such as hot cores (e.g., Olmi et al. 1996; Araya et al. 2005; Purcell et al. 2006), and the absence of a compact radio source (due to the free-free emission of the ionising star, which is typical of Hn regions) also agrees with a younger evolutionary stage than a UC Hn region or a deeply embedded UC Hn region. Higher sensitivity radio observations could test if the object might possibly be a hyper compact Hn region.

Core H-3b was found to contain a high-mass young stellar object identified by the massive collimated outflow observed in CO (3–2) emission (Section 3.5) and HCO\(^+\) (1–0) line (Pillai et al. 2011). We detected NH\(_3\) and CH\(_3\)CN emission and obtained the [CH\(_3\)CN]/[NH\(_3\)] abundance ratio, which can be used as a chemical clock (Charnley et al. 1992). The abundance ratio found toward core H-3b was \( 1 \times 10^{-3} \) (see Tables 3 and 4), which is similar to the typical hot core abundance ratio of [CH\(_3\)CN]/[NH\(_3\)] \( \sim 10^{-3} \) that corresponds to ages of \( \sim 6 \times 10^4 \) yr (Charnley et al. 1992). Literature ages for hot cores (high-mass protostellar objects) are between \( 2 \times 10^4 \) yr (Motte et al. 2007) and \( 5 \times 10^4 \) yr (Cazaux et al. 2003). These two age estimates are both lower than what we obtained from the L/M diagram.

Using several molecular lines and radio continuum data, we could infer that: 1) clump H-1 is indeed the most evolved and oldest object, 2) core H-2a is in an earlier evolutionary stage than clump H-1 and most likely also younger than clump H-1, while the L/M diagram based age estimation found similar ages for these two objects, and 3) that core H-3b is indeed the youngest and less evolved object. Given that all three objects are forming high-mass stars the evolutionary difference is likely to reflect the age difference, since high-mass stars evolve along similar tracks, i.e. change their envelope mass and luminosity on similar timescales. It is thus important to have higher-resolution submm or radio continuum data and molecular line data to complement the age estimations based on infrared continuum data (such as the L/M diagram).
5.3 Sequential star formation sequence around W 48A?

The top panel in Fig. 15 shows a synthesis of our sequential star formation results in W 48A. In this study we isolated at least four different stages of star formation and estimated their ages, the UC H II region (clump H-1), the maser core (core H-2a), the active star-forming YSO with outflow (core H-3b), and the cold streamers in which Pillai et al. (2011) found prestellar cores in NH$_2$D of 50–200 $M_\odot$. The ages of prestellar objects depend strongly on whether they will form low- or high-mass stars. Assuming that these massive prestellar cores will form high-mass stars, then they should have younger ages than hot cores $< 5 \times 10^4$ yr (Cazaux et al. 2003). As protostellar heating destroys NH$_2$D, detecting these prestellar cores indicates that no protostars have been formed yet, which was confirmed through the non-detection of protostellar objects in the 70$\mu$m PACS map. Interestingly, these four stages of star formation are aligned in a east (old) to west (young) configuration, forming a linear age gradient (Fig. 15 panel (a)). Large projection effects can, most likely, be disregarded since the objects are forming within the same molecular cloud, suggesting that this remarkable age arrangement is real and reflects the formation history of the W 48 molecular cloud.

Panels (b) and (c) in Fig. 15 show that the UC H II region formed on the border of the molecular cloud. Core H-2a could have either formed before the creation of the arc and was swept up by it, or formed in the material collected by the PDR in a collect and collapse mechanism (Elmegreen & Lada 1977; Whitworth et al. 1994). It is unlikely that the UC H II region triggered the formation of the W 48A molecular cloud nor the star formation inside it. In the previous section we mentioned the UC H II region is confined by the PDR (Roshi et al. 2005), but that the PDR is slowly expanding into the molecular medium. The distance between UC H II region and clump H-3 is about 1 pc. A shock wave propagating at $\sim 0.3$ km s$^{-1}$, which is the sound speed of the medium at 25 K, would take about 3 Myr. The age clump H-3 is about an order of magnitude younger than that, ruling out the UC H II region (and the PDR) as the triggering agent. Hence, the W 48A molecular cloud and its star-forming seeds were assembled before the onset of the UC H II region in such a way that star formation started in the east and progressed toward the west. The sequential star formation in W 48A reminds of the Aquila supershell expansion is oriented in the direction of Galactic latitude, $b$, indicated in the bottom panel.
age gradient seen in the DR 21 molecular cloud which was formed by colliding flows, and has a south (old) - north (young) age gradient (see Csengeri et al. 2014, Henemann et al. 2012). Given the location of W 48 at almost 100 pc below the Galactic plane, the presence of an age gradient suggests a large scale external force.

