SHOCK SYNTHESIS OF ORGANICS FROM SIMPLE ICE MIXTURES? M. C. Price1, M. J. Burchell1, K. Miljkovic2,3, A. T. Kearsley1 and M. J. Cole1. 1School of Physical Sciences, University of Kent, Canterbury, CT2 7NH, UK (mcp2@star.kent.ac.uk). 2Dept. of Earth Sciences, University College London, London, WC1E 6BT, 3PSSRI, The Open University, Milton Keynes, MK7 6AA. 4IARC, Department of Mineralogy, Natural History Museum, London, SW7 5BD.

Introduction: Several recent announcements of: i) the discovery of glycine on the comet 81P/Wild-2/P [1, 2]; ii) the successful shock synthesis of prebiotic compounds from liquid targets [3] and, iii) the results of molecular dynamics simulations demonstrating that amino acids could be created via shock synthesis of ices [4], have prompted a sequence of impact experiments using a light gas gun (LGG) at the Univ. of Kent [5]. The purpose of these experiments is to attempt to synthesise organic compounds from a mixture of simple ices (CO$_2$, NH$_3$ and H$_2$O).

NH$_3$ compounds, CO$_2$ and H$_2$O ices in the Saturnian system: According to observations made by Cassini’s Visual and Infrared Mapping Spectrometer (VIMS) instrument, Enceladus’ surface is composed mostly of nearly pure water ice except near its south pole, where there are light organics, CO$_2$, and amorphous and crystalline water ice. The absorptions near 3.44 and 3.53 µm could be due to short-chain organics, but other features in the spectrum are still unidentified [5]. Remote IR acquired new high-resolution spectra of Iapetus, Tethys, Enceladus and Rhea that show the absorption feature of ammonia hydrate [6]. It thus seems probable that there are conditions on the surfaces of bodies in the Saturnian system where ammonium compounds, CO$_2$ and water ice co-exist in a solid form. Impact of a bolide traveling with sufficient energy to promote shock synthesis of more complex organic compounds, including amino acids, from these ices.

Laboratory experiments: Targets were prepared as follows: CO$_2$ ice (commercially purchased from BOC Ltd., stored in a freezer at a temperature of -130 °C) was repeatedly passed through a clean domestic ice crusher until the fragments’ largest dimension was no bigger than ~3 mm. The crushed CO$_2$ ice was then placed in a lidded (135 mm x 95 mm x 110 mm) polystyrene box. 150 ml of previously chilled (~40°C) aqueous ammonia solution (Sigma-Aldrich, Cat. # 32, 014-5) was added, and the mix was shaken until the ammonia started to freeze, creating a well-mixed solid target in the box. The box was then returned to the low temperature freezer and cooled to -130°C.

The targets remained in the freezer until the LGG was ready to fire, and then they were removed and placed in the target chamber which was then evacuated to ~30 mbar. The time taken from removal from the freezer to impact was approx. 20 minutes. In order to reduce target contamination from carbon-bearing gun debris during the shot, the (impact) target was wrapped in a layer of cling film (~10 µm thick). A blank control target was also placed in the target chamber (out of the direct line of impact) and left uncovered for comparative analysis.

Preparation of residues: After the shot (Table 1), the impacted target and the control target were removed and placed into a vented oven at a temperature of approximately 90°C, to speed the sublimation/vapourisation of the remnant CO$_2$ ice and liquid NH$_3$ and also to decompose the ammonium bicarbonate which is inevitably produced via the reaction: NH$_3$+H$_2$O+CO$_2$→NH$_4$HCO$_3$. The water is captured by atmospheric condensation and from the ammonia solution itself. Both targets were left for 12 hours, then the fine powdery residue (coating the bottom and sides of the box) was dissolved in 20 ml of HPLC water (HPLC grade grade, Fisher Scientific) and filtered through grade 1 filter paper into a washed, sealable, glass container and the water left to evaporate in clean air.

