Herschel-ATLAS: dust temperature and redshift distribution of SPIRE and PACS detected sources using submillimetre colours

Journal Article

How to cite:


For guidance on citations see FAQs.

© 2010 ESO

Version: Version of Record

Link(s) to article on publisher’s website:

http://dx.doi.org/doi:10.1051/0004-6361/201014586

Copyright and Moral Rights for the articles on this site are retained by the individual authors and/or other copyright owners. For more information on Open Research Online’s data policy on reuse of materials please consult the policies page.
Herschel-ATLAS: Dust temperature and redshift distribution of SPIRE and PACS detected sources using submillimetre colours


(Affiliations are available in the online edition)

Received 30 March 2010 / Accepted 26 April 2010

ABSTRACT

We present colour–colour diagrams of detected sources in the Herschel-ATLAS science demonstration field from 100 to 500 μm using both PACS and SPIRE. We fit isothermal modified black bodies to the spectral energy distribution (SED) to extract the dust temperature of sources with counterparts in Galaxy And Mass Assembly (GAMA) or SDSS surveys with either a spectroscopic or a photometric redshift. For a subsample of 330 sources detected in at least three FIR bands with a significance greater than 3σ, we find an average dust temperature of (28 ± 8) K. For sources with no known redshift, we populate the colour–colour diagram with a large number of SEDs generated with a broad range of dust temperatures and emissivity parameters, and compare to colours of observed sources to establish the redshift distribution of this sample. For another subsample of 1686 sources with fluxes above 35 mJy at 350 μm and detected at 250 and 500 μm with a significance greater than 3σ, we find an average redshift of 2.2 ± 0.6.

Key words. submillimeter: galaxies – Galaxy: evolution – galaxies: high-redshift

1. Introduction

The total intensity of the extragalactic background light at far-infrared (FIR) wavelengths has been measured with absolute photometry (Puget et al. 1996; Fixsen et al. 1998; Dwek et al. 1998). Deep surveys are now starting to resolve the cosmic FIR background into discrete sources with the fraction resolved varying with wavelength (see reviews in Blain et al. 2002; Hauser & Dwek 2001; Lagache et al. 2005; Dye et al. 2009). Despite some successes, there are still large uncertainties about the nature and evolution of the submillimetre (submm) galaxy population, including their redshift distribution.

The Herschel Astrophysical Terahertz Large Area survey (H-ATLAS) is an open-time key program of the Herschel Observatory (Pilbratt et al. 2010) that will survey roughly 550 deg2 of sky over 600 h of observations (Eales et al. 2010). To cover the widest possible area, H-ATLAS uses the maximum possible scan rate for the telescope at 60 arcsec s−1 in the parallel mode of PACS (Poglitsch et al. 2010) and SPIRE (Griffin et al. 2010). The observations cover five photometric bands with 100 and 160 μm data from PACS, and 250, 350, and 500 μm data from SPIRE. The H-ATLAS fields include 1 field close to the northern Galactic pole (150 deg2), 3 fields (each 36 deg2) coinciding with the GAMA redshift survey fields (Driver et al. 2009), and 2 fields (total area of 250 deg2) near the southern Galactic pole.

During the science demonstration phase (SDP), H-ATLAS observed 14.4 deg2 in the GAMA-9 h field near the ecliptic plane to a 5σ depth of 35–90 mJy. In this Letter we present the colour–colour diagrams derived from the five bands from 100 to 500 μm for the H-ATLAS sources that are cross-identified in at least 3 of the SPIRE and PACS bands. We then discuss the dust temperature and redshift distributions based on certain assumptions about the spectral energy distribution (SED) and subject to selection effects associated with the sample selection.