First of all, the location of the other W 48 H"ii regions cannot explain the age gradient (by e.g., triggering) and it seems more likely that all these H"ii regions belong to one larger star-forming complex. Based on 21 cm HI line data, Maciejewski et al. (1996) discovered the Aquila supershell, centred at about four degrees below W 48A, at $l=35^\circ, b=-6^\circ$. This shell is expanding from below into the Galactic plane, creating, in addition to a superbubble, a cone-like shape directed at $l=34^\circ/6, b=+1^\circ 4$, very near ($\sim 40^\prime$ or $\sim 38$ pc) to the W 48 H"ii regions. The 21 cm HI spectra of the left and centre part of this cone show a clear peak at 40–45 km s$^{-1}$ (Figure 4, Maciejewski et al. 1996), which overlaps with the LSR velocity of the molecular material of W 48A and of the other H"ii regions in W 48: 41.2 km s$^{-1}$ for W 48B; 46.7 km s$^{-1}$ for W 48D; 45.5 km s$^{-1}$ for W 48E (Onello et al. 1994). It is therefore likely that the initial molecular cloud in which the W 48 complex of H"ii regions was formed through the compression of ambient material by the Aquila supershell. Maciejewski et al. (1996) estimated that the total energy of the events creating the Aquila supershell are about $1 \sim 5 \times 10^{52}$ erg, corresponding to 10–100 supernovae explosions powering the system over 10$^7$ yr. The W 48 H"ii regions were formed recently given the supershell’s age ($\sim$ estimated expansion velocity of about 15 km s$^{-1}$) and a distance between the centre of the supershell and the W 48 H"ii regions depending on the density of the material it passes through. The distance between the centre of the supershell and the W 48 H"ii regions is about 4 degrees (230 pc at a distance of 3.27 kpc) and a velocity of 15 km s$^{-1}$ it would take about 1.5 Myr for the shell to reach the W 48 H"ii regions, leaving ample time for the W 48 H"ii regions to evolve into their current state.

At the location of the W 48 H"ii regions, the supershell’s orientation is toward increasing Galactic latitudes. The shell would have therefore first reached the locations of W 48C and D, then W 48A and B, and finally W 48E. Hence, the density structures created by the shell, should be older in the eastern side of Fig. 15 and younger in the western side. Most of these H"ii regions are too extended than W 48D, but more extended that that of W 48A, implying an intermediate evolutionary stage. This is consistent with W 48B already containing a cluster of stars (clearly visible in the near-IR) and being surrounded by a small circular dust arc. Based on the 21 cm radio continuum we conclude that the H"ii regions have an evolutionary gradient along Galactic latitude, such as one would expect if these regions were formed by the Aquila supershell. In this discussion we have neglected W 48E, whose nature, either a H"ii region or a supernova remnant, is uncertain, since it emits extended weak radio emission, and does not contain many YSOs, nor high dust column densities.

In addition to the W 48 H"ii regions, IRAS 18586+0106, also known as Mol 87, is located between W 48A and W 48E. IRAS 18586 contains a few massive star-forming clumps (Beltr"an et al. 2006), but no centimetre continuum emission, which would indicate free-free emission of young stars, could be associated with them (Molinari et al. 1998), nor methanol or water masers (Fontani et al. 2010). With its NH$_3$ (1,1) LSR velocity of 38 km s$^{-1}$ (Molinari et al. 1996), IRAS 18586 is very likely to be a part of the W 48 complex. In Appendix D we analysed the infrared emission of the two clumps found in the Herschel maps. Source A, which coincides with the IRAS-coordinates centre, is the most evolved source with an age of $\sim5\times10^3$ yr forming an intermediate to high-mass star. The second source (B) is colder, more massive and younger. For the beginning of star formation in IRAS 18586, we only need to take the estimate of the oldest object into consideration, which is $\sim5\times10^3$ yr. This is younger than most evolved object in W48A (the UC H"ii region), and would agree with the predicted age gradient if all these star-forming regions were formed by the cone of the Aquila supershell.

In addition, the Herschel column density and temperature maps (Fig. 15) revealed that there exists a large reservoir of dense and cold gas west of W 48A, implying that this region has a high potential to form stars. The morphology of this dense gas suggests that it has not been strongly affected by H"ii regions or stellar winds, that are known to shape the material into bubbles and shells. East of W 48A, there is much less dense material, and when present it is shaped in shells, bubbles and ridges by H"ii regions. Hence, the large scale Herschel data support the hypothesis of an east-west evolutionary gradient across the W 48 H"ii regions, in a roughly similar orientation as the age gradient found around W 48A.