Residue analysis: Residues were analysed with a Jobin Yvon µ-Raman HR640 spectrometer ($h_{\text{laset}}$ = 632.8 nm) and a FTIR spectrometer (Nicolet 380 FT-IR)

Table 1. LGG shot parameters.

<table>
<thead>
<tr>
<th>Shot ID</th>
<th>Projectile</th>
<th>Dens. (g cm$^{-3}$)</th>
<th>Dia. (mm)</th>
<th>Vel. (km s$^{-1}$)</th>
</tr>
</thead>
<tbody>
<tr>
<td>G161009#1</td>
<td>SS 304</td>
<td>8.00</td>
<td>1.5</td>
<td>5.75</td>
</tr>
</tbody>
</table>

Figure 1: Example of residue grain (on gold substrate) from shot G161009#1
Fig. 1 shows an example of a grain found after washing and filtering the residue. As can be seen, the crystal is irregular and appears to contain smaller (unknown) grains within. It should be noted that several such crystals were found in the residue from G161009#1 but were NOT detected in the control residue.

![Raman spectrum of crystal from Fig. 1](image1)

Figure 2. Raman spectrum of crystal from Fig. 1 (upper line-black), pure glycine (middle line-pink), pure ammonium bicarbonate (bottom line-blue).

![FTIR spectra of residue](image2)

Figure 3. FTIR spectra of residue from G161009#1 (central curve-black). Also shown are the FTIR spectra of pure glycine (upper curve-pink) and ammonium bicarbonate (lower curve-blue).

Table 2: Raman line wavenumbers and relative intensities for pure glycine and the unknown shot residue.

<table>
<thead>
<tr>
<th>Compound</th>
<th>Raman lines and intensity(^v)</th>
</tr>
</thead>
<tbody>
<tr>
<td>Pure glycine(^\d) (on glass)</td>
<td>504(s), 562(w), 610(w), 689(m), 900(vs), 934(w), 1052(m), 1132(m), 1161(w), 1327(s), 1340(vs), 1400(m,b), 1441(s), 1505(vw), 1584(w), 1677(w), 2966(s), 3003(s)</td>
</tr>
<tr>
<td>Residue (on gold)</td>
<td>505(m), 602 (w, b), 689 (vw), 871(w), 901(vs), 981 (vw), 1128 (vw), 1356(w), 1395(vw), 1407(w), 1475(vs), 1512(vw), 1600 (w/vw), 1627(w/m)</td>
</tr>
</tbody>
</table>


From Fig. 2 (and Table 2) it can be seen that there are some matching Raman lines between the residue and pure glycine, but there is also a broad underlying fluorescence in the residue spectrum which may be obscuring some of the less intense lines. There is, however, a better match with the glycine spectra than with the ammonium bicarbonate spectra – which is the most likely contamination.

FTIR analysis was performed by scraping the residue into a pile using a clean spatula and placing it on the diamond anvil stage of the spectrometer. This analysis technique is potentially more ‘destructive’ of our sample than in-situ Raman spectroscopy, simply due to the loss of residue during transfer. Additionally, the method is ‘non-selective’ and will give an FTIR spectrum which is a summation of all the compounds within the residue. Unfortunately, the very small amount of control target residue made it impossible to obtain a FTIR signal above the background.

The FTIR spectra are difficult to accurately interpret, but do seem to suggest a mixture of ammonium bicarbonate and other organic matter (not excluding glycine). The small quantity of residue produced (<1 mg) makes a more accurate determination difficult using our current equipment.

**Conclusions:** Although more sensitive analytical testing of this type of shock residue is required (and this is ongoing), there are tantalising indications that we have synthesized organic compounds (including, possibly glycine) by shock impacting a simple ice mixture. This could provide a creation mechanism for prebiotic compounds in the outer solar system. Further experiments are currently underway and time has been obtained on a NMR spectrometer for further analysis.

**References:**

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