2. Colour–colour diagrams

We use source catalogues generated for the H-ATLAS consortium by Rigby et al. (in prep.), derived from SPIRE and PACS maps presented in Pascale et al. (in prep.) and Ibar et al. (in prep.), respectively. The source catalogues are supplemented with cross-identification information from the GAMA survey (Driver et al. 2009) and SDSS DR-7 (Abazajian et al. 2009) as...
We consider the H-ATLAS source sample with detections at 3.3 Dust temperature distribution
network code (Collister & Lahav 2004), trained with photometry
spectroscopic redshifts from GAMA or SDSS are not available,
described in Smith et al. (in prep.). For the sources for which
Dust temperature distribution
spectra from the GAMA spectroscopic survey (Driver et al. 2009),
DEEP2 (Davis et al. 2007), zCOSMOS (Lilly et al. 2007), and
the 2SLAQ-LRG (Cannon et al. 2006) survey.

In Fig. 1, we show the colour–colour plots of the H-ATLAS sources. We divide these plots in terms of colours based on either SPIRE only or SPIRE and PACS data. Our flux selection results in selecting 1686, 402, and 158 sources, from top to bottom of Fig. 1. For reference, the total ATLAS catalogue for this field contains ~6600 sources (Rigby et al., in prep.).

The colour–colour plots are filled with 10^6 black-body spectra at a single dust temperature, T_d, modified by a frequency-dependent emissivity function ε_ν ∝ ν^β, where the flux density f_ν is

\[ f_\nu = \epsilon_\nu B_\nu \propto \nu^{\beta} \left[ \frac{h \nu}{kT_d} \right] - 1. \]  

In generating these models, we consider uniform ranges of dust temperature from 10 to 60 K, emissivity parameter β from 0 to 2, and redshift from 0 to 5. The choice of 0 < β < 2 compared to 1 < β < 2 makes a minor difference since we also broaden the SED tracks in the colour diagram by adding an extra Gaussian standard deviation of 10% to the fluxes used to compute the model colour. This scatter accounts for the broadening of data in the colour–colour plane caused by flux uncertainties. As shown in Fig. 1, we find that the colour diagram of sources with fluxes only from SPIRE are well within the limits defined by the models we have considered. When we examine PACS colours, we find that some points lie outside the same set of tracks as used for the SPIRE-only colour diagram. While some of these outlier points may be caused by either the fractionally larger flux errors of PACS or contamination from a neighboring source, it is possible that some of these sources are not accurately described by our simple isothermal SED model, requiring for instance a second dust component (Dunne & Eales 2001) or a more complex SED model.

When fitting a simple modified black-body model to the data, we must keep in mind that there is a partial degeneracy between β and T_d and, more importantly, a perfect degeneracy between T_d and z. The peak of the SED is determined by the ν/T_d term in the exponential, so that a measurement of the colours alone constrains only the ratio (1 + z)/T_d. However, assuming reasonable priors on the free parameters of the SED model (in our case, β and T_d) it is still possible to estimate a qualitative redshift distribution for our sample of sources. Alternatively, if secure redshifts are known from optical cross-identifications, we can determine the dust temperatures.

3. Dust temperature distribution

We consider the H-ATLAS source sample with detections at 3σ in at least two bands and at 5σ in one band of either PACS and SPIRE that have been robustly (reliability parameter R_{LR} > 0.9, Smith et al., in prep.) identified with GAMA or SDSS DR-7 galaxies. We also require that there is a known spectroscopic redshift from either GAMA or SDSS, or from the photometric redshift catalog (with (1 + z)/σ_z > 5 and z/σ_z > 1) that was generated for this field and cross-identified with ATLAS sources (Smith et al., in prep.). We select 330 sources, which correspond to the low redshift subsample of the galaxies selected in the colour–colour diagrams. For each galaxy, we perform a single temperature fit from the above equation assuming β = 1.5. We found that an isothermal SED model generally is a good fit to the 330 galaxies. Fitting a two-component SED model does not, on average, seem to provide a closer fit. The relatively good fit provided by the isothermal SED model is probably due to most of our sample consisting of low redshift galaxies and therefore that we do not probe much of the Wien part of the galaxy SED.
In Fig. 2, we summarize our results and compare the H-ATLAS dust temperatures with dust temperatures in the literature for a variety of sub-mm bright galaxies. These samples are: (i) the sources in BLAST detected above 5σ in at least one of the BLAST bands with either a COMBO-17 (Wolf et al. 2004) or a SPIRE photometric redshift (Rowan-Robinson et al. 2008) and Spitzer-MIPS 70 and 160 μm fluxes (Dye et al. 2009; β = 1.5 fixed); (ii) local ULIRGS observed with SCUBA at 450 and 850 μm and complemented with IRAS 60 and 100 μm fluxes (Clements et al. 2010; β varied); (iii) SCUBA sub-mm galaxies detected at better than 3σ at 850 μm and having redshifts determined from Keck-I spectroscopy (Chapman et al. 2005; β = 1.5 fixed); and (iv) local IRAS-selected galaxies with 60 and 100 μm fluxes complemented with SCUBA 850 μm (Dunne et al. 2000; β varied).