### 6 SUMMARY

We have presented a combined approach of high-sensitivity Herschel continuum observations with high-resolution molecular line data from the BIMA array to study the sequential star formation around the W 48A UC H"ii region. For the three clumps found with Herschel, which include a UC H"ii region, a maser emitting YSO, and a YSO with a massive molecular outflow, we combined age estimations based on envelope mass and luminosity with those from radio continuum, molecular line data and the literature values, and found the ages to agree within an order of magnitude. Confusion of emission and multiplicity affect the L/M-based ages, by moving the star-forming object to more massive stellar tracks and to higher luminosities. We tried to correct for this by rescaling the continuum...
fluxes of clump H-2 and H-3 to the core-sized objects detected in the BIMA CH3CN maps. High resolution continuum observations with ALMA, SMA or CCAT are necessary to reach the fragmentation limit of the cores for obtaining more accurate masses and luminosities. In addition to dust continuum data, radio continuum and molecular line data are crucial to identify the various evolutionary phases and to verify the ages based on envelope masses and luminosities.

The W 48A molecular cloud has a particular shape: it is concentrated along a ridge (the main axis of the cloud) lying in front of an arc (which, from our perspective is oriented toward the cloud) that partially surrounds the UC H II region. 180(1–0) line data shows that this arc-like structure has two velocity components, the northern part which is receding and the southern part which is at rest/advancing. This arc, or part of it, was most likely created by the UC HII region, when this formed at the eastern edge of the ridge, and is maintained by the over-pressured PDR. Based on the age dating of the Herschel sources, we found a linear evolutionary gradient across the W 48A star-forming region that we interpreted as an age gradient. This age gradient starts in the east of W 48A with the UC HII region, the oldest object, and traverses the arc following the major axis of the W 48A molecular cloud. Since the UC HII region is pressure confined by the PDR, it could not have formed the W 48A molecular cloud nor triggered the formation of most of the YSOs inside of it.

From the centimetre radio continuum we found that a similar east (older) - west (younger) orientation, as in W 48A is also present across the W 48 H II regions. Also, large scale dust maps show that in the eastern part of W48 the H II regions shaped all the gas and dust around them (more evolved, older), while in the western part we find mostly molecular clouds (less evolved, younger). We therefore considered even larger scales and found that the Aquila supershell expansion is a plausible candidate to have swept up the material to form the W 48 complex, since the velocity of the shell matches the LSR velocities of the W 48A-E HII regions, and the age gradient of the W 48 HII regions is roughly aligned with the direction of the expansion of the shell.

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REFERENCES

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APPENDIX A: MAMBO 1.3 MM MAP

The 1.3 mm MAMBO data of W 48A shown in Fig. 2 were part of a larger map presented in Fig. 11. Observational details are given in Section 2.3 and Table 1.

APPENDIX B: DISTRIBUTION OF 107 GHZ METHANOL MASERS

The spectra of the 107-GHz masers compared to the 12.2-GHz (M. Gaylard, priv. comm.) and 6.7-GHz masers (Goedhart et al. 2004) are shown in Fig. 11. While the 6.7- and 12.2-GHz masers share the same velocity range of 39.5–46 km s$^{-1}$, the 107 GHz masers show emission out to 48 km s$^{-1}$, indicating that various transitions probe similar, but not the same physical conditions. The relative intensities of individual maser features are different at each transition, with the red-shifted features brighter at 107 GHz (see also the work of Val’tts & Lyubchenko 2002).

The 107 GHz methanal maser spots were detected across a velocity range of $\sim 8$ km s$^{-1}$ (see Fig. 11) and appear to be unresolved in the data cubes generated. We obtained the positions of the methanol maser spots by fitting two-dimensional Gaussians to the data cube from the (highest resolution) BIMA A-array data. The beam-size from A-array data alone is 0″.003 pc or 685 AU. The spatial and velocity distribution of the 107 GHz methanol masers is shown in Fig. 12.

Methanol maser kinematics can trace the small scale gas motions near the protostar, such as envelope expansion or a rotating disk (Moscadelli et al. 2011). AFGL 5142, is an example of a high-mass protostar where the 3D kinematics of methanol masers seem to be emanating from a rotation disk (Goedhart et al. 2011). The general morphology and angular distribution of the 107-GHz masers is similar to that found for the 6.7- and 12.2 GHz masers by Minier et al. (2000): spots belonging to the red-shifted group in the north-east and spots belonging to the blue-shifted group in the south-west. However, a one-to-one correspondence of spots in the different transitions is not apparent. The 107 GHz emission appears to be coming from a smaller area than seen in Minier et al. (2000): ∼0″.01 compared to the ∼0″.43. The 107 GHz masers show a tighter correlation in position and velocities than the 6.7 and 12.2 GHz masers. We find a correlation coefficient of 0.86 by performing a linear fit to the 107 GHz maser distribution. If we assume Keplerian rotation and a distance of 3.27 kpc (Zhang et al. 2004), the mass enclosed in 0″.01 (33 AU) is derived to be $\sim 2.2 \times 10^3 M_\odot$ from the mass enclosed in 0″.01 (33 AU) is derived to be $\sim 2.2 \times 10^3 M_\odot$. Furthermore, if we interpret the spatial distribution of 6.7- and 12.2 GHz masers in Minier et al. (2000) as due to the inclination of the disk, we obtain an inclination of 30°. Applying this inclination, the enclosed mass increases to 3 $M_\odot$, which about a factor 2 lower than the mass obtained by Minier et al. (2000). Class II methanol masers, to which the 6.7, 12.2 and 107 GHz masers belong, are associated with high-mass star formation (Menten 1991; Minier et al. 2003; Breen et al. 2013). Our derived mass of 3 $M_\odot$ is unlikely to form a high-mass star, indicating that the 107 GHz masers do not allow a mass calculation, because, for example, they do not trace the entire disk, as also suggested recently by Breen et al. (2013).
Figure A1. Full MAMBO 1.3 mm continuum map. Apart from W 48A and IRAS 18586, other mm-bright star-forming regions in the field are indicated.