In Table 1, we list average dust temperatures as a function of redshift for the 330 H-ATLAS sources reliably identified with GAMA and SDSS DR-7 galaxies having a known spectroscopic (black filled triangle) or a photometric (black empty triangle) redshift. The selected sources are: (i) the sources in BLAST detected above 5σ in at least one band of either PACS or SPIRE with 3σ flux measurements in at least 2 more bands. The isothermal spectral fits assume β = 1.5. For other data plotted, see Sect. 3.

In Fig. 3, we plot the dust temperature versus FIR luminosity for 330 H-ATLAS sources reliably identified with GAMA and SDSS DR-7 galaxies with a known spectroscopic (black filled triangle) or a photometric (black empty triangle) redshift. The FIR luminosity is obtained by fitting the PACS and SPIRE measurements with a modified black body (parameters are the dust temperature and the luminosity density, β is fixed to 1.5) and this model SED is integrated between 8 and 1100 μm. We also indicate BLAST and H-ATLAS best-fit $T_d$ vs. $L_{\text{FIR}}$ relation with a dashed blue line (from Dye et al. 2009) and a dashed red line, respectively.

### Table 1. Average dust temperatures as a function of redshift for the 330 H-ATLAS galaxies (Col. 2) and for all the data (Col. 4) presented in Figs. 2 and 3 (including H-ATLAS).

<table>
<thead>
<tr>
<th>z-range</th>
<th>H-ATLAS $N_{\text{obs}}$</th>
<th>$T_d$ (K)</th>
<th>All data $N_{\text{obs}}$</th>
<th>All data $T_d$ (K)</th>
</tr>
</thead>
<tbody>
<tr>
<td>All z</td>
<td>330</td>
<td>28 ± 8</td>
<td>657</td>
<td>30 ± 9</td>
</tr>
<tr>
<td>0 &lt; z &lt; 0.1</td>
<td>106</td>
<td>28 ± 5</td>
<td>635</td>
<td>32 ± 9</td>
</tr>
<tr>
<td>0.1 &lt; z &lt; 0.5</td>
<td>186</td>
<td>29 ± 8</td>
<td>620</td>
<td>28 ± 8</td>
</tr>
<tr>
<td>0.5 &lt; z &lt; 1</td>
<td>33</td>
<td>23 ± 5</td>
<td>67</td>
<td>24 ± 7</td>
</tr>
<tr>
<td>z &gt; 1</td>
<td>5</td>
<td>35 ± 4</td>
<td>95</td>
<td>37 ± 10</td>
</tr>
</tbody>
</table>

### 4. Redshift distribution

Since most of our sources lack redshifts, we consider another example application of our colour diagram and infer the statistical redshift distribution $N(z)$ for the source samples plotted in Fig. 1. We do this by first gridding the colour–colour plane along the SED tracks into redshift bins. We then convert the number of sources within a grid region of colours to a binned redshift distribution (Hughes et al. 2002). This method is equivalent to extracting the redshift probability distribution function for the whole sample if we had simply fitted SEDs to individual fluxes and taken the sum of the redshift probabilities of each source. While the redshift for an individual galaxy is largely uncertain,
identified sub-mm galaxies with SCUBA at 850 μm and spectroscopically at Keck (Chapman et al. 2005), the infrared ISGAI (Aretxaga et al. 2007), and a median redshift of greater than 1. Half Degree Extragalactic Survey determined photometrically is consistent with the average redshift of 2. Favour the high redshift-end of the distribution. Based on cross-correlation we extract should be a reasonable estimate of the true distribution of the source sample. Figure 4 shows the redshift distribution for SPIRE sources detected in all 3 bands (Fig. 1a). For the sample of 1686 sources with flux densities above 35 mJy at 350 μm and above 3σ at 250 and 500 μm, we find the average redshift to be 2.2 ± 0.6. This is consistent with the average redshift of 2.4 ± 0.4 for the radio-identified sub-mm galaxies with SCUBA at 850 μm followed up spectroscopically at Keck (Chapman et al. 2005), the interquartile redshift range of 1.8 to 3.1 for sources in SCUBA Half Degree Extragalactic Survey determined photometrically (Aretxaga et al. 2007), and a median redshift of greater than 1 for a subsample of 250 μm sources as faint as 35 mJy (Dunlop et al. 2009). Our distribution, however, has an average redshift that is higher than that of a larger sample of BLAST sources (Dye et al. 2009). This difference is not unexpected given the differences in the sample selection. We emphasize again that Figure 4 shows the redshift distribution for SPIRE sources with a flux cut of 35 mJy at 250 μm, and at least 3σ and 3.5σ, respectively. By imposing this colour-cut, we distinguish a redshift population. We find average redshifts of 2.2 ± 0.6 at 160 μm, and 3 ± 0.7 at 250 μm, and z = 0.8 ± 0.3 for sources detected in 100, 160, 250, and 350 μm with a flux cut of 35 mJy at 250 μm, and z = 0.7 ± 0.3 for sources detected in 100, 160, and 250 μm with a flux cut of >100 mJy at 160 μm. As is clear from Fig. 1, we find that lower redshift sources are more likely to be detected at shorter wavelengths, while brighter sources seen in longer SPIRE bands are likely to be on average at high redshifts. Some of the bright 500 μm sources that we detect, especially with a colour selection selecting red galaxies in the sub-mm, may also have their fluxes magnified by lensing and will be among the results forthcoming from H-ATLAS (Negrello et al., in prep.).

5. Conclusions

We have discussed the spectral energy distribution of sub-mm galaxies in a 14 deg², subset of the full H-ATLAS survey. We have used the colours in 5 bands to measure the dust temperature for the sub-sample of sources with known redshifts. We have also derived a qualitative estimate of the statistical redshift distribution of bright 350 μm selected galaxies assuming broad priors in temperature and the dust emissivity. We found that the average dust temperature of H-ATLAS sources is 28 ± 8 K, while 350 μm sources with fluxes above 35 mJy have an average redshift of 2.2 ± 0.6.

Acknowledgements. Ambland, Barton, Cooray, Leeuw, Serra and Temi acknowledge support from NASA funds for US participants in Herschel through JPL. Funding for the SDSS and SDSS-II has been provided by the Alfred P. Sloan Foundation, the Participating Institutions, the National Science Foundation, the US Department of Energy, the National Aeronautics and Space Administration, the Japanese Monbukagakusho, the Max Planck Society, and the Higher Education Funding Council for England. The SDSS Web Site is http://www.sdss.org/. The SDSS is managed by the Astrophysical Research Consortium for the Participating Institutions. The UKIDSS project is defined in Lawrence et al. (2007).

References


Fig. 4. Normalized redshift distribution (black thick solid line) of SPIRE sources in the colour–colour diagram of Fig. 1a, selected to satisfy the following criteria: $S_{350}/S_{250} > 0.75$ and at least 3σ detections at 250 and 500 μm. The dashed line and dotted-dash line represent Negrello et al. (2007) and Lagache et al. (2004) models for the same flux selection. Both model curves have similar but not identical distributions, although Lagache model predicts a larger number of high-redshift galaxies. In addition, we show $N(z)$ for the same subsample using a colour-cut of $S_{350}/S_{250} < 0.75$ (red thin dotted line). The distribution shown here is for 1686 sources, a subset of about 6600 H-ATLAS sources in the SDP field.