Figure B1. Spectra of 107 GHz (grey histogram), 12.2 GHz (dashed blue), and 6.7 GHz (green) methanol masers toward the clump H-2 in W 48A. The left y-axis is for the 107 and 12.2 GHz transitions, while the right y-axis belongs to the 6.7 GHz methanol maser.

Figure B2. 107 GHz methanol maser map obtained through gaussian fitting. Weaker features have higher positional uncertainties. Axes give the offset in arcseconds with respect to the pointing centre 19°01′45.448 +01°13′21″493 (J2000).

APPENDIX C: DISTRIBUTION OF $\text{NH}_3$ (2,2)/$\text{NH}_3$ (1,1) MAIN BEAM TEMPERATURES

The ratio of the $\text{NH}_3$ (2,2) to $\text{NH}_3$ (1,1) main beam temperature is a sign of increasing kinetic temperature [Torrelles et al. 1989].
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Figure C1. Ratio of NH$_3$(2,2) and NH$_3$(1,1) main beam temperatures. The beam is shown in the bottom right corner.

Zhang et al. (2002), and can be used to trace heating sources. Figure C1 shows the distribution of this ratio based on the VLA cubes. Near clump H-2 the ratio is high. This sign of heating is consistent with the presence of a PDR as seen in previous studies of PDRs near H$_{\text{ii}}$ regions (e.g., Palau et al. 2007). Note that the gas near clump H1 has a low ratio, which indicates that this gas is cold, and hence in front of the UC H$_{\text{ii}}$ region – this agrees with the NH$_3$ absorption seen at this location (see section 3.4).

APPENDIX D: IRAS 18586+0106

IRAS 18586+0106 originated from the IRAS-colour-selected list of low sources by Palla et al. (1991), which was consequently studied in NH$_3$, centimetre and submillimetre emission (Molinari et al. 1998, 1999, 2000, Beltrán et al. 2006). The object was found to have weak NH$_3$ and submillimetre emission, and no radio emission (the radio emission was found to be uncorrelated with the IRAS source. The IRAS resolutions at the various wavelengths (for example, the point spread function (PSF) at 60$\mu$m was $\sim$2$\times$5$''$) are nowadays surpassed by those of Herschel ($\sim$6$''$ at 70$\mu$m). Figure D1 shows the 70 and 250$\mu$m Herschel and 1.2 mm MAMBO maps of the IRAS source. Note that the image is completely within the IRAS PSF at 60$\mu$m, and that two clear clumps, marked as A and B are present. Of these, source A (which coincides with the centre of the IRAS source) is brighter at shorter wavelengths indicating a later evolutionary stage than that of source B, which is more luminous in the far-infrared to submillimetre. Performing the source extraction and SED fitting (Fig. D2, as described in Section 3.1.1) we find that source A has a temperature of 35 K, a envelope mass of 70 $M_\odot$ and a bolometric luminosity of 2140 $L_\odot$ (including the WISE data from Wright et al. 2010). Source B (SED not shown) has a temperatures of 17 K, mass of 300 $M_\odot$, bolometric luminosity of 470 $L_\odot$.

Source A is located between the evolutionary tracks of the 6.2 and 8.0 $M_\odot$ stars in the $L/M$ diagram (Molinari et al. 2008), suggesting it will form an intermediate to high mass star and that it has an age of $\sim$4.5$\times$10$^5$ yr. Source B lies on the 13.5 $M_\odot$ track and its combination of estimated mass and luminosity gives an age of 3$\times$10$^5$ yr.

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Figure D1. *Herschel* 70 μm, 250 μm and MAMBO 1.2 mm maps of IRAS 18586+0106 revealing two clear clumps, A and B. The position of the IRAS source is marked by a star. The radio source (Molinari et al. 1998) is marked by a cross. The beam FWHM is given in the bottom left corner.